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Photoelectric and CCD photometry of E and S0 galaxies

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ABSTRACT

We present *BR* photoelectric photometry for 352 E and S0 galaxies that are part of a large survey of the properties and peculiar motions of galaxies in distant clusters. Repeat measurements show our internal errors to be 2–3 per cent in *B* and *R* and 1–2 per cent in *B–R*. Comparisons of *BR* and *BVR* reductions for 10 galaxies also observed in *V* show small systematic errors due to differences between the spectral energy distributions of stars and galaxies. External comparisons with *B–V* colours in the literature confirm that these colours are good to 1 per cent. We also describe *R*-band CCD observations for 95 of the galaxies and place these on a *BR* photometric system for photoelectric and CCD photometry, with a common zero-point good to better than 1 per cent. We find the rms precision of both our photoelectric and CCD *R* magnitudes to be 2–3 per cent for galaxies as faint as $R \approx 15$. Errors in galaxy magnitudes of this order introduce errors of $\lesssim 2$ per cent into D_n - σ distance estimates, corresponding to errors in peculiar velocities for single galaxies of $\lesssim 200 \text{ km s}^{-1}$ at a distance of 10 000 km s^{-1} .

Key words: galaxies: distances and redshifts – galaxies: elliptical and lenticular, cD – galaxies: photometry.

1 INTRODUCTION

Over the past decade, CCD photometry has replaced photoelectric photometry as the method of choice for absolute photometry of galaxies. This preference is due to the imaging capability, high quantum efficiency and ready availability of CCDs. In 1986 we began a combined photometric and spectroscopic study of more than 500 E and S0 galaxies in distant clusters, with the dual scientific goals of understanding the physical properties of these galaxies and using these same properties to measure the distances and peculiar velocities of the clusters. At the outset of this project, we realized that most of the photometry could best and most easily be carried out using CCDs. However, we decided to obtain both CCD and photoelectric photometry for as many of the galaxies in our survey as possible. This decision was based partly on the critical need for accurate photometric zero-points in order for our programme to attain its scientific goals. As pointed

out by Dressler et al. (1987), errors in the photometric zero-points enter linearly into distance estimates for galaxies. Using the D_n - σ relation as defined by Dressler et al. to determine distances, a zero-point error of 0.01 mag (or 0.92 per cent) corresponds to an error in D_n of approximately 0.6 per cent. Thus zero-point errors of less than 0.03 mag are required to achieve distances accurate to better than 2 per cent (i.e. $< 200 \text{ km s}^{-1}$ at a distance of 10 000 km s^{-1}). The second reason for our decision was that it was highly desirable that the present survey be reconcilable with the photoelectric data assembled in previous studies, especially the extensive ‘7 Samurai’ (Faber et al. 1989).

Relatively little information exists in the literature on the degree of agreement between the zero-points and scales of CCD-calibrated and photoelectric-calibrated photometry of galaxies. Jørgensen, Franx & Kjørgaard (1992) have compared CCD and photoelectric measurements taken in the *B* band and have obtained results consistent with known observational errors of 1–2 per cent. Very little information exists concerning comparisons of galaxy observations in *R*, since relatively few photoelectric observations of galaxies have been made in this passband.

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Our goal was to obtain CCD-calibrated R -passband photometry of our entire galaxy sample, and a separate photoelectric calibration in R for at least half of these galaxies. The photoelectric data would also yield $B-R$ colours. Only in this manner could we be sure that (i) our data could be reliably placed on the existing photoelectric-based photometric system for galaxies, (ii) that the true extent of zero-point differences between CCD and photoelectric systems would be known, and (iii) that all of our data would be on an internally self-consistent system.

These goals for the photoelectric data have been achieved, and in this paper we present the results of all the photoelectric photometry that was carried out for this survey. We also now have sufficient overlap between the CCD and photoelectric photometry to permit a detailed comparison of the R magnitudes obtained from these two methods. We emphasize the accuracy of the magnitudes, especially in the R passband, as it is the error in magnitude that directly affects the accuracy of distance measurements. Section 2 presents the photoelectric photometry for a total of 344 programme galaxies and a further eight standard galaxies. The accuracy of these magnitudes is evaluated primarily by internal comparisons, although comparisons are also made, where possible, with $B-V$ colours from the literature. Our methodology for obtaining aperture magnitudes from CCD images is described in Section 3. The CCD and photoelectric magnitudes are compared in Section 4, and our results are summarized in Section 5.

2 PHOTOELECTRIC PHOTOMETRY

2.1 Methods

The scientific goal of our photoelectric observations was to obtain galaxy B and R magnitudes that were accurate to 0.03 mag or better. The choice of these passbands was dictated in part by the use of the R passband in the CCD observations, and in part by the desire to be able eventually to tie these data directly to the B -passband measurements of Burstein et al. (1987).

All photoelectric observations for this programme were made with the 1.3-m Kitt Peak National Observatory (KPNO) telescope, equipped with the Mark III photometer. The Landolt (1983) list of $UBVRI$ standard stars was used as the source for standardization of the photometry. A log of the observations is given in Table 1, listing the night of observation (local time), the photomultiplier (identified by both photocathode type and KPNO Cold Box designation) and filter system, the photometric quality, Q , of the night (defined below), the number of observations of Landolt standard stars that were made, the number of individual Landolt standard stars observed, the number of standard galaxies (defined below) observed, and the number of programme galaxies observed.

The programme galaxies were observed in the B and R filters. Two important sources of error in the determination of galaxy magnitudes are photon noise and zero-point errors in the transformation of magnitudes to a standard system. Adequate integration time can solve the first problem. The second problem requires that a set of well-observed bright galaxies, not in our programme, be established as 'standard' galaxies to help zero our magnitudes to an internally consistent BR system.

The standard galaxies and the standard stars were observed in BVR , as the addition of the V filter took little extra time. Separate integrations were taken every 10 s for the galaxies, and every 5 s for the standard stars. Sets of 10-s integrations were arranged so that no more than 4 min separated any galaxy observation from a sky observation in the same filter. No attempt was made to compensate for faint stars within the photoelectric apertures, although the smaller aperture sizes were used to avoid stars where possible. The photoelectric magnitudes therefore include everything on the sky within the aperture when it is centred on the galaxy.

The chief problem in obtaining accurate galaxy magnitudes is the same as for stellar magnitudes, namely monitoring of nightly extinction. To overcome this, standard stars were chosen with a range of $B-R$ colours that bracketed the galaxy colours. Nightly observations tried to balance a wide sample of individual stars with multiple observations of several stars over a range in airmass. Many different stars were observed within a given observing run, to average over the known errors in the standard star magnitudes (sometimes ≥ 0.01 mag).

On any given night, all galaxy observations were sandwiched between standard star observations, which were obtained approximately on an hourly basis. Observations were greatly enhanced by the fast slewing ability of the 1.3-m telescope and the on-line data reduction capability of the Mark III photometer. Standard star observations took an average of 2.7 min between any two points in the sky. We typically observed six standard stars at all airmasses, east and west, every hour, together with more stars at the beginning and end of the night. As is evident in Table 1, we observed between five and seven stars an hour, depending on the length of the night.

The on-line data reduction provided by the Mark III photometer, combined with our observing strategy, per-

Table 1. Log of photoelectric observations.

Date	Instrumental setup (MkIII photometer +)	Q	# std obsns	# std stars	# std gals	# gals
88.05.07	#27 S-20 + 'K' UBVR	B	58	28	2	18
88.05.08	#27 S-20 + 'K' UBVR	B	57	23	2	18
88.05.09	#51 GaAs + 'J' UBVR	A	50	18	1	26
88.05.10	#51 GaAs + 'J' UBVR	A	50	22	2	22
90.06.19	#54 S-20 + 'K' UBVR	B	31	21	1	7
90.06.20	#54 S-20 + 'K' UBVR	B	36	21	1	16
90.06.21	#54 S-20 + 'K' UBVR	B	29	21	1	12
90.08.20	#54 S-20 + 'K' UBVR	A	62	21	2	15
90.08.21	#54 S-20 + 'K' UBVR	A	59	25	0	15
90.08.22	#54 S-20 + 'K' UBVR	A	55	22	2	19
90.08.23	#54 S-20 + 'K' UBVR	A	57	29	1	19
91.05.15	#54 S-20 + 'K' UBVR	A	22	15	0	7
91.05.16	#54 S-20 + 'K' UBVR	A	39	20	2	15
91.05.17	#54 S-20 + 'K' UBVR	A	37	21	2	16
91.05.18	#54 S-20 + 'K' UBVR	A	36	18	0	9
91.07.09	#54 S-20 + 'K' UBVR	A	49	22	1	12
91.07.10	#54 S-20 + 'K' UBVR	A	41	23	2	16
91.07.11	#54 S-20 + 'K' UBVR	A	20	12	1	8
91.07.12	#54 S-20 + 'K' UBVR	B	20	12	1	6
91.10.03	#54 S-20 + 'K' UBVR	A	23	17	0	6
91.10.04	#54 S-20 + 'K' UBVR	A	36	21	1	15
91.10.05	#54 S-20 + 'K' UBVR	A	33	21	2	22
91.10.10	#54 S-20 + 'K' UBVR	A	56	28	2	15
91.10.12	#54 S-20 + 'K' UBVR	A	30	13	0	9
91.10.13	#54 S-20 + 'K' UBVR	A	45	20	0	22
91.10.14	#54 S-20 + 'K' UBVR	A	51	25	2	24
Total:			1082	539	31	389

mitted us to monitor extinction variations during each night on an hourly basis. On most nights we found that the Kitt Peak sky experienced temporal variations of 1–3 per cent in the zero-point of extinction that were relatively grey and adiabatic, in that they were consistent over the whole sky at all airmasses and in all filters. On several nights, these adiabatic changes in transparency were confirmed by comparing the 1.3-m measures to those being taken by other observers on the 0.9-m KPNO telescope. For each run, the uniformity of the photomultiplier was established by scanning a bright star across the field of view.

