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Aaron Dotter
Dartmouth College

Brian Chaboyer
Dartmouth College

Darko Jevremović
University of Oklahoma

E. Baron
University of Oklahoma

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THE ACS SURVEY OF GALACTIC GLOBULAR CLUSTERS. II. STELLAR EVOLUTION TRACKS, ISOCHRONES, LUMINOSITY FUNCTIONS, AND SYNTHETIC HORIZONTAL-BRANCH MODELS¹

AARON DOTTER AND BRIAN CHABOYER

Department of Physics and Astronomy, Dartmouth College, Hanover, NH 03755, USA

DARKO JEVREMOVIĆ² AND E. BARON³

Homer L. Dodge Department of Physics and Astronomy, University of Oklahoma, Norman, OK 73019-2061, USA

JASON W. FERGUSON

Physics Department, Wichita State University, Wichita, KS 67260-0032, USA

ATA SARAJEDINI

Department of Astronomy, University of Florida, Gainesville, FL 32611, USA

AND

JAY ANDERSON

Department of Physics and Astronomy, Rice University, Houston, TX 77005, USA

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ABSTRACT

The ACS Survey of Galactic Globular Clusters, an *HST* Treasury Project, will deliver high-quality, homogeneous photometry of 65 globular clusters. This paper introduces a new collection of stellar evolution tracks and isochrones suitable for analyzing the ACS survey data. Stellar evolution models were computed at $[\text{Fe}/\text{H}] = -2.5, -2.0, -1.5, -1.0, -0.5$, and 0 ; $[\alpha/\text{Fe}] = -0.2, 0, 0.2, 0.4, 0.6$, and 0.8 ; and three initial He abundances for masses from 0.1 to $1.8 M_{\odot}$ and ages from 2 to 15 Gyr. Each isochrone spans a wide range in luminosity, from $M_V \sim 14$ up to the tip of the red giant branch. These are complemented by a set of He-burning tracks that extend from the zero-age horizontal branch to the onset of thermal pulsations on the asymptotic giant branch. In addition, a set of computer programs are provided that make it possible to interpolate the isochrones in $[\text{Fe}/\text{H}]$, generate luminosity functions from the isochrones, and create synthetic horizontal-branch models. The tracks and isochrones have been converted to the observational plane with two different color- T_{eff} transformations, one synthetic and one semiempirical, in ground-based B , V , and I , and F606W and F814W for both ACS WFC and WFPC2 systems. All models and programs presented in this paper are available at the Dartmouth Stellar Evolution Database and the Multimission Archive at the Space Telescope Science Institute.

Key words: globular clusters: general — stars: evolution

Online material: color figures

1. INTRODUCTION

The ACS Survey of Galactic Globular Clusters (Sarajedini et al. 2007; J. Anderson et al. 2007, in preparation) will set photometric standards for years to come. The survey has obtained high-quality photometry of 65 globular clusters from the horizontal branch (HB) down the main sequence to at least $\sim 0.2 M_{\odot}$. The data from this survey will allow for revised absolute and differential age analyses and studies of the stellar luminosity, mass functions, and dynamical evolution of a large number of clusters in a homogeneous data set, among many other things. The data will provide a critical test for some aspects of current stellar evolution calculations, in particular the shape and luminosity function of the lower main sequence. The survey will eventually combine ACS WFC and WFPC2 photometry in the F606W and F814W bands and ground-based B , V , and I photometry for most of the clusters. In order to analyze these data, it is preferable to have theoretical models transformed to the appropriate photo-

metric system that can be applied directly to the observations. Sirianni et al. (2005) provide a detailed analysis of all aspects of the ACS, including filter transmission curves that account for the performance of the system from optics to electronics, thereby allowing accurate synthetic photometry to be produced from synthetic fluxes (e.g., Bedin et al. 2005).

There are a number of groups currently producing isochrones similar to those introduced in this paper; chief among these are Bergbusch & Vandenberg (2001); Vandenberg et al. (2006, hereafter Victoria-Regina); Girardi et al. (2002); Bonatto et al. (2004); Girardi et al. (2004, hereafter Padova); Pietrinferni et al. (2004, 2006, hereafter BaSTI); Cordier et al. (2007); Yi et al. (2003, 2004); and Demarque et al. (2004, hereafter Y^2 [Yonsei-Yale]). All of these provide isochrones for ground-based B , V , and I bands, but only the Padova and BaSTI libraries include isochrones on the *HST* photometric systems. Only the Padova library extends past the lower mass limit of the ACS survey data; the others terminate at $0.4 M_{\odot}$ (Y^2) or $0.5 M_{\odot}$ (Victoria-Regina and BaSTI).

In order to analyze the full complement of photometry, a new set of stellar evolution models has been produced. This paper describes in detail the stellar evolution code used to create the models in § 2; the color- T_{eff} transformations used to convert the models to the observational plane in § 3; properties of the model

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² Current address: Astronomical Observatory, 11160 Belgrade, Serbia.

³ Computational Research Division, Lawrence Berkeley National Laboratory, Berkeley, CA 94720, USA.

TABLE 1
INITIAL COMPOSITIONS

[Fe/H]	$Y = 0.245 + 1.54Z$						$Y = 0.33$		$Y = 0.4$	
	$[\alpha/\text{Fe}] = -0.2$	0	0.2	0.4	0.6	0.8	0	0.4	0	0.4
-2.5	0.7549	0.7548	0.7548	0.7547	0.7546	0.7543	0.6699	0.6699	0.6000	0.5999
	4.09E-5	5.48E-5	7.63E-5	1.11E-4	1.65E-4	2.52E-4	4.85E-5	9.85E-5	4.34E-5	8.82E-5
-2.0	0.7547	0.7545	0.7544	0.7541	0.7536	0.7529	0.6698	0.6697	0.5999	0.5997
	1.30E-4	1.72E-4	2.41E-4	3.50E-4	5.24E-4	7.97E-4	1.53E-4	3.50E-4	1.37E-4	2.79E-4
-1.5	0.7539	0.7536	0.7530	0.7521	0.7507	0.7484	0.6695	0.6690	0.5996	0.5991
	4.10E-4	5.47E-4	7.62E-4	1.11E-3	1.65E-3	2.50E-3	4.85E-4	9.83E-4	4.34E-4	8.81E-4
-1.0	0.7516	0.7505	0.7487	0.7459	0.7415	0.7346	0.6685	0.6669	0.5986	0.5972
	1.29E-3	1.72E-3	2.40E-3	3.46E-3	5.15E-3	7.77E-3	1.53E-3	3.46E-3	1.37E-3	2.78E-3
-0.5	0.7444	0.7441	0.7355	0.7270	0.7139	0.6942	0.6652	0.6603	0.5957	0.5913
	4.05E-3	5.37E-3	7.44E-3	1.07E-2	1.57E-2	2.32E-2	4.82E-3	9.70E-3	4.31E-3	8.69E-3
0.0	0.7174	0.7071	0.6882	0.6628	0.6247	0.5710	0.6550	0.6402	0.5866	0.5734
	1.43E-2	1.89E-2	2.55E-2	3.52E-2	4.97E-2	7.02E-2	1.50E-2	2.98E-2	1.34E-2	2.66E-2

NOTE.—For each value of [Fe/H], the first row gives the X -value and the second row gives the Z -value.

database, including descriptions of the tracks, isochrones, luminosity functions, and synthetic HB models, in § 4; comparisons of the isochrones with other sets of isochrones and data from the ACS survey in § 5; and, finally, a summary of the results in § 6. Following the main text is an Appendix that describes in detail the stellar evolution track, isochrone, luminosity function, and synthetic HB model files and outlines the use of computer codes distributed with the models.

2. THE DARTMOUTH STELLAR EVOLUTION PROGRAM

The stellar evolution code described here is the Dartmouth Stellar Evolution Program (DSEP). Many details of DSEP have been previously described by Chaboyer et al. (2001) and Bjork & Chaboyer (2006). This section is intended to explain the physics and numerical methods used in the code that have not been covered by Chaboyer et al. (2001) and Bjork & Chaboyer (2006).

Stellar evolution models have been created for [Fe/H] = -2.5, -2.0, -1.5, -1.0, -0.5, and 0 with $[\alpha/\text{Fe}] = -0.2, 0, 0.2, 0.4, 0.6, \text{ and } 0.8$. In the models, α -enhancement refers to enhancements in the α -capture elements O, Ne, Mg, Si, S, Ca, and Ti by the same amount as specified in $[\alpha/\text{Fe}]$. The models assume that the initial He mass fraction follows $Y = 0.245 + 1.54Z$, where $Y = 0.245$ is the primordial He abundance from big bang nucleosynthesis (Spergel et al. 2003) and the required initial He abundance of the calibrated solar model is $Y = 0.274$. In addition, sets of models were produced with $Y_{\text{init}} = 0.33$ and 0.4 for the same range of [Fe/H] with $[\alpha/\text{Fe}] = 0$ and 0.4. Complete details of the initial compositions for all models are listed in Table 1.

For future reference, the solar abundances were adopted from Grevesse & Sauval (1998) and the global solar parameters luminosity, radius, surface composition, and age from Bahcall et al.