On 1988 May 8 we sadly witnessed the demise of the very sensitive Cold Box # 27 S-20 photomultiplier, which lost its stability during this night. This was unfortunate, as the substitute S-20 cold box we used for most of the remainder of the survey (# 54) was 1 mag less sensitive than Box # 27 in *B*, and 1.7 mag less sensitive in *R*. The loss of Cold Box # 27 cost us the ability to obtain two aperture observations per galaxy, as originally planned, so that most galaxies were observed through one aperture only.

Of the first 59 nights allotted to this programme, only 11 nights (19 per cent) were usable; of the last 20.5 nights, 14.5 (71 per cent) were usable. Thus, while our overall success rate is somewhat better than average for KPNO (32 per cent), it came at the expense of many discouraging nights at the beginning of the programme.

The extinction and transformation coefficients for each night were derived in a graphical, iterative manner. A graphical manner was chosen for the reduction, as it is often found that ‘wild’ observations bias the result if a blind multiparameter least-squares fit is made. These wild observations result from a number of factors, including pointing at the wrong star, loss of aperture centring, poor focus, variable standard stars (several were found), or poor standard star magnitudes. All nights within each run were reduced both separately and as part of the run, to improve the precision of the transformation to the standard system. Differential changes in extinction from night to night were absorbed into the temporal terms for a given night. The transformation coefficient for *B* magnitudes was the same for all runs with Cold Box # 54 (the only photomultiplier used on more than one run), while that for the *B* – *R* term differed by 0.02 mag between the 1990 and 1991 runs.

For 19.5 nights the temporal changes in extinction during a night were less than 0.030 mag, and the resulting *total* dispersions of *B*, *V* and *R* magnitudes for the standard stars were less than 0.015 mag. These nights are given a quality code of ‘A’ (see Table 1). On six nights the stability of the sky (or, on 1988 May 8, the photometer), once temporal changes were removed, was 0.015–0.030 mag in all filters. These nights are given a quality code of ‘B’. Observations were also made on seven other nights, but during reduction it was found that the standard star observations had a scatter of > 0.03 mag, even after removing temporal changes. These nights we classify as quality ‘C’. As discussed in the following section, comparison of galaxy observations taken on C nights with those taken on A or B nights show the C-night observations to be of inferior quality, so they are only used for comparison purposes in this paper.

Due to the fact that the programme galaxies were only observed in *BR*, while the standard stars and standard galaxies were observed in *BVR*, the standard stars were reduced in both *BR* and *BVR*. Galaxy magnitudes were then

predicted from the *BR* transformations obtained in each reduction. The reliability of this transformation is evaluated below by both internal and external comparisons.

All the photoelectric photometry taken on A- or B-quality nights is presented in Tables 2 and 3. In each table are given the ID number of the galaxy used in these studies, the galaxy’s name and (in Table 2) position, the date, time and airmass of the observation, the diameter of the photometric aperture used (in arcsec), the quality of the observation (A or B), the measured *B* and *R* magnitudes and *B* – *R* colour of the galaxy (Table 2), or *BVR*, *B* – *V* and *B* – *R* (Table 3).

The positions for all programme galaxies were determined by one of us (DB) using the Guide Star Astrometric Support System (GASP) of the Space Telescope Science Institute (STScI). These positions were originally calculated in J2000 coordinates and then precessed to B1950 coordinates. The positions are quoted to ~ 1-arcsec precision, which is about the expected uncertainty due to centring errors (and which dominates the errors of ~ 0.5 arcsec from the GASP astrometry). Positions for non-programme galaxies were obtained from either the RC2 (de Vaucouleurs, de Vaucouleurs & Corwin 1978) or, for the Coma cluster galaxies, Faber et al. (1989).

For most galaxies, the *B* and *R* magnitudes are formally determined to better than 1 per cent in terms of photon-statistical errors. These magnitudes are corrected to zero airmass, but have no corrections applied for either Galactic extinction or the K-correction. Table 2 lists a total of 397 *BR* photoelectric observations of 344 programme galaxies (including two galaxies in the Coma cluster) and 38 observations of eight standard galaxies. There are two or more *BR* photoelectric observations for 43 of the programme galaxies and for six of the standard galaxies. Table 3 gives *BVR* data from 31 observations of eight standard galaxies and two observations of two programme galaxies.

2.2 Internal comparisons

Two types of internal comparison can be made within the photoelectric data set. The first is a straightforward comparison of observations of the same galaxy made with the same aperture on different nights. The second is a comparison of the photometry when it is reduced in two different ways. One reduction procedure is used for all galaxies, both programme and standard, and uses the results of the *BR* reduction of the standard stars. The other reduction procedure uses the results of the *BVR* reduction of the standard stars to predict *BR* magnitudes for those few galaxies (mostly standard galaxies) with *BVR* observations. This second type of comparison cross-checks the degree to which the standard stars can consistently transform the galaxy colours. This is always of concern, given that the spectral energy distributions of galaxies are not the same as those of stars with the same colours.

2.2.1 Repeat photoelectric observations

Nights of quality A and B were stable to 1–3 per cent, while nights of quality C were much less stable. Of the 435 independent observations of programme and standard galaxies which were of quality A or B, 98 were duplicated with the same galaxy and the same aperture, but on *different* nights of quality A or B. There were also 96 duplicate observations that combined nights of quality C with other nights. Included

Table 2 - continued

ID#	Name	R.A. (1950)	Dec.	Date	UT	Air	D (")	Q	B	R	B-R
1	A76 A	00 36 51.2	+06 27 35	91.10.05	05:46	1.175	29.9	A	15.225	13.379	1.846
2	A76 B	00 37 53.0	+06 26 43	91.10.05	05:57	1.159	29.9	A	15.901	14.206	1.696
4	A76 D	00 37 55.5	+06 38 35	91.10.14	06:11	1.108	29.9	A	15.778	14.023	1.755
5	A76 E	00 36 19.7	+06 47 17	91.10.14	06:23	1.103	19.2	A	15.972	14.117	1.855
6	A76 F	00 36 29.4	+07 13 05	91.10.14	06:39	1.101	19.2	A	16.035	14.199	1.835
10	J29 C	00 40 22.9	-09 30 16	91.10.04	06:56	1.331	29.9	A	16.980	15.153	1.827
13	A119 A	00 53 52.3	-01 31 59	90.08.22	09:33	1.214	29.9	A	15.959	13.631	1.728
14	A119 B	00 53 42.7	-01 31 33	90.08.22	09:45	1.205	29.9	A	15.380	13.595	1.785
15	A119 C	00 55 01.6	-01 39 40	90.08.22	09:59	1.199	29.9	A	15.526	13.789	1.737
17	A119 E	00 53 29.3	-01 36 17	90.08.22	10:16	1.197	29.9	A	15.934	14.253	1.681
21	A119 I	00 54 54.3	-00 44 31	91.10.05	06:09	1.252	29.9	A	16.056	14.187	1.869
24	A119 2	00 52 48.1	-01 37 24	91.10.05	06:27	1.228	29.9	A	16.619	15.415	1.205
25	J3 A	00 59 22.8	+26 40 59	91.07.10	09:56	1.335	29.9	A	15.614	13.789	1.825
26	J3 B	00 56 05.3	+26 42 29	91.07.10	10:14	1.254	29.9	A	16.209	14.431	1.779
27	J3 C	00 56 42.2	+26 47 23	91.07.10	10:30	1.202	29.9	A	15.662	14.011	1.651
31	J4 A	00 56 13.6	+12 42 11	91.10.03	07:59	1.063	29.9	A	15.168	13.380	1.788
32	J4 B	00 56 58.5	+12 43 00	91.10.03	08:10	1.069	29.9	A	16.010	14.303	1.708
38	A150 C	01 05 37.6	+01 55 38	91.10.04	07:28	1.153	29.9	A	15.831	14.064	1.767
39	A150 D	01 05 40.2	+01 54 36	91.10.04	08:09	1.164	29.9	A	16.387	14.756	1.631
40	A150 E	01 06 08.9	+01 55 32	91.10.04	08:29	1.186	29.9	A	16.730	15.023	1.708
43	A160 A	01 09 47.1	+15 29 47	90.08.20	10:11	1.047	29.9	A	15.685	13.993	1.692
44	A160 B	01 09 44.1	+15 27 39	90.08.20	09:26	1.067	19.2	A	17.180	15.330	1.850
45	A160 C	01 07 48.5	+15 55 29	90.08.22	10:54	1.047	29.9	A	16.148	14.281	1.867
46	A160 D	01 09 39.0	+16 03 39	90.08.23	09:07	1.091	29.9	A	16.563	14.694	1.869
47	A160 E	01 10 36.3	+15 15 05	90.08.22	11:13	1.059	29.9	A	15.732	13.989	1.743
47	A160 E	01 10 36.3	+15 15 05	90.08.23	10:44	1.046	29.9	A	15.712	13.983	1.729
50	A160 H	01 11 08.2	+15 14 36	90.08.23	10:13	1.044	19.2	A	16.812	15.119	1.693
51	A160 I	01 08 45.6	+15 18 26	90.08.23	10:28	1.043	19.2	A	16.722	14.809	1.913
52	A160 J	01 11 34.8	+15 38 02	90.08.20	10:24	1.042	29.9	A	15.653	14.068	1.585
56	A168 A	01 10 26.5	-00 01 05	91.10.12	08:58	1.328	29.9	A	14.319	12.708	1.610
58	A168 C	01 11 47.8	-00 05 05	91.10.13	05:36	1.271	19.2	A	16.981	14.644	1.737
59	A168 D	01 12 23.8	+00 09 58	91.10.12	08:42	1.276	29.9	A	15.594	13.801	1.793
60	A168 E	01 12 20.6	+00 02 19	91.10.12	09:09	1.356	29.9	A	16.453	14.623	1.830
62	A168 G	01 13 39.2	-00 22 20	91.10.13	05:52	1.247	19.2	A	16.605	14.849	1.787
70	J30 B	01 18 27.4	-14 05 49	91.10.13	06:50	1.442	19.2	A	16.915	15.167	1.748
73	J30 E	01 17 52.5	-14 09 06	91.10.10	07:30	1.437	19.2	A	16.680	15.010	1.670
73	J30 E	01 17 52.5	-14 09 06	91.10.13	07:06	1.436	19.2	A	16.707	15.042	1.665
79	A193 B	01 22 34.7	+08 23 44	91.10.10	05:52	1.177	29.9	A	16.776	14.966	1.810
81	A193 D	01 22 01.4	+08 14 58	91.10.10	06:13	1.145	29.9	A	15.988	14.238	1.750
85	J32 B	01 34 54.1	-09 25 23	91.10.03	08:45	1.343	29.9	A	15.686	14.023	1.662
96	A260 A	01 47 49.4	+32 50 04	90.08.23	10:58	1.001	29.9	A	15.116	13.340	1.776
98	A260 C	01 48 30.0	+32 47 02	90.08.23	11:09	1.000	29.9	A	15.472	13.712	1.760
100	A260 E	01 46 19.6	+32 50 51	91.10.10	06:31	1.046	29.9	A	15.755	14.020	1.734
101	A260 F	01 47 21.7	+33 14 52	91.10.14	07:37	1.000	29.9	A	15.670	13.935	1.736
102	A260 G	01 48 51.4	+33 17 26	91.10.14	07:48	1.001	19.2	A	16.035	14.273	1.762
105	A262 A	01 49 50.0	+35 54 21	91.10.05	07:45	1.010	29.9	A	14.652	12.981	1.671
106	A262 B	01 47 55.1	+36 01 42	91.10.05	08:02	1.005	29.9	A	13.367	11.858	1.509
107	A262 C	01 47 37.1	+36 07 24	91.10.05	08:13	1.003	29.9	A	14.240	12.537	1.702
108	A262 D	01 46 48.2	+35 32 15	91.10.05	08:19	1.002	29.9	A	14.209	12.472	1.737
110	A262 F	01 50 11.5	+36 34 25	91.10.14	08:03	1.006	19.2	A	14.995	13.178	1.817
113	A262 I	01 49 43.3	+35 55 30	91.10.05	08:25	1.003	29.9	A	15.766	13.318	1.766
115	J7 A	02 22 35.1	+36 44 21	91.10.03	09:02	1.004	29.9	A	15.350	13.586	1.764
117	J7 C	02 20 53.0	+36 50 11	91.10.04	09:29	1.011	29.9	A	15.683	14.014	1.669
118	J7 D	02 23 24.2	+36 48 56	91.10.04	09:40	1.018	29.9	A	15.399	14.667	1.732
119	J7 E	02 23 11.0	+37 04 02	91.10.04	10:02	1.038	29.9	A	15.967	14.245	1.722
122	J8 A	02 27 25.5	+22 55 54	91.10.10	08:31	1.012	29.9	A	15.544	13.741	1.803
127	J8 F	02 26 44.3	+22 43 14	91.10.10	08:42	1.014	29.9	A	16.087	14.184	1.903

Table 2. BR photoelectric data.

Table 2 - continued

ID#	Name	R.A. (1950)	Dec.	Date	UT	Air	D (")	Q	B	R	B-R
306	J11 E	13 03 17.4	+53 51 32	91.05.17	04:41	1.077	29.9	A	15.547	14.177	1.371
307	J11 F	13 05 14.1	+53 51 11	88.05.09	04:41	1.084	39.5	A	15.470	13.899	1.570
310	J12 A	13 41 08.0	+30 19 13	88.05.07	04:46	1.128	29.9	B	15.388	13.755	1.633
311	J12 B	13 40 38.4	+30 07 11	91.05.16	04:15	1.039	29.9	A	15.376	13.718	1.658
312	J12 C	13 40 31.4	+29 57 24	91.05.16	04:28	1.024	29.9	A	16.053	14.429	1.624
313	J12 D	13 41 57.2	+30 03 16	91.05.16	04:44	1.014	29.9	A	16.329	14.590	1.740
314	J12 E	13 40 06.1	+30 05 57	91.05.16	05:02	1.004	29.9	A	15.746	14.141	1.605
315	J12 F	13 40 06.3	+30 04 34	88.05.07	04:33	1.056	29.9	B	15.660	13.947	1.713
318	J13 A	13 52 54.7	+25 17 45	88.05.07	05:55	1.010	39.5	B	14.783	13.148	1.635
318	J13 B	13 52 54.7	+25 17 45	88.05.07	05:50	1.012	60.2	B	14.568	12.977	1.591
321	J13 D	13 52 53.8	+25 15 56	88.05.07	06:02	1.008	29.9	A	15.253	13.866	1.472
321	J13 E	13 52 53.8	+25 15 56	88.05.07	06:02	1.008	29.9	A	15.638	13.866	1.772
322	J13 F	13 51 21.8	+25 19 28	88.05.07	06:09	1.007	39.5	B	15.591	13.879	1.712
323	J13 G	13 50 25.9	+25 17 08	88.05.07	06:17	1.009	39.5	B	15.349	13.831	1.519
324	J13 H	13 50 18.2	+24 59 40	88.05.07	06:25	1.009	39.5	B	14.843	13.255	1.588
325	J13 I	13 48 46.2	+25 20 24	88.05.07	06:33	1.011	39.5	B	14.837	13.249	1.588
327	J13 J	13 48 12.5	+25 13 13	88.05.07	06:45	1.016	39.5	B	15.389	13.904	1.485
328	J13 K	13 48 11.5	+25 12 35	91.07.10	04:37	1.232	29.9	A	15.710	14.007	1.703
329	J13 L	13 48 05.7	+25 09 21	91.07.10	04:53	1.289	29.9	A	16.652	14.930	1.722
330	J13 M	13 47 33.3	+25 26 17	88.05.07	06:53	1.021	39.5	B	15.474	14.069	1.406
331	J13 N	13 53 29.5	+25 02 57	91.07.10	05:09	1.334	29.9	A	16.572	14.924	1.648
333	J13 P	13 53 50.8	+25 26 21	91.05.17	06:03	1.013	29.9	A	16.090	14.465	1.625
337	J14 A	14 05 27.7	-08 50 06	88.05.07	07:17	1.360	29.9	B	15.031	13.376	1.655
336	J14 B	14 05 27.7	-08 50 06	88.05.07	07:26	1.384	39.5	B	15.494	13.920	1.574
339	J14 C	14 03 34.3	+13 52 35	88.05.08	04:36	1.264	39.5	B	15.198	13.593	1.605
343	J14 D	14 44 39.3	+13 53 05	88.05.08	04:48	1.130	29.9	A	15.172	13.567	1.605
344	J14 E	14 44 16.3	+13 53 05	88.05.08	04:48	1.232	29.9	B	15.670	14.128	1.652
345	J14 F	14 44 24.9	+13 54 12	91.05.17	06:20	1.053	29.9	A	16.746	15.265	1.481
352	J14-1 A	14 44 41.7	+11 48 08	88.05.08	05:34	1.110	39.5	A	15.046	14.028	1.618
353	J14-1 B	14 44 29.2	+11 50 05	91.05.17	06:44	1.069	29.9	A	15.466	13.864	1.602
354	J14-1 C	14 44 21.9	+11 46 43	88.05.08	05:35	1.139	29.9	B	15.875	13.898	1.677
355	J14-1 D	14 44 03.0	+11 42 58	88.05.08	05:42	1.128	29.9	B	15.253	13.683	1.571
355	J14-1 E	14 44 03.0	+11 42 58	88.05.08	06:10	1.088	39.5	A	15.115	13.640	1.475
356	J14-1 F	14 46 42.6	+11 11 11	91.07.11	04:44	1.197	19.2	A	15.987	14.396	1.611
357	J14-1 G	14 44 41.9	+11 12 29	88.05.08	05:49	1.122	29.9	B	15.887	14.276	1.611
362	A1983-1 A	14 52 03.3	+16 33 27	88.05.08	05:57	1.084	39.5	B	15.058	13.326	1.732
362	A1983-1 B	14 52 03.3	+16 33 27	88.05.08	06:38	1.044	39.5	A	15.003	13.314	1.689
363	A1983-1 C	14 50 35.3	+16 54 23	88.05.08	06:06	1.067	29.9	B	16.202	14.594	1.609
363	A1983-1 D	14 50 35.3	+16 54 23	88.05.08	06:50	1.038	39.5	A	15.964	14.308	1.657
364	A1983-1 E	14 50 36.9	+16 55 53	91.05.16	06:03	1.045	29.9	A	16.233	14.525	1.709
365	A1983-1 F	14 50 23.5	+17 06 27	88.05.08	06:19	1.056	29.9	B	15.766	14.021	1.746
365	A1983-1 G	14 50 23.5	+17 06 27	88.05.08	07:02	1.035	39.5	A	15.877	13.931	1.646
366	A1983-1 H	14 50 23.5	+17 06 32	91.05.16	06:20	1.036	29.9	A	15.911	14.217	1.693
367	A1983-1 I	14 47 39.0	+17 00 58	91.05.16	06:42	1.037	29.9	A	16.748	15.137	1.611
372	A1983 A	14 52 13.3	+18 50 41	88.05.16	07:30	1.052	29.9	A	15.968	14.176	1.792
375	A1983 B	14 48 56.0	+18 57 45	91.07.11	05:00	1.178	29.9	A	16.336	14.616	1.720
376	A1983 C	14 48 59.3	+18 53 31	88.05.10	05:21	1.101	39.5	A	15.464	13.801	1.663
377	A1983 D	14 51 21.2	+18 16 23	91.07.11	05:21	1.230	29.9	A	16.460	14.671	1.789
383	J16 A	15 16 31.9	+04 42 05	88.05.08	06:59	1.139	29.9	B	15.239	13.538	1.701
383	J16 B	15 16 31.9	+04 42 05	88.05.09	07:16	1.128	39.5	A	14.996	13.324	1.672
383	J16 C	15 16 31.9	+04 42 05	88.05.08	07:04	1.136	60.2	B	14.616	13.032	1.584
384	J16 D	15 16 33.7	+04 30 51	88.05.08	07:12	1.134	29.9	B	15.445	13.732	1.714
384	J16 E	15 16 33.7	+04 30 51	88.05.09	07:26	1.128	29.9	A	15.440	13.692	1.748
385	J16 F	15 15 58.3	+04 51 24	90.06.19	05:14	1.136	29.9	B	16.205	14.630	1.575