(2005). Table 2 defines the requirements for a calibrated solar model and gives the values obtained from the DSEP solar-calibrated model. The calibrated solar model has a mixing length parameter $\alpha_{\text{ML}} = 1.938$ with initial composition $X = 0.7071$ and $Z = 0.01885$. Further details of the physics employed by DSEP can be found in the remainder of this section and in Chaboyer et al. (2001) and Bjork & Chaboyer (2006).

2.1. Convective Core Overshoot

The adopted convective core overshoot (CCO) treatment is that of Y^2 . The basic idea behind the method is that the extent of CCO depends on the size of the convective core. The lowest mass main-sequence stars that develop convective cores have relatively small ones. As stellar mass increases beyond this minimum, the convective cores become fully established and the amount of overshoot increases.

The extent of CCO is treated as a fraction of the pressure scale height at the core boundary. At the minimum mass (listed for all compositions in Table 3) the amount of CCO is 5% of the pressure scale height. Between the minimum and $0.2 M_{\odot}$ above the minimum, the extent of CCO is ramped linearly from 5% to 20%. For masses greater than $0.2 M_{\odot}$ above the minimum, the extent of CCO is held constant at 20% of the pressure scale height. CCO plays a role in only the youngest isochrones presented in this paper, those with ages less than about 5 Gyr.

2.2. Microscopic Diffusion and Gravitational Settling

Chaboyer et al. (2001) described the implementation of diffusion and gravitational settling in DSEP; the following is a brief summary. The basic formalism is that of Thoul et al. (1994), but diffusion has been inhibited in the outer $0.1 M_{\odot}$. In the outermost $0.05 M_{\odot}$ diffusion has been completely stopped, and in the $0.05 M_{\odot}$ below that diffusion has been linearly ramped from zero to the full effect. This approach is meant to satisfy the lack of observational evidence for diffusion in metal-poor stars (e.g., Gratton et al. 2001; Ramírez & Cohen 2003), although recently, evidence for some diffusion has been reported in NGC 6397 (Korn et al. 2006).

2.3. Equation of State

The basic equation of state (EOS) employed by DSEP is a general ideal gas EOS with the Debye-Hückel correction and performs sufficiently well for stellar models with masses greater than about $0.7 M_{\odot}$ (Chaboyer & Kim 1995). Because this EOS

TABLE 2
SOLAR PARAMETERS

Parameter	Adopted Value	Solar Model Value
Age (Gyr).....	4.57	...
M_{\odot} (g).....	1.9891×10^{33}	...
R_{\odot} (cm).....	6.9598×10^{10}	$\log(R/R_{\odot}) = 0.0001$
L_{\odot} (ergs s ⁻¹).....	3.8418×10^{33}	$\log(L/L_{\odot}) = 0.0005$
R_{CZ}/R_{\odot}	0.713 ± 0.001	0.716
Z/X	0.0229	0.02288

NOTE.—Values are from Bahcall et al. (2005) and Grevesse & Sauval (1998).

TABLE 3
MINIMUM MASS IN M_{\odot} FOR CONVECTIVE CORE OVERSHOOT

[Fe/H]	$Y = 0.245 + 1.54Z$						$Y = 0.33$		$Y = 0.4$	
	$[\alpha/\text{Fe}] = -0.2$	0	0.2	0.4	0.6	0.8	0	0.4	0	0.4
-2.5	1.6	1.6	1.6	1.5	1.3	1.3	1.5	1.5	1.5	1.5
-2.0	1.4	1.4	1.4	1.4	1.3	1.3	1.5	1.5	1.5	1.5
-1.5	1.4	1.4	1.4	1.4	1.3	1.3	1.4	1.4	1.4	1.4
-1.0	1.4	1.4	1.4	1.4	1.2	1.2	1.4	1.4	1.4	1.4
-0.5	1.2	1.2	1.2	1.2	1.2	1.0	1.2	1.2	1.2	1.2
0.0	1.2	1.2	1.1	1.1	1.0	0.9	1.2	1.1	1.2	1.1

neglects many nonideal effects that become important at lower temperatures and higher densities (such as are found in low-mass main-sequence stars), the basic EOS was used only for models with masses greater than $0.8 M_{\odot}$.

For stellar models with $0.1 M_{\odot} \leq M \leq 0.8 M_{\odot}$ the detailed EOS code developed by A. Irwin, FreeEOS (in the EOS4 configuration), was used. FreeEOS takes into account many nonideal effects and is generally superior for high-density, low-temperature environments. Furthermore, FreeEOS takes into account the makeup of Z . Stellar models produced with this EOS are in good agreement with observations of low-mass stars (see § 5). Irwin has written extensively about the numerical methods and physical models employed by FreeEOS; the interested reader should consult the FreeEOS Web site.⁴

Beginning at approximately $0.7 M_{\odot}$, stellar evolution tracks produced with FreeEOS and the default EOS begin to differ in the H-R diagram. As stellar mass decreases, these differences increase, becoming substantial at $0.1 M_{\odot}$. The current version of FreeEOS (2.0) does not operate in the parameter range occupied by main-sequence models below $0.1 M_{\odot}$.

2.4. Opacities

Opacities were adopted from two sources.⁵ In both cases, separate sets of tables were generated for each of the $[\alpha/\text{Fe}]$ mixtures listed above. The OPAL opacities (Iglesias & Rogers 1996) were used at high temperatures ($\log T > 4.5$), while opacities computed for this effort following Ferguson et al. (2005) were used at low temperatures ($\log T < 4.3$). Between $\log T = 4.3$ and 4.5 the opacity was ramped between the two sources. Ferguson et al. (2005) have shown that OPAL opacities agree with their own to within 5% in this temperature range.

During core and shell He fusion (HB and asymptotic giant branch [AGB] evolution) the OPAL type 2 opacity tables were used. The type 2 opacity tables allow mass fractions of two individual elements (in this case C and O) to vary independent of overall Z . These opacity tables are particularly helpful in the case of He-burning evolution, where the He core gradually becomes a mixture of almost entirely C and O over the course of ~ 100 Myr. Conductive opacities were adopted from Hubbard & Lampe (1969) in the nonrelativistic limit and Canuto (1970) in the relativistic limit via the fitting formulae of Sweigart (1973).

2.5. Surface Boundary Conditions

It is typical in stellar evolution codes to see surface boundary conditions treated as simple mathematical models of surface temperature as a function of optical depth. The almost universal

choices are the Eddington approximation or the empirical fit to the Sun done by Krishna Swamy (1966).

However, particularly in the case of cool, low-mass main-sequence stars, it is common practice to derive surface boundary conditions from model atmosphere calculations (Baraffe et al. 1997). Model atmosphere boundary conditions are an important ingredient in models of cool or very low mass stars ($M < 0.4 M_{\odot}$) but make little or no difference for hotter, more massive stars.

DSEP has been modified to accept surface boundary conditions derived from PHOENIX model atmospheres (Hauschildt et al. 1999a, 1999b). One of the strengths of the PHOENIX code is its attention to the effects of molecules and grains that appear at low temperatures. Indeed, the low-temperature opacities used in this paper were computed with a modified version of the PHOENIX code (Ferguson et al. 2005). A grid of stellar atmosphere models was produced covering the range of $[\text{Fe}/\text{H}]$ and $[\alpha/\text{Fe}]$ listed above based on the Grevesse & Sauval (1998) solar abundances. The grid spans surface gravity $-0.5 \leq \log g \leq 5.5$ and effective temperature $2500 \text{ K} \leq T_{\text{eff}} \leq 10,000 \text{ K}$. For the hottest tracks, model atmosphere grids from Castelli & Kurucz (2003) were supplemented for $T_{\text{eff}} > 10,000 \text{ K}$. A coarser grid of model atmospheres was generated to explore the effect of enhancing the He abundance on atmosphere structure and synthetic flux. While additional He in the atmosphere has a measurable effect on the gas pressure, there is essentially no change in the synthetic flux. Because He is a noble gas, it is difficult to excite for T_{eff} below $10,000 \text{ K}$ and thus contributes little to the flux. At present, He-enhancement is treated as an adjustment to the gas pressure that serves as the surface boundary condition but is ignored in the synthetic colors.

The surface boundary condition is satisfied by specifying pressure as a function of T_{eff} and $\log g$. To obtain the pressure from each model atmosphere, the location within the atmosphere structure where $T = T_{\text{eff}}$ was found, and the gas pressure was extracted from that location. This approach differs from those in which the pressure is extracted from the model atmosphere at a constant optical depth. Before deciding to follow the adopted method, tests were performed by comparing surface boundary conditions extracted from the atmospheres at a constant optical depth ($\tau = 2/3$ and 10), but the tests showed that the $T = T_{\text{eff}}$ method produced the best results when compared with observations. The physical approximations made by model atmosphere and stellar evolution codes are different enough that any choice of where to connect the interior to the atmosphere will encounter some inconsistencies.

Calibrated solar models using the Eddington, Krishna Swamy, and PHOENIX surface boundary conditions were produced in order to compare the three boundary conditions in the H-R diagram. The Krishna Swamy model required the largest mixing length ($\alpha_{\text{ML}} = 2.02$) and had the hottest red giant branch. The

⁴ See <http://freeeos.sourceforge.net>.

⁵ For OPAL opacities, see <http://www-phys.lnl.gov/Research/OPAL/>. For Ferguson et al. (2005) opacities, see <http://webs.wichita.edu/physics/opacity/>.