Table 2 - continued

Table 2 - continued

Table 2 - continued

ID#	Name	R.A. (1950)	Dec.	Date	UT	Air	D (")	Q	B	R	B-R
681	A2589 A	23 21 27.0	+16 30 08	90.08.20	07:52	1.087	29.9	A	15.418	13.637	1.781
682	A2589 B	23 21 28.7	+16 32 11	90.08.22	08:57	1.296	29.9	A	16.281	14.607	1.674
683	A2589 C	23 21 17.7	+16 29 38	90.08.22	06:15	1.227	29.9	A	16.631	14.965	1.666
684	A2589 D	23 21 23.9	+16 24 21	90.08.20	07:51	1.063	29.9	A	16.663	15.009	1.655
687	A2589 G	23 21 17.2	+16 34 39	91.10.04	03:45	1.186	29.9	A	17.087	15.131	1.956
691	A2593-N A	23 21 49.2	+14 22 20	91.07.10	09:21	1.206	29.9	A	15.511	13.740	1.770
692	A2593-N B	23 21 41.3	+14 20 38	91.10.10	04:26	1.081	29.9	A	16.192	14.532	1.660
693	A2593-N C	23 22 01.2	+14 21 51	90.08.20	08:51	1.048	29.9	A	16.458	14.752	1.705
694	A2593-N D	23 22 06.3	+14 22 02	90.08.22	07:14	1.113	29.9	A	16.072	14.376	1.696
695	A2593-N E	23 22 10.0	+14 21 29	91.07.10	09:40	1.161	29.9	A	15.508	13.791	1.717
696	A2593-N F	23 21 52.1	+14 23 01	91.10.04	05:44	1.049	29.9	A	16.508	14.814	1.694
697	A2593-N G	23 22 08.8	+14 23 32	91.10.04	04:29	1.111	29.9	A	16.155	14.967	1.188
698	A2593-N H	23 22 04.9	+14 24 28	91.10.04	04:56	1.076	29.9	A	16.595	14.868	1.727
699	A2593-N I	23 21 52.7	+14 25 55	90.08.22	06:56	1.145	29.9	A	16.348	14.635	1.713
700	A2593-N J	23 21 41.1	+14 08 56	90.08.20	09:10	1.084	29.9	A	16.326	14.612	1.714
702	A2593-N L	23 23 06.6	+14 17 55	90.08.22	06:36	1.196	29.9	A	16.351	14.548	1.803
703	A2593-N M	23 22 53.8	+14 28 53	91.10.13	04:33	1.064	19.2	A	16.547	14.847	1.701
704	A2593-N 1	23 22 05.9	+14 18 31	91.10.05	04:46	1.086	29.9	A	16.728	15.115	1.612
706	A2593-N 3	23 21 30.5	+14 12 24	91.10.05	05:07	1.066	29.9	A	17.445	15.692	1.753
707	A2593-N 4	23 21 02.6	+14 18 07	91.10.13	04:51	1.053	19.2	A	17.177	15.637	1.540
709	A2593-S A	23 21 55.0	+13 41 49	90.08.20	09:27	1.064	29.9	A	15.204	13.461	1.743
711	A2593-S C	23 21 45.6	+13 41 30	91.10.04	06:10	1.054	29.9	A	16.461	14.763	1.698
712	A2593-S D	23 22 00.8	+13 59 49	91.07.09	08:55	1.309	29.9	A	16.327	14.625	1.702
713	A2593-S E	23 20 55.7	+13 51 24	91.10.13	04:16	1.081	29.9	A	16.023	14.492	1.531
715	A2593-S G	23 21 24.4	+13 38 08	91.07.09	09:49	1.147	29.9	A	17.222	15.551	1.671
717	A2634 B	23 36 08.5	+26 44 02	90.08.22	08:03	1.024	29.9	A	15.833	14.080	1.753
718	A2634 C	23 35 58.9	+26 42 05	90.08.21	10:22	1.059	29.9	A	15.735	14.012	1.723
719	A2634 D	23 35 56.5	+26 42 28	91.07.09	10:13	1.071	29.9	A	15.832	14.065	1.738
720	A2634 E	23 36 20.4	+26 59 24	90.08.21	11:06	1.130	29.9	A	15.329	13.502	1.827
721	A2634 F	23 35 27.6	+26 59 14	91.10.14	03:27	1.113	19.2	A	16.342	14.556	1.785
722	A2634 G	23 37 30.3	+26 51 21	90.08.21	09:35	1.012	29.9	A	14.823	12.940	1.883
723	A2634 H	23 37 30.3	+26 51 21	91.10.14	05:26	1.004	29.9	A	14.824	12.917	1.908
724	A2634 I	23 37 19.0	+27 05 54	90.08.22	08:15	1.015	29.9	A	15.873	14.077	1.795
725	A2634 J	23 37 48.6	+27 17 02	90.08.21	10:40	1.079	29.9	A	15.476	13.699	1.776
726	A2634 K	23 38 23.7	+27 13 54	90.08.22	08:31	1.008	29.9	A	15.813	14.078	1.734
727	A2634 L	23 35 52.5	+26 52 51	91.10.14	05:10	1.005	19.2	A	16.096	14.365	1.730
728	A2634 2	23 35 48.2	+26 36 34	91.10.14	04:54	1.011	19.2	A	16.187	14.448	1.738
730	A2657 B	23 42 10.9	+08 56 16	90.08.23	08:45	1.086	29.9	A	16.836	14.850	1.986
731	A2657 C	23 41 57.5	+08 59 09	91.10.14	07:20	1.206	29.9	A	16.087	14.214	1.873
733	A2657 E	23 41 51.7	+08 53 12	90.08.23	07:20	1.176	29.9	A	16.367	14.519	1.847
734	A2657 F	23 41 43.2	+08 46 15	90.08.22	08:46	1.088	29.9	A	16.281	14.408	1.873
735	A2657 G	23 42 28.7	+08 46 00	90.08.23	08:31	1.093	19.2	A	17.376	15.594	1.782
736	A2657 H	23 42 44.3	+08 59 37	90.08.23	07:38	1.144	29.9	A	16.505	14.641	1.864

Table 3. *BVR* photoelectric data.

ID#	Name	Date	UT	Air	D (")	Q	B	V	R	B-V	B-R
209	P597-1 A	91.10.12	11:08	1.186	29.9	A	14.560	13.446	12.793	1.114	1.767
580	A2199 B	90.08.20	04:06	1.136	29.9	A	15.838	14.844	14.258	0.994	1.580
773	N3379	88.05.09	04:25	1.127	29.9	A	11.828	10.849	10.244	0.979	1.584
773	N3379	88.05.10	03:52	1.088	29.9	A	11.837	10.851	10.245	0.986	1.592
773	N3379	91.05.16	05:19	1.347	29.9	A	11.857	10.866	10.274	0.991	1.583
773	N3379	91.05.17	04:26	1.187	29.9	A	11.848	10.848	10.258	1.000	1.590
773	N3379	88.05.09	04:30	1.135	60.2	A	11.327	10.354	9.750	0.973	1.577
773	N3379	88.05.10	03:57	1.093	60.2	A	11.337	10.362	9.761	0.975	1.576
774	N936	90.08.23	11:20	1.200	29.9	A	12.661	11.685	11.087	0.976	1.574
794	N584	90.08.22	10:32	1.288	29.9	A	12.558	11.577	10.995	0.981	1.563
794	N584	91.10.10	08:18	1.307	29.9	A	12.554	11.574	10.990	0.981	1.565
795	N596	90.08.20	10:04	1.330	29.9	A	13.016	12.066	11.498	0.950	1.518
795	N596	91.10.14	05:38	1.448	29.9	A	13.008	12.076	11.509	0.932	1.498
797	N5846	90.06.19	04:59	1.163	29.9	B	13.093	12.043	11.403	1.051	1.690
797	N5846	90.06.20	05:07	1.169	29.9	B	13.085	12.038	11.401	1.047	1.685
797	N5846	90.06.21	05:11	1.175	29.9	B	13.077	12.023	11.391	1.054	1.686
797	N5846	91.07.10	05:26	1.371	29.9	A	13.008	11.963	11.340	1.045	1.668
797	N5846	91.07.11	04:35	1.234	29.9	A	13.011	11.969	11.338	1.042	1.674
797	N5846	91.07.12	04:45	1.265	29.9	B	13.043	11.994	11.368	1.049	1.675
797	N5846	90.06.19	05:03	1.165	39.5	B	12.807	11.759	11.139	1.048	1.668
797	N5846	90.06.20	05:11	1.172	39.5	B	12.774	11.727	11.096	1.047	1.677
797	N5846	90.06.21	05:15	1.179	39.5	B	12.762	11.715	11.091	1.047	1.671
799	N7626	90.08.20	07:25	1.163	29.9	A	13.625	12.561	11.930	1.063	1.694
799	N7626	90.08.22	09:03	1.098	29.9	A	13.642	12.579	11.951	1.063	1.691
799	N7626	91.07.09	08:42	1.444	29.9	A	13.654	12.582	11.948	1.072	1.706
799	N7626	91.07.10	10:46	1.110	29.9	A	13.611	12.554	11.926	1.057	1.685
799	N7626	91.10.04	06:46	1.116	29.9	A	13.616	12.564	11.950	1.052	1.666
799	N7626	91.10.05	04:33	1.156	29.9	A	13.625	12.570	11.933	1.056	1.692
799	N7626	91.10.10	03:48	1.192	29.9	A	13.626	12.570	11.943	1.056	1.683
799	N7626	91.10.14	04:03	1.139	29.9	A	13.642	12.586	11.954	1.056	1.688
800	N4486	91.05.16	05:25	1.102	29.9	A	12.075	11.059	10.453	1.016	1.622
800	N4486	91.05.17	04:34	1.064	29.9	A	12.086	11.068	10.455	1.018	1.631
801	N224	91.10.05	07:30	1.015	29.9	A	10.250	9.204	8.579	1.046	1.671

in these comparisons are observations of both standard galaxies and programme galaxies. In the results given below, each observation is counted as a separate measurement.

An intercomparison of the 98 duplicate observations taken on quality A or B nights yields 1σ errors per observation of 0.027 mag in *B*, 0.023 mag in *R* and 0.018 mag in *B-R* colour. In comparison, errors from photon statistics alone are 0.007, 0.006 and 0.009 mag respectively. While the scatter in the galaxy magnitudes is dominated by effects other than photon noise, the scatter in the colours is only twice as great as the photon noise. This suggests that the scatter in the repeat measurements is dominated by the photometric stability of the night, consistent with our estimates of 1–3 per cent stability from the standard star observations.

In contrast, the 96 comparisons of nights of quality C both with themselves and with better quality nights yield much larger 1σ errors of 0.063 mag in *B*, 0.042 mag in *R* and 0.034 mag in *B-R*. All of these errors are large compared to the photon statistics. Inspection of the individual comparisons reveals that most of the difference between the quality C nights and the better nights comes in the form of a number of large deviations, often 0.15–0.30 mag in both *B* and *R*. Even when the magnitude deviation is large, however, the amplitude is similar in both passbands, so that the colour is little affected. This implies that extinction variations are the cause of the large scatter in the magnitudes from the quality C nights.

2.2.2 *BVR* versus *BR* photoelectric reductions

For the galaxies with *BR* photometry alone, *R* magnitudes were obtained from the standard stars using only the *B* and *R* magnitudes. For those galaxies with *BVR* photometry, *R* magnitudes were obtained using both the *BR* reduction and a separately determined *BVR* reduction. The 33 *BVR* galaxy observations are given in Table 3, and can be compared with *BR* reductions for the same observations in Table 2. Unfortunately, on the nights of 1988 May 7 and 8, the only nights observed with Cold Box # 27 which were of quality A or B, standard galaxies were not observed in the *V* filter. The observations of NGC 4486 on the night of 1988 May 10 also did not include the *V* filter.

The data for all standard stars were reduced using both a *BR* reduction and a *BVR* reduction. The reduction procedures ensured that there were no systematic differences between the two reduction schemes for the magnitudes and the colours of the standard stars.

The result for the 10 galaxies with *BVR* photometry is somewhat different. In Fig. 1 we plot for the 33 observations of these galaxies the difference in *B*, *R* and *B-R* magnitudes (in the sense of *BR* reduction minus *BVR* reduction) as a function of *B-R* colour from the *BR* reduction. The data are separated into two groups according to photomultiplier and filter set used: four observations made with the GaAs system (circles) and 29 observations made with the S-20 system (squares). The GaAs observations were made during

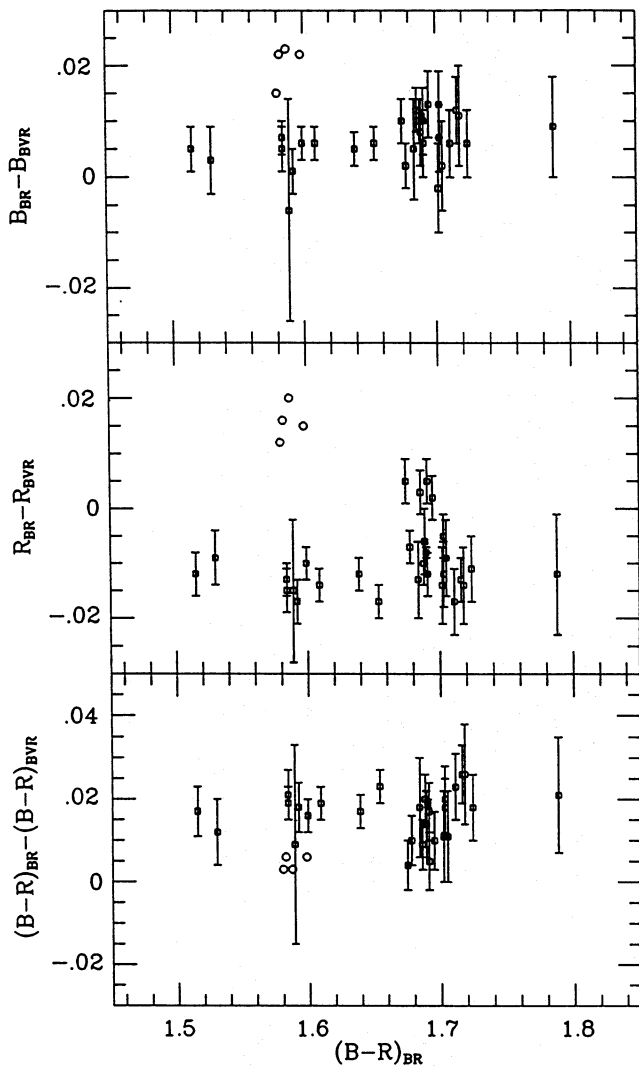


Figure 1. Comparison of *BR* and *BVR* reductions of photoelectric observations of galaxies (circles = GaAs, squares = S-20).

one run of two nights, while the S-20 data were obtained over many runs and many nights. The *BR* reduction for the S-20 system predicts galaxy *B* magnitudes that are 0.007 mag too faint relative to the *BVR* reductions, *R* magnitudes that are -0.012 mag too bright, and *B*–*R* colours that are 0.02 mag too red. The same comparisons for the GaAs observations yield *BR*-predicted *B* magnitudes that are 0.02 mag fainter than the *BVR* *B* magnitudes, *R* magnitudes that are 0.018 mag fainter, and *B*–*R* colours that are 0.005 mag redder.

These systematic differences cannot be due to differences in extinction, as they are consistent for all nights observed, and are not seen for the standard stars. Rather, we believe that these kinds of systematic difference arise from the fact that galaxy spectral energy distributions are not those of stars. The 800-Å widths of the *BVR* passbands apparently allow 1–2 per cent differences in the transformation coefficients between galaxies and stars to occur.

All of the programme galaxy CCD data have been obtained using a *BR* system. Hence it was decided to publish

the photoelectric observations from the *BR* reduction to minimize systematic differences between the photoelectric and CCD data sets. As will be shown later in this paper, the systematic difference in *R* magnitude between the GaAs system and the S-20 system is of similar size to the errors that exist between the CCD and photoelectric data.

2.3 External comparisons

Of all the photoelectric data given in this paper, very little can be compared in a straightforward manner to published data. Few *R*-passband data exist on galaxies, especially on the Landolt photometric system. An accurate external check on the reliability of the photoelectric (or CCD) magnitude requires either a precisely matched aperture or the use of a luminosity profile or growth curve for each galaxy. The latter type of comparison will be deferred to a future paper, in which we will present all of our CCD data.

Only the *B*–*V* colours of those galaxies with measured *BVR* magnitudes are suitable for external comparison in this paper. *BVR* photometry is available, of course, for the standard galaxies. In addition, one programme galaxy, P597-1 A (NGC 1713), was observed in *BVR* and has published *BV* data. The *B*–*V* comparisons are summarized in Table 4. Sources of external photometry are given in the notes to Table 4, and are the same as were used in the analysis of Burstein et al. (1987). The number code for each source is the same as given in table 4 of Burstein et al., and the transformations of the *BV* observations of each source to a ‘standard’ *BV* system are those used in that paper.

A total of 11 separate aperture observations of nine galaxies can have the new measurements of *B*–*V* colour made here compared to existing measurements in the literature. The criterion for comparison is the same as used in Burstein et al., namely that the difference in aperture size between our data and a literature source cannot be greater than 0.05 dex. Most galaxies in this comparison have multiple observations in our survey, and multiple comparisons to literature data. Of the 45 aperture measurements taken from the literature, 24 come directly from the Burstein et al. survey while 21 come from other sources. The mean of the zero-point differences between our new measurements and those from the literature is only -0.007 ± 0.005 mag. The standard deviation of the differences about this value is 0.016 mag. It should also be noted that the galaxy with only one published colour and one observation by us (P597-1 A) also has the largest difference in *B*–*V* colour, which is consistent with the dispersion observed between single observations by two different observers in Burstein et al. (1987).

The accuracy implied for the individual measurements is 0.011 mag, within a factor of 2 of the accuracy obtained for the best standard star observations. The fact that no significant *B*–*V* zero-point offset is found in our data relative to the literature can be principally attributed to the fact that the *BVR* reduction package used here was also used for the data given in Burstein et al. (1987).

To summarize the results of this section, we find by internal comparisons that the *B* and *R* photoelectric magnitudes for our galaxies should be accurate to 2–3 per cent, while the *B*–*R* colours should be accurate to 1–2 per cent. An external check of *B*–*V* colours indicates that these colours are of 1 per cent accuracy when multiply observed.