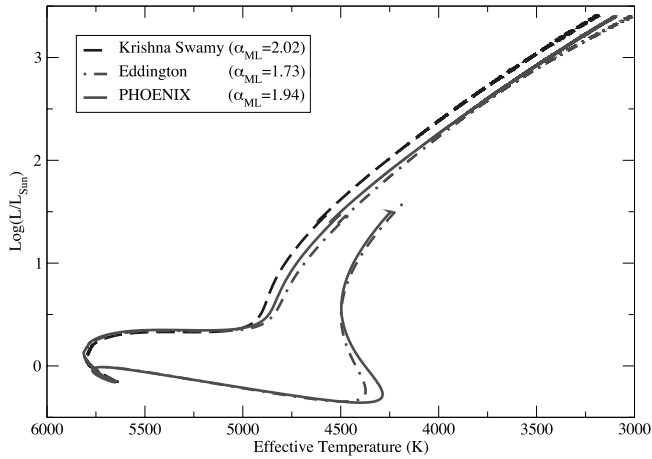


FIG. 1.— Evolution of calibrated solar models from the pre-main sequence to the tip of the red giant branch in the theoretical H-R diagram (except for Krishna Swamy, where no pre-main sequence is plotted). Each model employs a different surface boundary condition and requires a different solar-calibrated mixing length. The PHOENIX model atmosphere boundary condition adopted in this paper lies between the Eddington and Krishna Swamy models, both in terms of the mixing length and, as a result, the temperature of the red giant branch. Note also that while the Eddington and Krishna Swamy models have a similar amount of curvature above $\log L/L_{\odot} = 1$, the PHOENIX model red giant branch is relatively straight. [See the electronic edition of the Journal for a color version of this figure.]

Eddington model required the smallest mixing length ($\alpha_{ML} = 1.73$) and had the coolest red giant branch. The PHOENIX model was located between the others ($\alpha_{ML} = 1.94$), with a slightly flatter red giant branch in the H-R diagram (see Fig. 1).

2.6. Nuclear Reactions and Neutrino Cooling Rates

Nuclear reaction rates were adopted from Adelberger et al. (1998), with two exceptions: the $^{14}\text{N}(p, \gamma)^{15}\text{O}$ reaction (Imbriani et al. 2004) and the $^{12}\text{C}(\alpha, \gamma)^{16}\text{O}$ reaction (Kunz et al. 2002). Neutrino cooling rates were adopted from Haft et al. (1994).

2.7. He-Burning Evolution

Stellar evolution calculations that follow the He flash in low-mass stars are possible but numerically difficult and time-consuming. Furthermore, studies have shown that so-called

zero-age HB (pseudo-ZAHB) models that are concocted from red giant branch tip models are nearly indistinguishable from full He flash calculations in terms of the subsequent evolution (Vandenbergh et al. 2000; Piersanti et al. 2004; Serenelli & Weiss 2005). Beginning with a number of these pseudo-ZAHB models, a complete set of core and shell He-burning tracks were created based on the red giant branch tip surface compositions listed in Table 4 and He core masses listed in Table 5. Since He core mass varies slightly with age or stellar mass, the core mass was taken from the stellar track that reached the red giant branch tip closest to 12 Gyr; in each case, if no track for a given composition reached the tip within 11–13 Gyr, the surface composition and He core mass were interpolated between the two nearest tracks. The core and shell He-burning tracks allowed for the population of the HB and AGB through the use of individual tracks or the construction of synthetic HB models (for more on this, see § 4.4). The remainder of this section describes the physics in DSEP specific to core and shell He-burning processes.

Semiconvection has long been known to play an important role in stars with steep composition gradients such as occur in stars undergoing core He fusion. Instead of including a full treatment of a semiconvection zone, DSEP uses a simple model developed by Castellani et al. (1971) to treat the expansion of the core boundary during He-burning. Thermal energies have been artificially reduced when the core He mass fraction falls below 0.1 to discourage small loops in the H-R diagram, known as “breathing pulses,” as suggested by A. Sweigart (2004, private communication) and first noted by Dorman & Rood (1993). Breathing pulses occur when the convective core expands outward into He-rich material and are widely believed to be unphysical artifacts of the models. Neglecting thermal energies in the core during the final 10% of core He fusion effectively freezes the size of the core, thereby removing the breathing pulses. High-temperature opacities that allow for varying levels of C and O (as described in § 2.4) allow for correct opacity coverage of stellar cores during the entire core and shell He-fusion phases.

Time steps during core He fusion are based on the depletion of He in the core such that the mass fraction Y_{core} cannot fall by more than a fixed amount during each time step. When Y_{core} falls to 0.1, the time step is set by the percentage change in Y_{core} rather than by a fixed amount. During shell He fusion (the AGB phase) time steps are further constrained by the total He-burning luminosity: the sum of the triple- α and $^{12}\text{C}(\alpha, \gamma)^{16}\text{O}$ luminosities.

TABLE 4
SURFACE COMPOSITION FOR ZERO-AGE HORIZONTAL-BRANCH MODELS

[Fe/H]	$Y_{\text{init}} = 0.245 + 1.54Z_{\text{init}}^a$						$Y_{\text{init}} = 0.33$		$Y_{\text{init}} = 0.4$	
	$[\alpha/\text{Fe}] = -0.2$	0	0.2	0.4	0.6	0.8	0	0.4	0	0.4
−2.5	0.7564	0.7565	0.7561	0.7549	0.7549	0.7546	0.6835	0.6828	0.5967	0.6212
	3.97E−5	5.31E−5	7.40E−4	1.60E−4	1.08E−4	2.44E−4	4.63E−5	9.40E−5	8.40E−3	8.33E−5
−2.0	0.7547	0.7548	0.7551	0.7541	0.7524	0.7510	0.6820	0.6810	0.6208	0.6193
	1.26E−4	1.67E−4	2.33E−4	3.41E−4	5.06E−4	7.69E−4	1.46E−4	2.97E−4	1.30E−4	2.64E−4
−1.5	0.7529	0.7517	0.7509	0.7491	0.7464	0.7427	0.6785	0.6763	0.6180	0.6156
	3.96E−4	5.29E−4	7.37E−4	1.07E−3	1.60E−3	2.43E−3	4.64E−4	9.44E−4	4.12E−4	8.37E−4
−1.0	0.7472	0.7453	0.7423	0.7383	0.7325	0.7238	0.6732	0.6698	0.6125	0.6076
	1.25E−3	1.67E−3	2.33E−3	3.37E−3	5.02E−3	7.59E−3	1.47E−3	3.00E−3	1.31E−3	2.67E−3
−0.5	0.7350	0.7309	0.7242	0.71485	0.7010	0.6810	0.6641	0.6600	0.6034	0.5967
	3.94E−3	5.24E−3	7.27E−3	1.05E−2	1.54E−3	2.28E−2	4.67E−3	9.48E−3	4.16E−3	8.41E−3
0.0	0.7166	0.6928	0.6741	0.6497	0.6131	0.5640	0.6477	0.6303	0.5877	0.5854
	1.41E−2	1.85E−2	2.50E−2	3.46E−2	4.88E−2	6.87E−2	1.47E−2	2.91E−2	1.31E−2	2.68E−2

NOTE.—For each value of [Fe/H], the first row gives the X-value and the second row gives the Z-value.

^a Y_{init} and Z_{init} refer to the initial Y and Z values; for the initial compositions, see Table 1. The Z values in this table differ from those in Table 1 as a result of diffusion.

TABLE 5
He CORE MASS IN M_{\odot} FOR ZERO-AGE HORIZONTAL-BRANCH MODELS

[Fe/H]	$Y = 0.245 + 1.54Z$						$Y = 0.33$		$Y = 0.4$	
	$[\alpha/\text{Fe}] = -0.2$	0	0.2	0.4	0.6	0.8	0	0.4	0	0.4
-2.5	0.506	0.505	0.504	0.502	0.500	0.497	0.489	0.485	0.474	0.473
-2.0	0.500	0.499	0.498	0.494	0.492	0.491	0.483	0.479	0.469	0.468
-1.5	0.495	0.493	0.491	0.489	0.488	0.486	0.478	0.474	0.464	0.463
-1.0	0.489	0.488	0.486	0.484	0.481	0.478	0.473	0.469	0.460	0.458
-0.5	0.483	0.482	0.480	0.477	0.473	0.469	0.469	0.465	0.457	0.455
0.0	0.475	0.473	0.469	0.464	0.458	0.447	0.465	0.460	0.454	0.451

These time steps are used in addition to the time steps that govern the H shell.

3. COLOR- T_{eff} TRANSFORMATIONS

The stellar evolution models described in § 2 have been transformed to the observational plane by two different methods. One is based on the semiempirical colors and bolometric corrections of Vandenberg & Clem (2003) and the transformation equations found in Appendix D of Sirianni et al. (2005), specifically those in Tables 22 and 24, hereafter referred to as the “empirical” transformation. The other is based entirely on synthetic fluxes from PHOENIX (the same model atmosphere grids described in § 2.5) with the appropriate filter transmission curves and is hereafter referred to as the “synthetic” transformation.

Both transformations have been normalized to Vega and include ground-based B , V , and I magnitudes on the Bessell (1990) system and the *HST* filters F606W and F814W on both the WFPC2 and the ACS WFC systems from Sirianni et al. (2005).⁶

3.1. Empirical Transformation

Vandenberg & Clem (2003) introduced a new set of semiempirical color transformations and bolometric corrections for standard B , V , R , and I bandpasses. The colors and bolometric corrections were empirically constrained so that the isochrones of Bergbusch & Vandenberg (2001) were able to accurately reproduce the color-magnitude diagrams (CMDs) of various Galactic globular, open clusters and assorted other stars spanning the range of metallicities commonly observed in the Milky Way. The bolometric correction to V , $B - V$, and $V - I$ of Vandenberg & Clem (2003) make up the ground-based portion of the empirical transformation.

As part of an exhaustive study of the ACS, Sirianni et al. (2005) derived equations to convert between ground-based BVR magnitudes and *HST* ACS WFC (and HRC) filters, as well as between ACS and WFPC2 magnitudes. The V and I magnitudes from Vandenberg & Clem (2003) were used to obtain ACS F606W and F814W magnitudes using the $(V, V - I) \rightarrow \text{F606W}$ and $(I, V - I) \rightarrow \text{F814W}$ equations. The ACS magnitudes were then used to obtain the corresponding WFPC2 magnitudes. The Vegamag zero points were used in each case.

3.2. Synthetic Transformation

Filter transmission curves in B , V , and I from Bessell (1990) were used along with curves from both ACS WFC and WFPC2 filters F606W and F814W from Sirianni et al. (2005) to create a synthetic color- T_{eff} transformation. The synthetic colors were generated with fluxes from the same PHOENIX model atmospheres used for the surface boundary conditions described in

§ 2.5. These colors account for the effects of α -enhancement—something not possible with the empirical transformation. The effects are small but noticeable, particularly below ~ 4000 K.

As mentioned in § 2.5, model atmospheres with enhanced He were produced to properly treat He-enhancement in the surface boundary condition. Tests have been performed to address the possible changes to synthetic colors due to the enhancement of He in the atmosphere for the $Y = 0.33$ and 0.40 isochrones. As discussed in § 2.5, the flux is virtually unchanged by enhancing He, and this carries through to the colors. Therefore, He-enhancement has been omitted from the synthetic color transformation.

The procedure used to derive the synthetic transformation closely follows that of Bedin et al. (2005). The bolometric correction in V and all colors were normalized by assuming that Vega has $V = 0.026$ and all colors equal to zero.

3.3. Discussion

Isochrones constructed from the two color transformations tend to agree within ~ 0.05 mag in all filters (the only noteworthy exception occurs at high Z , where the synthetic B is flawed). Figure 2 compares the two color transformations over the full range of the isochrones for the case of $[\text{Fe}/\text{H}] = -2.0$, $[\alpha/\text{Fe}] = +0.4$ at 12 Gyr (*left*) and $[\text{Fe}/\text{H}] = 0$, $[\alpha/\text{Fe}] = 0.2$ at 5 Gyr (*right*). In each panel the transformed isochrones were plotted with no adjustments in either direction; thus, any difference is due entirely to the color transformation. The empirical isochrones (*dashed lines*) generally have a narrower color difference between the turnoff and the base of the red giant branch than the synthetic

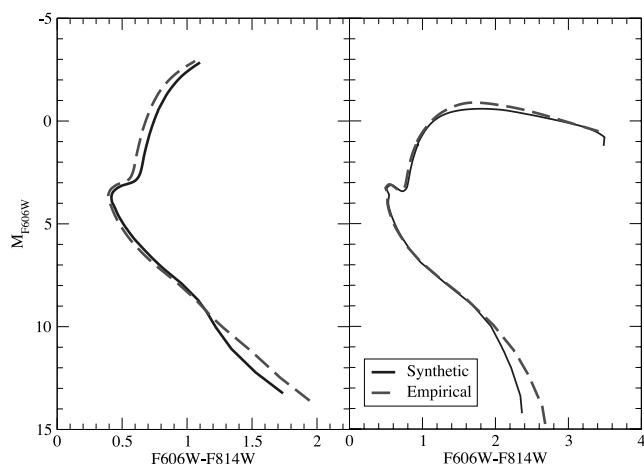


FIG. 2.— Color comparison for the ACS bands. *Left*: Isochrones with $[\text{Fe}/\text{H}] = -2.0$, $[\alpha/\text{Fe}] = +0.4$ at 12 Gyr. *Right*: Isochrones with $[\text{Fe}/\text{H}] = 0$, $[\alpha/\text{Fe}] = +0.2$ at 5 Gyr. The general trend is for the empirical colors (*dashed lines*) to be bluer on the giant branch and redder on the lower main sequence than the synthetic colors (*solid lines*). [See the electronic edition of the *Journal* for a color version of this figure.]

⁶ Readers interested in obtaining isochrones in other filters should contact the authors.

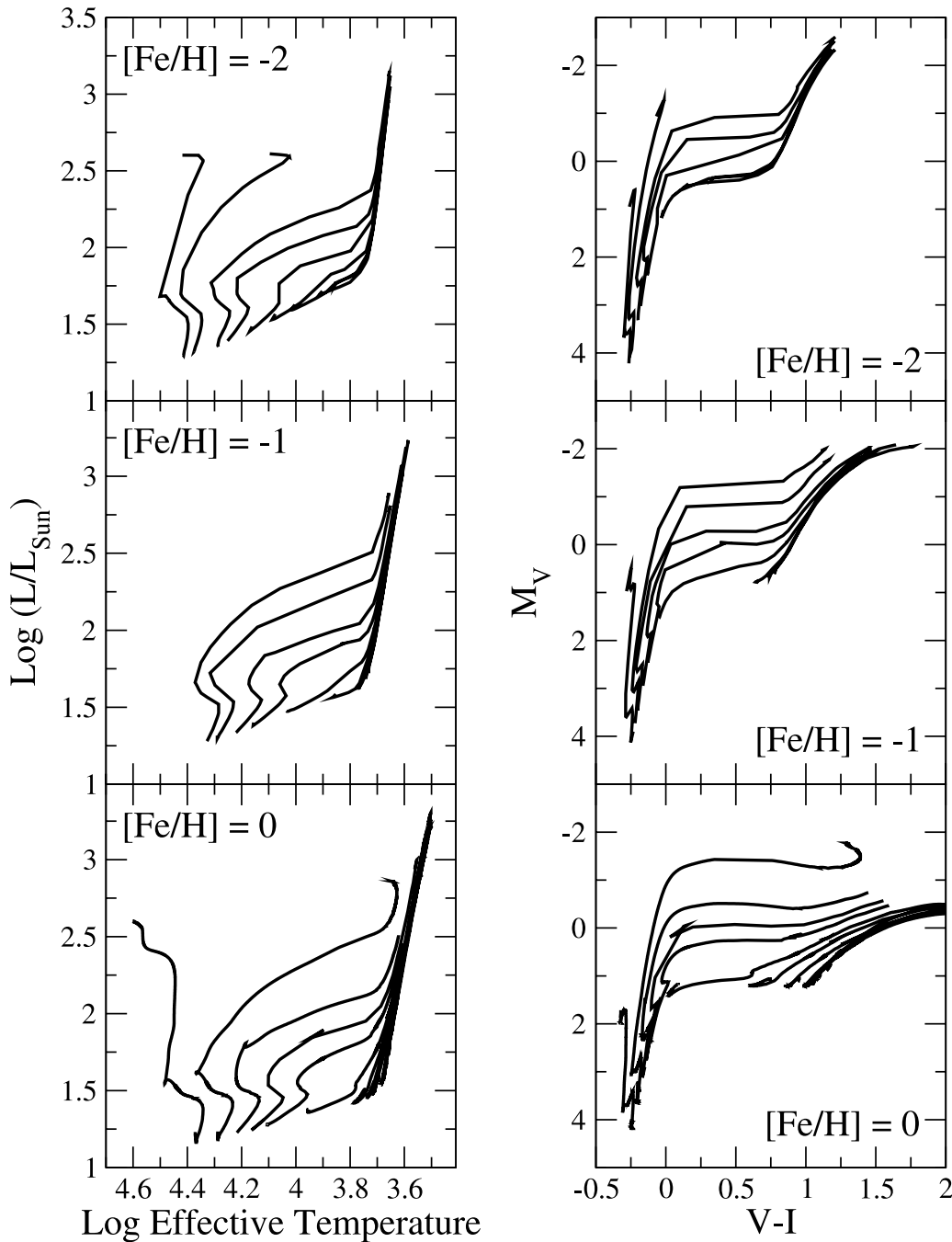


FIG. 3.—He-burning tracks for $[\text{Fe}/\text{H}] = -2$ (top panels), -1 (middle panels), and 0 (bottom panels), all with scaled-solar abundances. All masses from the minimum on the HB (see Table 5 for He core masses) of $0.51 M_{\odot}$ at $[\text{Fe}/\text{H}] = -2$, $0.5 M_{\odot}$ at $[\text{Fe}/\text{H}] = -1$, and $0.48 M_{\odot}$ at $[\text{Fe}/\text{H}] = 0$ up to $0.9 M_{\odot}$ are shown. The left panels show $\log T_{\text{eff}}$ vs. $\log L/L_{\odot}$, and the right panels show the V vs. $V-I$ CMD using empirical color transformation. The $[\text{Fe}/\text{H}] = 0$ CMD has been truncated in order to show the hotter stars more clearly; the tracks extend out to $V-I \sim 4$.

isochrones (solid lines). The red giant branches move closer together until they overlap (or nearly so) at the tip. There is good agreement for the first ~ 5 mag below the turnoff, and then the empirical transformation moves to the red of the synthetic.