Table 4. External comparisons of photoelectric photometry.

Name	N_{obs}	log A	$\Delta B-V$	$B-V_{new}$	$\langle B-V_{cat} \rangle$	$B-V_{cat}$	Source
N224	1	0.70	0.001	1.046	1.045	1.04	1
						1.05	20
N584	2	0.70	-0.001	0.980	0.981	0.98	1
						0.98	4
						0.99	5*
						0.97	15
						0.985	20
						0.98	22
						0.98	25
N596	2	0.70	-0.026	0.941	0.967	0.97	1
						0.98	5*
						0.98	15
						0.98	20
						0.95	20
						0.94	23
						1.005	1
N936	1	0.70	-0.028	0.976	1.004	1.000	5*
						1.03	1
						0.98	20
N3379	4	0.70	-0.012	0.988	1.000	1.02	1
						0.99	3
						0.99	21
						1.00	37
						0.975	1
N3379	2	1.00	-0.002	0.977	0.979	0.966	5*
						0.97	7
						1.00	15
						0.97	32
						0.99	37
N4486	2	0.70	0.010	1.017	1.007	1.046	2*
						0.99	3
						0.99	5*
						1.00	21
						1.01	26
						1.06	21
N5846	6	0.70	0.008	1.048	1.040	0.99	5*
						1.07	37
						1.07	37
N5846	3	1.00	0.016	1.051	1.035	1.035	2*
						1.04	7
						1.02	15
						1.04	21
						1.04	37
N7626	7	0.70	-0.011	1.059	1.070	1.05	15
						1.07	20
						1.09	22
						1.06	23
						1.145	27†
P597-1 A	1	0.70	-0.031	1.114	1.145	1.145	27†

Sources: (1) de Vaucouleurs & de Vaucouleurs (1972); (2) Tifft (1961, 1963, 1969, 1973); (3) de Vaucouleurs et al. (1978); (4) Persson, Frogel & Aaronson (1979); (5) Sandage & Visvanathan (1978); (7) Sandage (1972, 1973, 1975); (15, 20–27, 37) Burstein et al. (1987); (32) Kormendy (1977). *Transformed following Burstein et al. (1987). †N1713, average of 0.51 and 0.81 log A apertures.

3 CCD PHOTOMETRY

3.1 Observations

The photometric CCD data for this paper were obtained during two observing runs on the 1.0-m Jacobus Kapteyn Telescope (JKT) at La Palma over the nights of 1990 October 15–17 and 1991 May 9–13. The detectors were two similar GEC CCDs (GEC3 in 1990, GEC6 in 1991), both with a pixel size of 0.30 arcsec, a field of 2.9×1.9 arcmin² and a gain of $1.0 \text{ e}^- \text{adu}^{-1}$. As spatial resolution was not critical, the CCDs were binned 2×2 on-chip to reduce read-out noise. Mould B and R filters (Argyle et al. 1988) were used for both runs.

Since the aim of this project was to obtain photometry for the programme galaxies accurate to better than 3 per cent,

care was taken with both the observing strategy and the reduction procedure. However, there is clearly a trade-off between the amount of observing time spent on standard stars and that spent on the objects of interest. The following observing strategy was therefore adopted to ensure enough calibration data to parametrize the photometric conditions sufficiently, whilst leaving most of the observing time free for our programme galaxies.

Bias frames were taken in the usual way at the beginning and end of each night. Sky flats were also taken each night during evening and morning twilight, as these were found to be superior to dome flats which leave residual structure after flat-fielding. Standard stars, selected from Landolt's (1983) tables of $UBVRI$ standards, were observed in B and R throughout each night. These stars were usually observed out of focus in order to maximize the signal whilst remaining in the linear regime of the CCD's response. A minimum of four or five standard stars were observed at least twice in both B and R at the beginning of the night. These were selected to give a good sky coverage at a range of airmasses and a broad span of $B-R$ colours. Using software available at the telescopes, instrumental magnitudes were computed as the observations were taken. Standard stars were repeatedly observed until we were confident that the atmosphere was stable and photometric across the whole sky.

Consistent with our experience with the photoelectric data, analysis of various CCD runs gave a rule of thumb that fluctuations of more than 0.03 mag in repeat observations of standard stars indicated that conditions were insufficiently photometric. Groups of standard stars would be observed every 60 to 80 min during the night. These groups would typically consist of three stars observed twice in B and R , again selected to give good sky and colour coverage. Five or more standard stars would be observed in the same way at the end of the night. For an observation of a programme galaxy to be declared photometric, it had to be bracketed by two sets of standard star measurements, each with rms errors of less than 0.03 mag.

Table 5 is a log of the observations for each of the nine photometric nights, giving the telescope and CCD together with the photometric periods of the night. The table also lists the number of standard star observations, the number of different standard stars observed, the rms residual of these observations about the fitted photometric transformation (see below) and the number of galaxies imaged in photometric conditions. Exposure times for galaxies ranged from 150 to 600 s. The galaxies were generally only observed in the R filter, although some previously well-studied galaxies were also observed in B in order to permit comparisons with earlier work.

3.2 Data reductions

All CCD images of both standard stars and galaxies underwent the same basic reduction procedure. The images were first bias-subtracted, using the median of the 10 or so bias images taken before and after each night's run scaled to the mean level of the CCD's overscan strip. No correction for dark current was made, as this was negligible on both CCDs. Tests of the CCD response curves showed no evidence for non-linearity over the relevant range of counts. Flat-field corrections were made by dividing each image by the median

Table 5. Log of CCD observations.

Date	Telescope	CCD	Photometric period (UT)	# std obsns	# std stars	rms residual (mag)	# gal obsns
90.10.15	JKT 1m	GEC3	0105–0631	38	9	0.017	12
90.10.16	"	binned 2x2	1952–0259	65	13	0.027	17
90.10.17	"	"	2049–0523	36	7	0.026	9
91.05.09	JKT 1m	GEC6	2347–0521	61	12	0.020	7
91.05.10	"	binned 2x2	2054–0528	57	13	0.010	18
91.05.11	"	"	2037–0351	46	12	0.009	6
91.05.13	"	"	2047–0535	77	17	0.007	28
Total:				380	83		97

of the twilight sky images taken in the appropriate filter. The resulting images had residual flat-field variations of less than 1 per cent rms.

Instrumental magnitudes for the standard stars were determined using either the FIGARO FOTO routine or the IRAF QPHOT task. In both cases the magnitudes were measured within synthetic apertures much larger than the stars' images, which were usually taken deliberately out of focus in order to obtain as high a signal as possible without saturating the CCD. The sky level about each standard star was estimated as the mode of the pixel values in an annulus at large radius about the star. The results of these measurements were in excellent agreement with the instrumental magnitudes obtained using the PHOT on-line software during the observations.

In order to reduce the instrumental magnitudes to the standard *BR* photometric system, we parametrized the photometric transformations as

$$M_{BR} = M_{\text{obs}} - \alpha \times \text{airmass} - \beta \times (B - R) - \gamma - \delta_i,$$

where M_{obs} is the observed stellar magnitude, α is the coefficient of extinction, β is the colour correction, γ is an overall zero-point, and the δ_i are zero-point corrections determined at various times during the night. The δ_i terms are simply additive corrections to the fit for each group of standard star observations, and so provide a means of tracking temporal changes in atmospheric conditions. A linear interpolation between the two δ_i terms bracketing a given galaxy observation is used in computing a photometric zero-point for that observation. This procedure allows changing atmospheric conditions to be modelled as well as the standard star data permit, but care must be taken that there are enough standard stars to ensure that the δ_i terms represent real changes in the extinction.

The parameters of the photometric transformation were determined using an iterative multilinear regression technique. This procedure was as follows. (i) Solve for α , β and γ using multilinear regression. (ii) The δ_i term for each group of standard stars is then simply the mean difference between the measured magnitudes for each group and this solution. On the second and subsequent iterations the δ_i terms from the previous iteration are included in the fit. (iii) The fit is subtracted from the data and the process is repeated either until further iterations do not improve the rms residual by more than some small amount (here 0.001 mag), or until a maximum of 100 iterations is reached. Graphical checks of the solutions were made to avoid bias from 'wild' observations. We have tested this procedure with simulated data and established that it leads to accurate and reliable solutions.

Our photometric data typically required 8–12 iterations to converge to a solution.

The method we use to determine a galaxy's integrated magnitude from a CCD image is quite different from the direct measurement of the light in circular apertures employed by photoelectric photometry. Although this latter approach could be mimicked by using synthetic apertures, the real aim of the CCD photometry is to determine the galaxy's surface brightness profile accurately, uncontaminated by superimposed stars or galaxies. Before determining magnitudes, we therefore delete contaminating objects by setting all the pixels in regions fully covering their images to a 'magic' value, indicating that the flux in these pixels is undetermined. A comparison of the CCD magnitudes with and without deletions is made in Section 4. The sky level is estimated as the mode of the pixel values in selected areas on the CCD frame, excluding the deleted regions and the area covered by the target galaxy. Given that almost all the galaxies are small compared to the size of the CCD frame (with the exception of # 773 = NGC 3379), this method gives robust and precise estimates of sky – for the 1990 October run the mean estimated error in the sky level was 0.8 per cent, while for the 1991 May run it was 1.0 per cent. The centre of the galaxy's image is then established by centroiding over a small region about the peak of the light. The circularized surface brightness profile about this centre is computed as the azimuthal average of the flux in radial steps of one pixel. The flux at any given radius and azimuth is estimated by bilinear interpolation using the four nearest pixels. If the flux in any of these four pixels is undetermined (i.e. the pixel is set to the 'magic' value), then the flux being estimated is also considered to be undetermined. At a radius of N pixels, the azimuthally averaged flux is computed as the mean over approximately $2\pi N$ points uniformly distributed in azimuth (so the sampled points are approximately one pixel apart). Points at which the flux is undetermined are omitted when calculating the mean. The resulting circularized surface brightness profile is converted to instrumental magnitudes and then integrated out to the appropriate radius (using linear interpolation) in order to compare the CCD magnitude to the photoelectric magnitude within the corresponding aperture.

Instrumental CCD magnitudes determined in this way were transformed on to the standard system using the extinction coefficient, colour term and overall zero-point for the night. A zero-point correction was also applied, as interpolated from the corrections derived for the two groups of stars bracketing the galaxy observation. Since only an *R* image was obtained for the galaxies, we assumed a *B*–*R* colour of 1.7

in every case. As the photoelectric data show, this is the mean colour for these E and S0 galaxies. The bluest galaxy in the sample has $B-R=1.37$, while the reddest has $B-R=1.96$; the standard deviation about the mean colour is 0.11 mag. The mean absolute value of the colour terms we obtain is 0.010; the maximum is 0.021. Thus the typical error we make by assuming all our sample have $B-R=1.7$ is 0.001 mag, and the maximum possible error for a single galaxy is 0.007 mag.

4 COMPARISON OF PHOTOELECTRIC AND CCD PHOTOMETRY

There are 95 programme galaxies for which we can compare photoelectric and CCD photometry. Table 6 summarizes this comparison. The name, date, time and airmass of the CCD observation are given, together with the diameter (in arcsec) of the photometric aperture. This diameter is chosen to match the aperture used in the photoelectric observations of the galaxy. There are 97 comparisons listed in Table 6: NGC 3379 (# 773) was observed with two different photoelectric apertures and separate CCD magnitudes were computed for each; J11 E (# 306) has two photometric CCD observations, so differences from the photoelectric magnitude are given for both. For galaxies with more than one photoelectric observation at the same aperture, the observations were averaged before being compared to the CCD magnitude. Table 6 lists the mean R_{PE} and $(B-R)_{PE}$, the difference $\Delta R = R_{CCD} - R_{PE}$ for each comparison and the quality of the nights on which the photoelectric and CCD data were taken (indicated by class A or B for the photoelectric data and by the rms residual of the calibration for the CCD data).

The distribution of these differences is shown in Fig. 2. The mean difference between the CCD and photoelectric magnitudes for all 97 comparisons is $\overline{\Delta R} = -0.004$ mag and the standard deviation about the mean is $\sigma_{\Delta R} = 0.036$ mag. Although the distribution appears double-peaked, a battery of tests for deviations from a Gaussian distribution (part of the ROSTAT software: Beers, Flynn & Gebhardt 1990) shows no inconsistency with a single Gaussian distribution at even the 10 per cent confidence level. Since the rms errors in the photoelectric magnitudes estimated from repeat observa-

tions are 0.02–0.03 mag (Section 2.2.1), the scatter of 0.036 mag in ΔR implies that the rms precision of the CCD magnitudes is also 0.02–0.03 mag.

For the 38 comparisons using data from the 1990 October CCD run, we find $\overline{\Delta R} = -0.010$ mag and $\sigma_{\Delta R} = 0.030$; for the 59 comparisons using the 1991 May CCD data, $\overline{\Delta R} = 0.000$ mag and $\sigma_{\Delta R} = 0.039$. The mean ΔR values for these two runs differ by only 1.4 times their joint standard error. Fig. 3 shows how ΔR varied with time during each night of CCD observations. Rank correlation tests (Spearman's ρ and Kendall's τ) show no evidence for any night having a significant trend of ΔR with time. Two nights, both with fewer than 10 comparisons, show small but statistically significant overall offsets: 1990 October 17 has $\overline{\Delta R} = -0.022$ mag (2.8 times the standard error in the mean) and 1991 May 9 has $\overline{\Delta R} = +0.031$ mag (2.6 times the standard error in the mean). The scatter ΔR on all nights was in the range $\sigma_{\Delta R} = 0.02$ –0.04 mag.

Both quality A and quality B photoelectric observations are consistent with zero offset between the two photometric systems. For the 90 comparisons using photoelectric data of quality A, $\overline{\Delta R} = -0.004$ mag and $\sigma_{\Delta R} = 0.036$ mag, while for the seven quality B comparisons $\overline{\Delta R} = -0.010$ mag and $\sigma_{\Delta R} = 0.043$ mag.

Fig. 4(a) shows that the fainter ($14 < R < 15$) galaxies in the sample have a scatter $\sigma_{\Delta R} = 0.038$ mag, close to the value for the sample as a whole, while the brighter ($R < 14$) galaxies show significantly less scatter: $\sigma_{\Delta R} = 0.026$ mag. The increase in the scatter with R magnitude is consistent with the estimated error of ~ 1 per cent in the sky level, since in the 19.2- and 29.9-arcsec apertures the sky flux becomes comparable to that from the galaxy at around $R = 14$. By $R = 15$ the error in sky contributes a couple of per cent to the overall magnitude error, resulting in the observed increase in the scatter. There is no trend of $\overline{\Delta R}$ with R magnitude.

Neither the mean nor the scatter of the differences varies with the galaxy's $B-R$ colour (Fig. 4b), the airmass of the CCD observation (Fig. 4c) or the aperture size used to make the measurement (Fig. 4d). The relatively large difference found for the one comparison at an aperture diameter of 60.2 arcsec (# 773 = NGC 3379, $\Delta R = +0.065$ mag) is probably due to the fact that this large galaxy nearly fills the

Table 6. CCD photometry and comparison with photoelectric photometry.

ID#	CCD Image	Date	UT	Air	D (")	R_{CCD}	R_{PE}	$(B-R)_{PE}$	$R_{CCD}-R_{PE}$	Q	rms
1	A76_A_582	90.10.16	00:38	1.09	29.9	13.375	13.379	1.846	-0.004	A	0.027
2	A76_B_583	90.10.16	00:40	1.09	29.9	14.171	14.206	1.695	-0.035	A	0.027
4	A76_D_585	90.10.16	00:48	1.09	29.9	14.034	14.023	1.755	+0.011	A	0.027
96	A260_A_350	90.10.15	01:55	1.01	29.9	13.322	13.340	1.776	-0.018	A	0.017
98	A260_C_351	90.10.15	02:03	1.01	29.9	13.723	13.712	1.760	+0.011	A	0.017
100	A260_E_353	90.10.15	02:17	1.02	29.9	14.012	14.020	1.735	-0.008	A	0.017
101	A260_F_354	90.10.15	02:24	1.03	29.9	13.924	13.935	1.735	-0.011	A	0.017
102	A260_G_355	90.10.15	02:30	1.03	19.2	14.292	14.273	1.762	+0.019	A	0.017
122	J8_A_601	90.10.16	01:54	1.00	29.9	13.778	13.741	1.803	+0.037	A	0.027
127	J8_F_604	90.10.16	02:08	1.01	29.9	14.177	14.184	1.903	-0.007	A	0.027
128	J8_G_605	90.10.16	02:12	1.01	19.2	14.225	14.240	1.826	-0.015	A	0.027
130	J8_I_607	90.10.16	02:21	1.01	19.2	14.665	14.690	1.937	-0.025	A	0.027
137	A376_C_728	90.10.17	02:24	1.01	29.9	14.886	14.904	1.759	-0.018	A	0.026
151	A397_A_733	90.10.17	02:47	1.03	19.2	14.403	14.414	1.859	-0.011	A	0.026
153	A397_C_734	90.10.17	02:51	1.03	19.2	14.444	14.467	1.808	-0.023	A	0.026
154	A397_D_735	90.10.17	02:55	1.04	19.2	14.639	14.705	1.831	-0.066	A	0.026
155	A397_E_734	90.10.17	02:51	1.03	29.9	14.816	14.793	1.782	+0.023	A	0.026
158	A397_H_732	90.10.17	02:43	1.03	19.2	14.915	14.942	1.909	-0.027	A	0.026
160	A400_A_367	90.10.15	03:16	1.10	29.9	13.775	13.798	1.869	-0.023	A	0.017
163	A400_D_368	90.10.15	03:22	1.11	29.9	14.370	14.374	1.900	-0.004	A	0.017

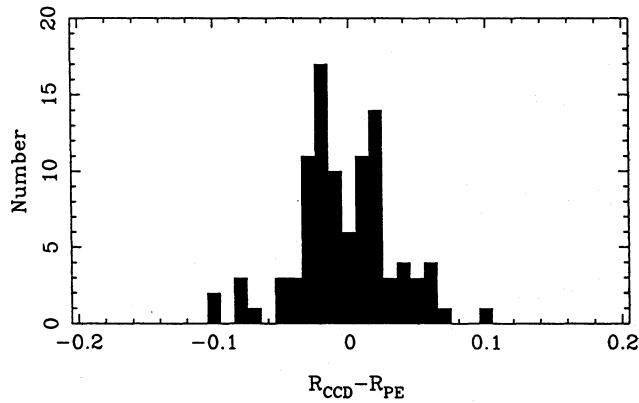


Figure 2. The distribution of the difference between CCD and photoelectric R magnitudes for 97 comparisons at various apertures of 95 galaxies.

CCD frame. The estimate of the sky level may therefore be biased upwards, leading to over-subtraction of sky and an overly faint CCD magnitude, especially in larger apertures (in the 29.9-arcsec aperture this galaxy has $\Delta R = +0.006$).

We have checked the effect on the photoelectric magnitudes of contaminating sources within the photometric apertures by also deriving CCD magnitudes *without* deleting nearby objects. We find that only nine of the galaxies in Table 6 have their R_{CCD} changed by more than 0.02 mag. When the CCD magnitude *without* deletions is used in the comparison with the photoelectric magnitudes, only six of these nine actually show a smaller ΔR and the overall scatter is not improved. Examination of the CCD images shows that not deleting other objects makes little difference to the comparison simply because the original apertures used in the photoelectric photometry were well chosen to avoid bright contaminating sources.