Further side-by-side comparisons of the two color transformations for a variety of metallicities and ages, including comparisons to photometry with distance and reddening estimates, can be found in Sarajedini et al. (2007).

4. STELLAR EVOLUTION MODELS

This section lists and describes the different types of models accompanying this paper, either as data files or files that may be

generated with one of several computer programs distributed along with the data files. The contents of each type of data file and the use of the accompanying programs are described in the Appendix.

4.1. Stellar Evolution Tracks

Stellar evolution tracks of individual masses for each composition are available in two formats: one listing the theoretical quantities and one listing basic global parameters and absolute magnitudes in each of the seven filters listed in the previous section. The tracks include evolution from the pre-main sequence to the tip of the red giant branch in the range $0.1 M_{\odot} < M < 1.8 M_{\odot}$ in steps of $0.05 M_{\odot}$.

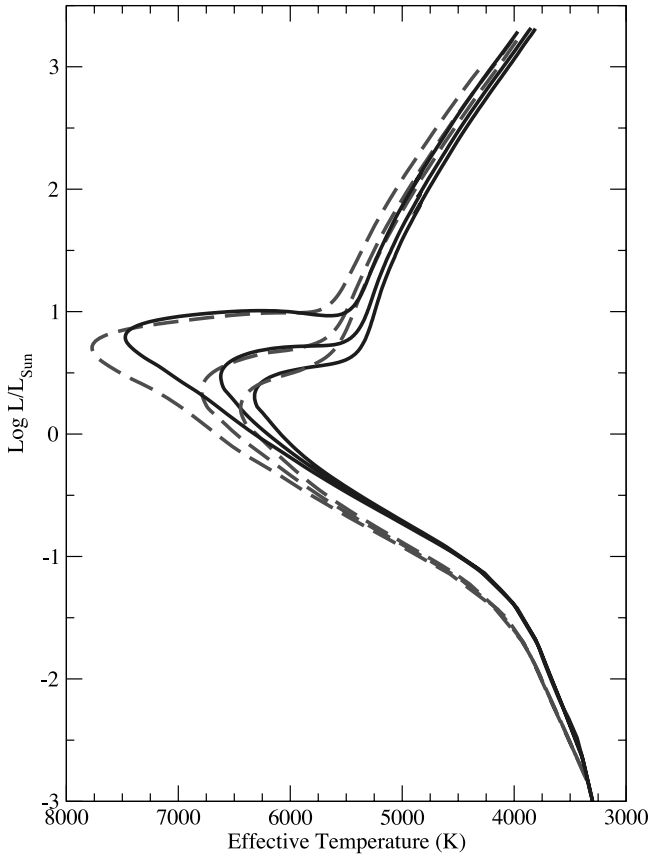


FIG. 4.—Effect of He-enhancement at constant $[\text{Fe}/\text{H}] = -1$ and $[\alpha/\text{Fe}] = 0$ in the H-R diagram. Ages shown are 4, 8, and 12 Gyr. The solid lines have $Y_{\text{init}} = 0.248$, while the dashed lines have $Y_{\text{init}} = 0.4$. [See the electronic edition of the *Journal* for a color version of this figure.]

Separate tracks follow the He-burning evolution from the ZAHB to the onset of the thermal pulsation (TP-AGB) phase; these tracks are produced at masses from $\sim 0.01 M_{\odot}$ above the He core mass (see Table 5) up to $0.51 M_{\odot}$ in steps of $0.01 M_{\odot}$, from 0.52 to $0.56 M_{\odot}$ in steps of $0.02 M_{\odot}$, from 0.6 to $0.9 M_{\odot}$ in steps of $0.1 M_{\odot}$, and, finally, from 1 to $1.75 M_{\odot}$ in steps of $0.25 M_{\odot}$. Figure 3 shows all of the He-burning tracks from the minimum mass up to $0.9 M_{\odot}$ for $[\text{Fe}/\text{H}] = -2, -1$, and 0 with scaled-solar abundances in both the theoretical H-R diagram and the V versus $V - I$ CMD. The $[\text{Fe}/\text{H}] = 0$ tracks have been truncated in the CMD in Figure 3 in order to more clearly show the hotter stars; these tracks extend out to $V - I \sim 4$.

Each track file containing theoretical quantities lists the following for each time step: age in years, log effective temperature in kelvins, log surface gravity in cgs , log luminosity in L_{\odot} , log radius in R_{\odot} , core He mass fraction Y_{core} , core heavy-element mass fraction Z_{core} , surface heavy element-to-hydrogen ratio $(Z/X)_{\text{surf}}$, H luminosity in L_{\odot} , He luminosity in L_{\odot} , He core mass in M_{\odot} , and CO core mass in M_{\odot} . All logarithms are base 10; see Table 2 for the adopted solar values.

Each track file containing observable quantities lists at each time step the first four quantities from the theoretical track file (age, $\log T_{\text{eff}}$, $\log g$, and $\log L/L_{\odot}$) plus absolute magnitudes in each of the seven filters listed above. Samples of both types of track file are given in the Appendix.

4.2. Isochrones

Isochrones were built using the equivalent evolutionary phase (EEP) principle. Each isochrone consists of ~ 230 points extend-

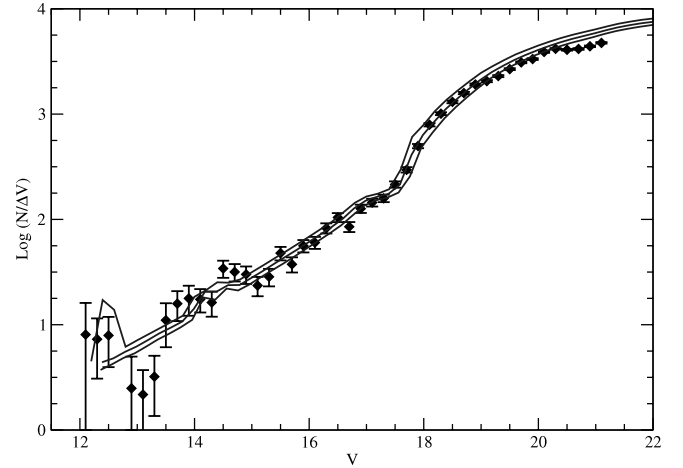


FIG. 5.—Example of theoretical luminosity functions compared to the observed luminosity function of M92 (Paust et al. 2007). The theoretical luminosity functions shown were derived from synthetic color isochrones at $[\text{Fe}/\text{H}] = -2.3$ and $[\alpha/\text{Fe}] = 0.4$ for ages 12, 13, and 14 Gyr with a Salpeter IMF (see also Fig. 11). The assumed distance is $DM_V = 14.6$. The full theoretical luminosity functions extend several magnitudes fainter than the data shown in the plot. [See the electronic edition of the *Journal* for a color version of this figure.]

ing from the lowest mass point at $0.1 M_{\odot}$ up to the tip of the red giant branch. EEPs are assigned as a function of the core He mass fraction (Y_{core}) on the main sequence and as a function of the He core mass on the red giant branch. Isochrones are provided for ages ranging from 2 to 15 Gyr in increments of 0.5 Gyr for all compositions.

Figure 4 demonstrates the effects of He-enhancement at $[\text{Fe}/\text{H}] = -1$, $[\alpha/\text{Fe}] = 0$. Isochrones are shown for 4, 8, and 12 Gyr with $Y_{\text{init}} \sim 0.25$ (solid lines) and $Y_{\text{init}} = 0.4$ (dashed lines). Enhancing He at the expense of H decreases the opacity and increases the mean molecular weight. An isochrone with $Y = 0.40$ is significantly hotter at a given age and $[\text{Fe}/\text{H}]$ than its counterpart with $Y \sim 0.25$, as demonstrated in Figure 4.

The model grid has a relatively fine spacing in both age and $[\alpha/\text{Fe}]$, such that parameters of clusters being analyzed can be well constrained by the ages and $[\alpha/\text{Fe}]$ values provided. The only parameter that could be more tightly constrained by observations is $[\text{Fe}/\text{H}]$, and so a computer program has been distributed with the isochrones to allow interpolation in $[\text{Fe}/\text{H}]$ (see the Appendix for details). The interpolation code has been extensively tested and provides results that are reasonably accurate over the range of metallicities provided in the isochrone database.

4.3. Luminosity Functions

Rather than adopt one or more standard filters and initial mass functions (IMFs) and provide tabulated luminosity functions, the distribution of isochrones comes with a program to generate luminosity functions from any of the isochrone files. Luminosity functions may be constructed from any bandpass with a choice of bin size and IMF exponent. The IMF is assumed to be a simple power law, where $dN/dM \propto M^x$. Figure 5 demonstrates the luminosity functions when compared to the observed luminosity function of M92 (Paust et al. 2007). The isochrones on which the luminosity functions were based have $[\text{Fe}/\text{H}] = -2.3$ and $[\alpha/\text{Fe}] = 0.4$ at ages 12, 13, and 14 Gyr. The assumed IMF is the Salpeter IMF ($x = -2.35$), a good fit in this case. The distance modulus is $DM_V = 14.6$, as used in Figure 11 (discussed in § 5.2) for the synthetic isochrone. The theoretical luminosity

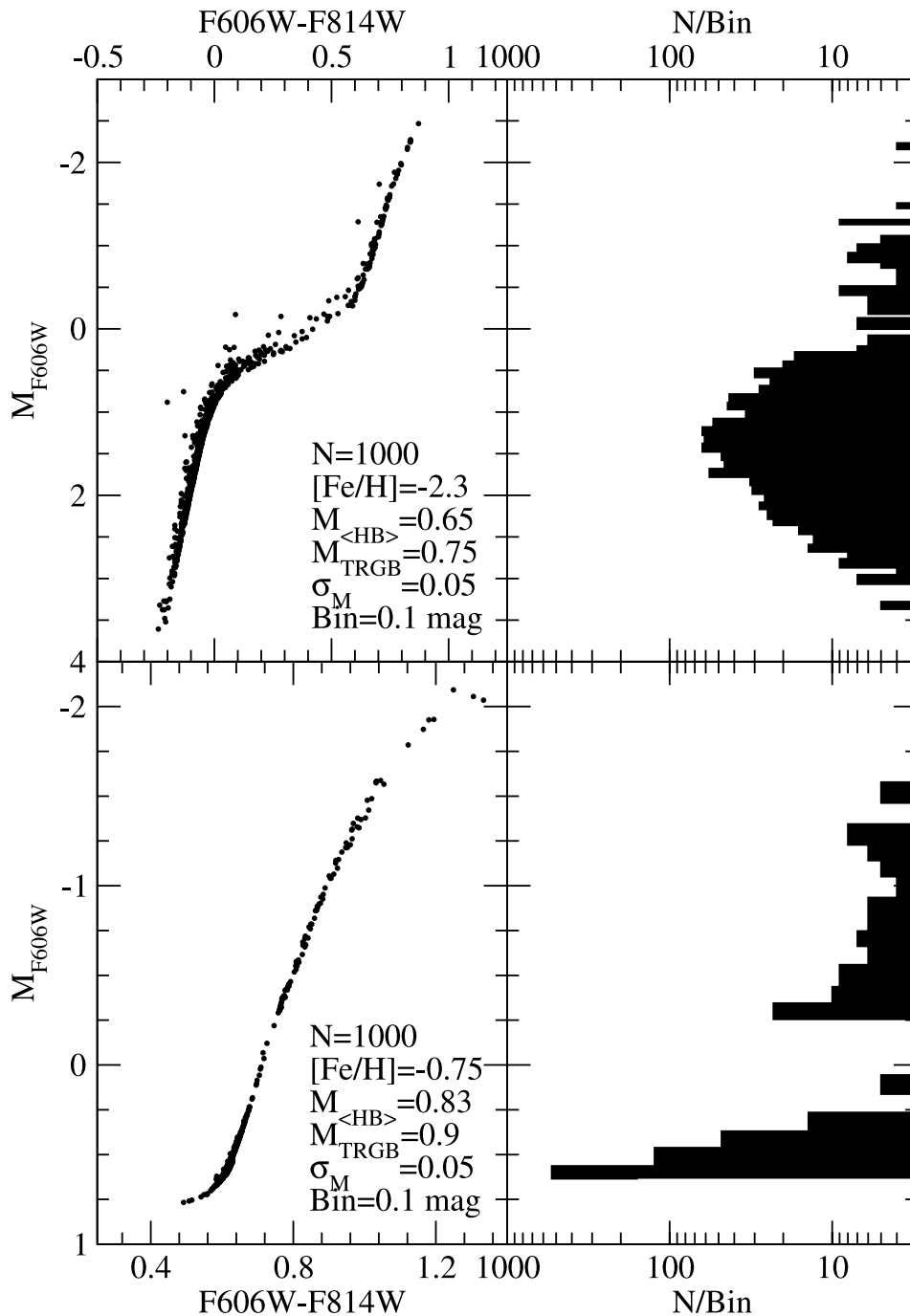


FIG. 6.— SHB models with $N = 1000$ stars for $[\text{Fe}/\text{H}] = -2.3$ (top panels) and -0.75 (bottom panels). The top panels represent a blue HB in a metal-poor cluster, while the bottom panels represent a red HB in a metal-rich cluster. The left panels show the SHB models in the CMD, while the right panels show the corresponding luminosity function.

functions extend approximately 7 mag fainter than the data in that figure.

4.4. Synthetic HB Models

The formalism for constructing synthetic HB (SHB) models is based on that of Lee et al. (1990). Accompanying the He-burning evolutionary tracks is a program for creating SHB models. The SHB models are constructed from the tracks based on a set of parameters supplied by the user. These parameters include the composition ($[\text{Fe}/\text{H}]$, $[\alpha/\text{Fe}]$, Y), average HB mass ($\langle M_{\text{HB}} \rangle$), the mass of the red giant branch tip (M_{RGT} , effectively the upper bound to the mass distribution), the mass dispersion on the HB (σ_M), and

the total number of stars (N). Based on these parameters, the program constructs a random Gaussian distribution of stellar masses, chooses a random age, and interpolates within the tracks for mass, age, and composition. Interpolation is linear within the tracks (age) and quadratic between the tracks (mass and composition). Because the actual mass distribution is finite and truncated on both ends by the finite mass range, the code calculates the actual values of $\langle M_{\text{HB}} \rangle$ and σ_{HB} and reports these at the end.

In addition to the SHB-generating program, another program is supplied to generate luminosity functions and fiducial lines from each individual SHB model. The program works on the raw output from the SHB code by binning the SHB model data,

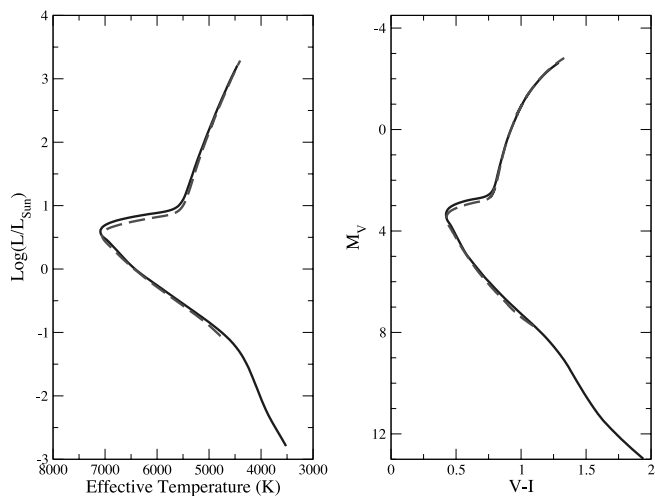


FIG. 7.—BaSTI (*dashed lines*) and DSEP (*solid lines*, with synthetic color transformation) isochrones at 10 Gyr with $Z = 0.0001$ (scaled-solar mixture). [See the electronic edition of the Journal for a color version of this figure.]

determining the mean and standard deviation of magnitude and color in a given bin, and writing results from all bins to a file. Use of this program in conjunction with the isochrone-based luminosity function code allows the user to create full luminosity functions that include all phases of the evolution covered by the models presented in this paper.

Figure 6 (*left panels*) demonstrates the effect of metallicity and the mass variables mentioned above in the CMD for two SHB models. Figure 6 (*right panels*) shows luminosity functions constructed from the corresponding SHB model. The top panels show the CMD and number distribution of a blue, metal-poor SHB model with $[\text{Fe}/\text{H}] = -2.3$. The bottom panels show the same information for a red, metal-rich SHB model with $[\text{Fe}/\text{H}] = -0.75$. These models, shown alone here, are the same as those used in Figures 11 and 12, discussed below.

5. COMPARISONS

5.1. Comparisons with Other Isochrones

Several groups currently produce stellar evolution tracks and isochrones similar to those presented in this paper. Here the new DSEP isochrones are compared to isochrones from the other groups in both the theoretical H-R diagram and the V versus

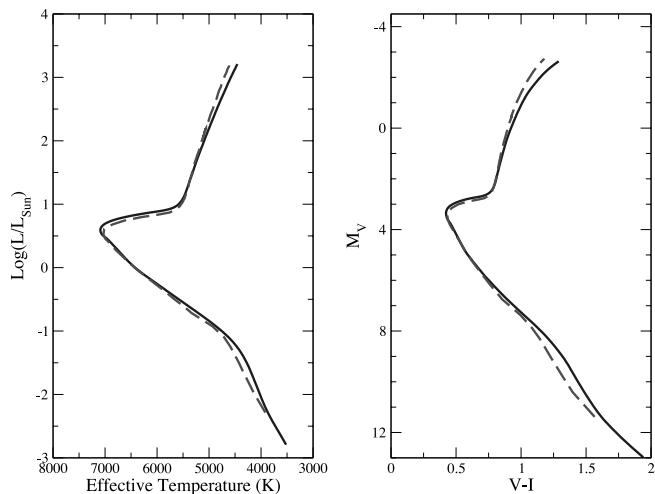


FIG. 8.—Same as Fig. 7, but for Padova and DSEP isochrones. [See the electronic edition of the Journal for a color version of this figure.]