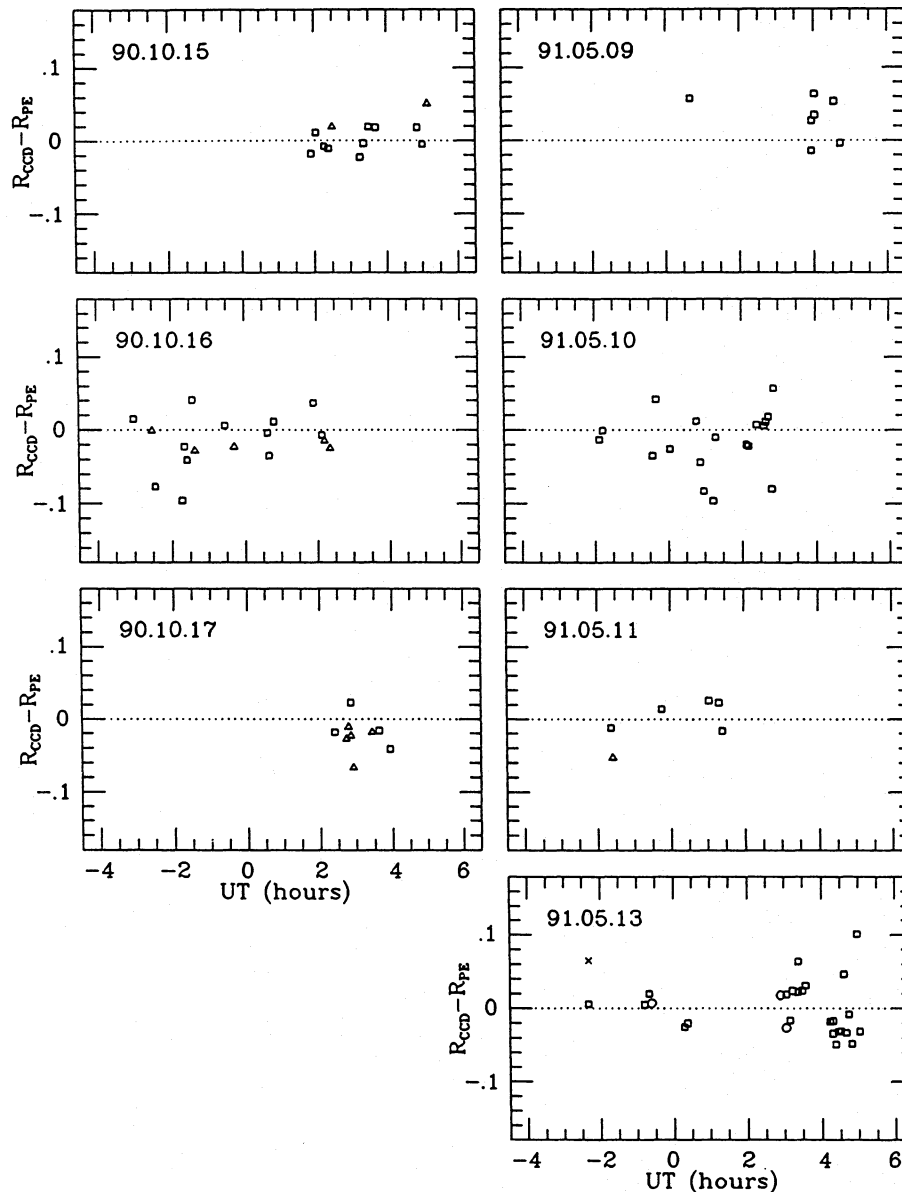


Figure 3. The temporal variation in R_{CCD} on each of the nights with CCD photometry, as shown by the differences $R_{\text{CCD}} - R_{\text{PE}}$. The different symbols indicate aperture sizes: triangles = 19.2 arcsec; squares = 29.9 arcsec; circles = 39.5 arcsec, and crosses = 60.2 arcsec.

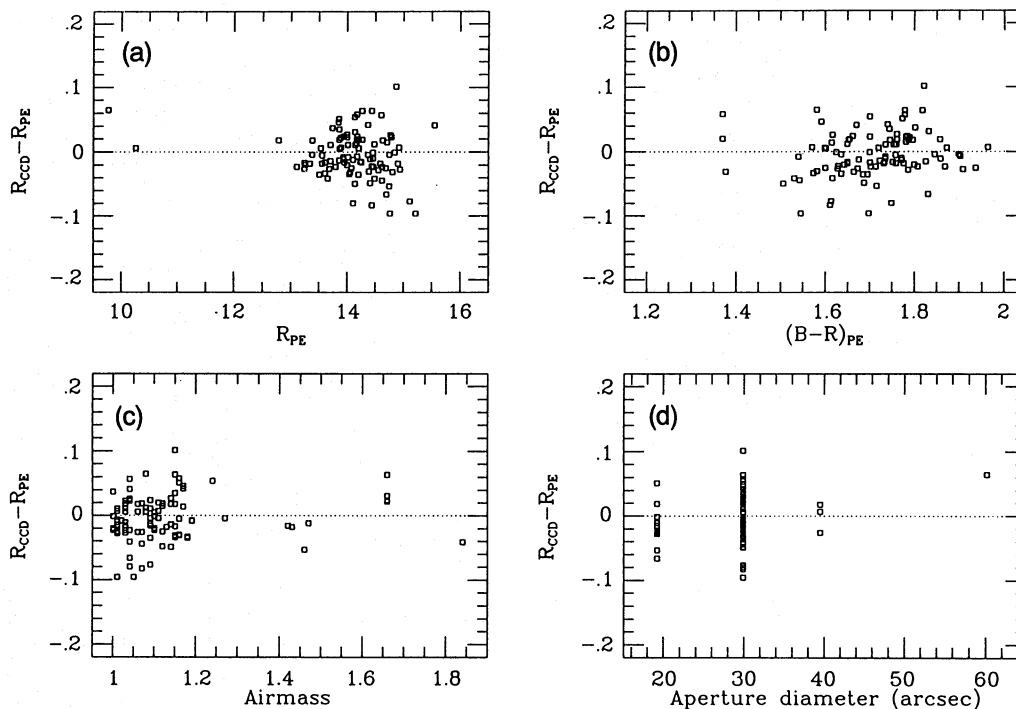


Figure 4. The distribution of differences $R_{\text{CCD}} - R_{\text{PE}}$ with (a) photoelectric magnitude R_{PE} , (b) $(B-R)_{\text{PE}}$ colour, (c) the airmass of the CCD observation, and (d) the aperture diameter (in arcsec) used when measuring the magnitudes.

5 CONCLUSIONS

In this paper we have presented *BR* photoelectric photometry for 352 E and S0 galaxies and *BVR* photoelectric photometry for 10 galaxies. These observations are part of a large programme to study the properties and peculiar motions of these galaxies. The aims of this paper are to show, first, that even for relatively distant, faint galaxies ($R \lesssim 15$) it is possible to obtain magnitudes with a precision of at least 3 per cent and, secondly, that a photometric system can be established for both photoelectric and CCD photometry that has a common zero-point to within 1 per cent and negligible colour terms for E and S0 galaxies.

A comparison of the 98 repeat *BR* photoelectric observations in our data set shows the rms errors on individual *B* and *R* magnitudes to be 2–3 per cent and on *B* – *R* colours to be 1–2 per cent. We have compared the results obtained for the 33 *BVR* observations using both a *BVR* reduction and a *BR* reduction. We find small (0.01–0.02 mag) systematic effects that are attributable to the differences between the spectral energy distributions of stars and galaxies. We have compared the *B* – *V* colours of the galaxies with *BVR* photometry to values from the literature. The mean zero-point difference is only -0.007 ± 0.005 mag. The inferred precision for the colours of individual galaxies is 0.01 mag.

We have also compared photoelectric and CCD *R* magnitudes (both reduced on the *BR* system) for 95 galaxies. We give a detailed description of our CCD observing strategy and our methods for reduction of the CCD data and derivation of aperture magnitudes free from contaminating sources. We show that observation of the galaxies in the *R* band only and assumption of a mean *B* – *R* colour introduces negligible errors in our CCD magnitudes. We find that there is excellent overall agreement (better than 1 per cent) between the zero-

points of the photoelectric and CCD *R* magnitudes. Two of the seven nights of CCD observations, however, show zero-point differences from the photoelectric photometry of 2–3 per cent. The rms difference between the photoelectric and CCD magnitudes is 0.036 mag which, together with the estimated rms internal error of 2–3 per cent for the photoelectric photometry, implies that the rms error on individual CCD measurements is also 2–3 per cent. There is no trend in either the mean or the scatter of the differences with magnitude, colour, airmass or photometric aperture, except that brighter galaxies have slightly less scatter than the sample as a whole.

We thus have established a *BR* photometric system for photoelectric photometry which can also be applied to *R*-band CCD photometry of early-type galaxies, with a zero-point agreement that is good to better than 1 per cent. We have also shown that we can obtain aperture magnitudes, both photoelectrically and with CCDs, which have a precision of 2–3 per cent even for distant galaxies as faint as $R \approx 15$. Errors in galaxy magnitudes of this order introduce errors of $\lesssim 2$ per cent into D_n and the resulting distance estimates from the D_n – σ relation (Dressler et al. 1987), corresponding to errors in peculiar velocities of $\lesssim 200$ km s $^{-1}$ for galaxies at 10 000 km s $^{-1}$. However, photometric errors are unlikely to be the limiting factor in the precision of our D_n – σ distance estimates, since a precision of 2 per cent would be required in the measurement of velocity dispersions to achieve an equally small contribution to the errors from σ .

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