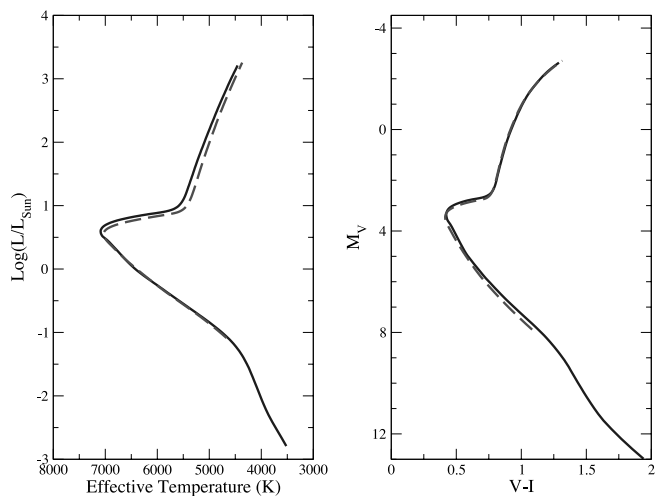


FIG. 9.—Same as Fig. 7, but for Victoria-Regina and DSEP isochrones. [See the electronic edition of the Journal for a color version of this figure.]

$V - I$ CMD. Comparisons in the H-R diagram allow for differences in the luminosity and effective temperature scales to be assessed, while comparisons in the CMD display additional differences due to the color- T_{eff} transformations adopted by each group. The DSEP isochrones shown in the following figures employ the synthetic color transformation. Figure 2 shows the differences between synthetic and empirical colors.

Comparisons were performed between DSEP and BaSTI, Padova, Victoria-Regina, and Y^2 at 10 Gyr with $Z = 0.0001$ assuming a scaled-solar heavy-element mixture. Scaled solar was chosen because the definition and amount of α -enhancement is not consistent among the groups. Neither the primordial He abundance nor the adopted $\Delta Y/\Delta Z$ is consistent at $Z = 0.0001$: BaSTI has $Y = 0.245$, DSEP has $Y = 0.02452$, Padova has $Y = 0.23$, Victoria-Regina has $Y = 0.2352$, and Y^2 has $Y = 0.2302$. Aside from differences in composition, the groups do not all use the same assumptions concerning the surface boundary condition, solar-calibrated mixing length, EOS, gravitational settling, mass loss, and other physics; each of these discrepancies will result in differences among the isochrones at constant age and Z .

DSEP is compared to BaSTI in Figure 7, Padova in Figure 8, Victoria-Regina in Figure 9, and Y^2 in Figure 10. Both BaSTI and

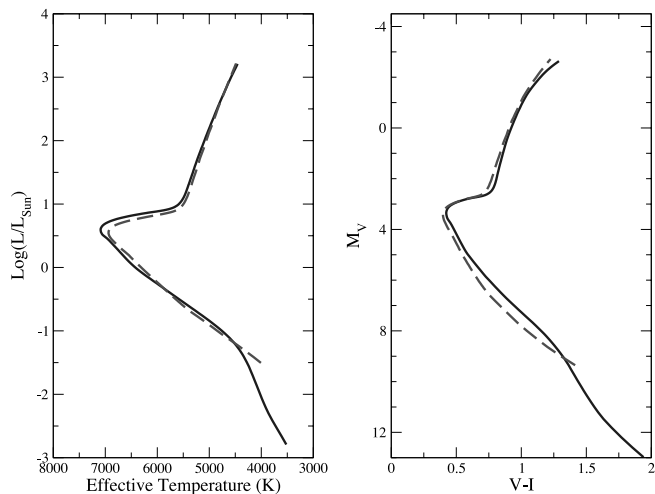


FIG. 10.—Same as Fig. 7, but for Y^2 and DSEP isochrones. [See the electronic edition of the Journal for a color version of this figure.]

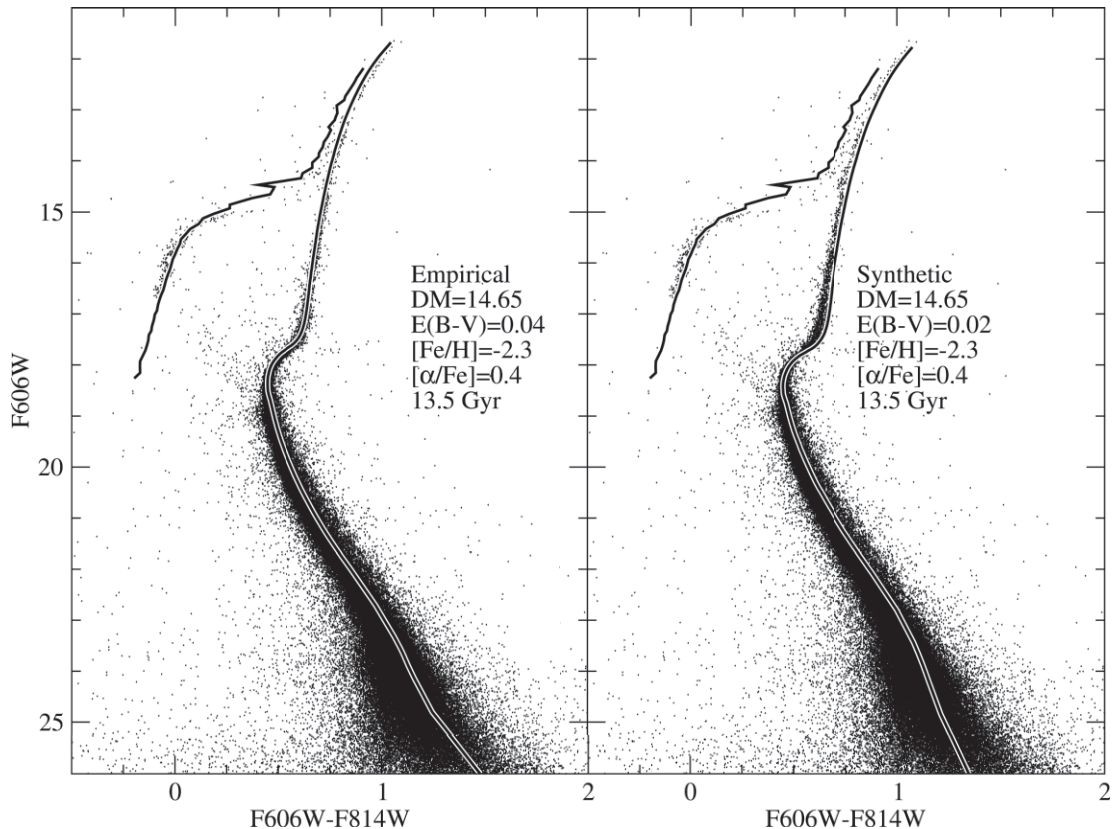


FIG. 11.— ACS data from M92 compared to isochrones with both empirical (*left*) and synthetic (*right*) color transformations. Details are listed on each panel. Data are from Sarajedini et al. (2007). The fiducial line from the metal-poor SHB model of § 4.4 (Fig. 6) is plotted alongside both isochrones. [See the electronic edition of the *Journal* for a color version of this figure.]

Padova include core He-burning evolution in their isochrones, but these portions have been removed from the figures for clarity.

In the H-R diagram, the DSEP isochrone is slightly hotter and more luminous than the others around the main-sequence turnoff and subgiant branch, but this does not carry over to the CMD in all cases due to the differences in adopted color- T_{eff} transformations. There is good agreement on the red giant branch tip location in the H-R diagram, except with Padova, which is hotter than all the others.

5.2. Comparisons with Data

ACS photometry of two classic globular clusters, M92 and 47 Tuc, from Sarajedini et al. (2007) and J. Anderson et al. (2007, in preparation) that span a wide range of $[\text{Fe}/\text{H}]$ have been chosen for comparison. Many of the brightest giants in Figures 11 and 12 include at least one saturated pixel. For these stars, the point-spread function photometry used for unsaturated stars has been replaced by aperture photometry with excellent results. Full details of the photometry will be given by J. Anderson et al. (2007, in preparation).

Isochrones for reasonable assumptions of the age and metallicity of M92 and 47 Tuc are shown in Figures 11 and 12, respectively. The purpose of these plots is not to constrain any of these parameters for either cluster but merely to demonstrate the performance of the isochrones when compared with high-quality photometry. To show the complete evolutionary picture, ridges from the SHB models shown in Figure 6 are plotted alongside the photometry, the metal-poor SHB model ($[\text{Fe}/\text{H}] = -2.3$) with M92 and the metal-rich SHB model ($[\text{Fe}/\text{H}] = -0.75$) with 47 Tuc, assuming the adopted reddening and distance modulus in each figure.

Figure 11 shows little difference between the color transformations, and both provide reasonable fits over the extent of the data. The largest difference occurs on the subgiant branch: in this region the empirical color transformation is clearly a better fit to the data.

Figure 12 shows more clearly the difference between the two color transformations. The empirical transformation provides a good fit from the red giant branch to approximately 3 mag below the main-sequence turnoff. Below that point, the empirical isochrone becomes increasingly redder than the data. In contrast, the synthetic isochrone becomes bluer than many of the data a few magnitudes below the turnoff. The difference between the synthetic isochrone and the data is smaller than in the case of the empirical isochrone, and in the opposite sense, but still visible. For this reason, the synthetic color transformation is preferable for examining the lower main sequence, particularly at higher metallicity. Disagreement between the lower main sequence data and the isochrones is likely due to errors in both the effective temperature scale of the stellar evolution models and the color transformations, but it is not possible to quantify these errors.

6. CONCLUSIONS

The stellar evolution code DSEP has been modified to significantly improve the accuracy of isochrones on the lower main sequence and to self-consistently treat the outer layers of stars for a wide range of compositions. The synthetic and empirical color transformations employed have strengths and weaknesses and highlight the amount of uncertainty in current color- T_{eff} transformations in the CMD.

A new set of stellar evolution models has been computed with DSEP for a range of initial He abundance, $[\text{Fe}/\text{H}]$, and $[\alpha/\text{Fe}]$

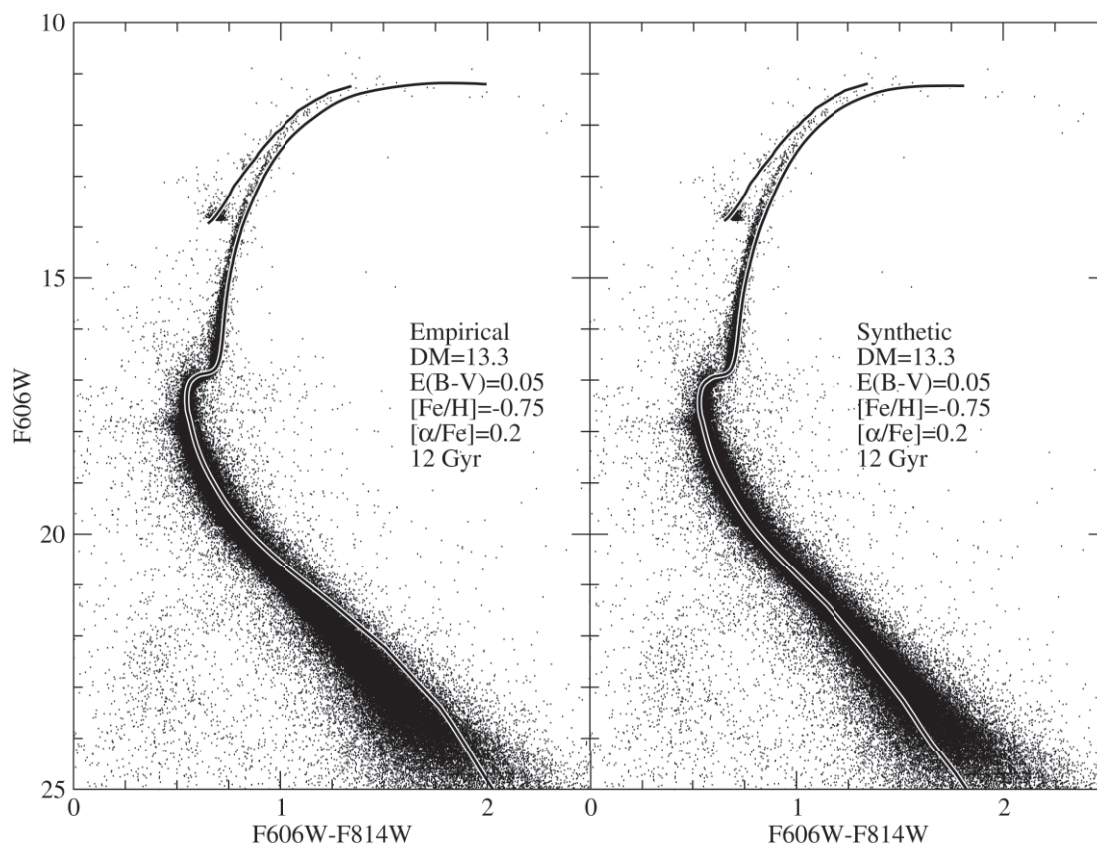


FIG. 12.— ACS data from 47 Tuc compared to isochrones with both empirical (*left*) and synthetic (*right*) color transformations. Details are listed on each panel. Data are from Sarajedini et al. (2007). The fiducial line from the metal-rich SHB model of § 4.4 (Fig. 6) is plotted alongside both isochrones. [See the electronic edition of the *Journal* for a color version of this figure.]

encompassing ages and abundances observed in Galactic globular clusters. The models are presented as isochrones (2–15 Gyr) and stellar evolution tracks (0.1 – $1.8 M_{\odot}$) with accompanying programs to create luminosity functions and SHB models suitable for analysis of old stellar populations. Details of the stellar evolution calculations and all steps leading up to the final products have been discussed. These models should prove useful in analyzing the ACS Survey of Galactic Globular Clusters data and in many other applications. The models extend to $0.1 M_{\odot}$ on the main sequence, making them useful for deep photometric studies such as the ACS Survey of Galactic Globular Clusters, and at the same time they provide nearly full evolutionary coverage of the stars found in globular clusters and old open clusters.

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APPENDIX

A. COMPOSITION, COLOR TRANSFORMATIONS, AND FILE NAMES

Composition is specified in the file names in terms of $[\text{Fe}/\text{H}]$ as *feh* and $[\alpha/\text{Fe}]$ as *afe*. The values of $[\text{Fe}/\text{H}]$ and $[\alpha/\text{Fe}]$ are listed using the letters *m* for minus (<0) and *p* for plus (>0) followed by two digits for $[\text{Fe}/\text{H}]$ and one digit for $[\alpha/\text{Fe}]$. For example, $[\text{Fe}/\text{H}] = -2.5$, $[\alpha/\text{Fe}] = +0.4$ would be represented by the string *feh*m25*afep*4. Zero is set to positive, and thus $[\text{Fe}/\text{H}] = 0$, $[\alpha/\text{Fe}] = 0$ is listed as *feh*p00*afep*0. Tracks and isochrones with enhanced He, either $Y = 0.33$ or $Y = 0.4$, have an additional string following *feh*XXX*afe*XX in their file names: *y*33 or *y*40, respectively.

Color transformations are listed as a suffix to the isochrone and evolutionary track file names. The synthetic color transformation is indicated by the suffix *.phx*, and the semiempirical color transformation is indicated by *.cmd*. As a final example, the synthetic color isochrone with $[\text{Fe}/\text{H}] = -1.5$, $[\alpha/\text{Fe}] = +0.4$, and $Y = 0.33$ is called *feh*m15*afep*4*y*33.*phx*.

B. STELLAR EVOLUTION TRACK FILES

The stellar evolution tracks begin on the pre-main-sequence and terminate at the He flash or begin on the ZAHB and terminate at the onset of thermal pulsations on the AGB (TP-AGB).

TABLE 6
EXAMPLE OF A STELLAR EVOLUTION TRACK WITH THEORETICAL QUANTITIES

$M = 1.00$, $X = 7.0710E - 01$, $Z = 1.8850E - 02$, COMP = GS98, $[\alpha/\text{Fe}] = 0.00$, MIXL = 1.9380, CCO = 0.00											
Age (yr)	$\log T$	$\log g$	$\log L$	$\log R$	Y_{core}	Z_{core}	$(Z/X)_{\text{surf}}$	L_{H}	L_{He}	$M_{\text{He core}}$	$M_{\text{CO core}}$
5.27000000E+02.....	3.62873	2.41053	1.49574	1.01367	2.7402E-01	1.8850E-02	2.6658E-02	0.0000E+00	0.0000E+00	0.0000E+00	0.0000E+00
5.94200000E+02.....	3.62878	2.41191	1.49456	1.01298	2.7402E-01	1.8850E-02	2.6658E-02	0.0000E+00	0.0000E+00	0.0000E+00	0.0000E+00
6.64760000E+02.....	3.62884	2.41335	1.49335	1.01225	2.7402E-01	1.8850E-02	2.6658E-02	0.0000E+00	0.0000E+00	0.0000E+00	0.0000E+00
7.38848000E+02.....	3.62890	2.41487	1.49209	1.01150	2.7402E-01	1.8850E-02	2.6658E-02	0.0000E+00	0.0000E+00	0.0000E+00	0.0000E+00
8.16640400E+02.....	3.62897	2.41645	1.49077	1.01071	2.7402E-01	1.8850E-02	2.6658E-02	0.0000E+00	0.0000E+00	0.0000E+00	0.0000E+00
8.98322420E+02.....	3.62904	2.41810	1.48940	1.00988	2.7402E-01	1.8850E-02	2.6658E-02	0.0000E+00	0.0000E+00	0.0000E+00	0.0000E+00
9.84088541E+02.....	3.62911	2.41983	1.48796	1.00902	2.7402E-01	1.8850E-02	2.6658E-02	0.0000E+00	0.0000E+00	0.0000E+00	0.0000E+00
1.07414297E+03.....	3.62919	2.42163	1.48646	1.00811	2.7402E-01	1.8850E-02	2.6658E-02	0.0000E+00	0.0000E+00	0.0000E+00	0.0000E+00

Tracks are provided for masses ranging from 0.1 to 1.8 M_{\odot} in increments of 0.05 M_{\odot} . For the mass range that experiences the He flash, separate He-burning tracks are provided. The content of these tracks is identical to that of the H-burning tracks as described below.

The track files come in two formats: one containing theoretical quantities and one containing observable quantities. The theoretical track files contain information about the initial composition and physics used in a header followed by data lines. Each data line lists age in years, $\log T_{\text{eff}}$, $\log g$, $\log L/L_{\odot}$, $\log R/R_{\odot}$, Y_{core} , Z_{core} , $(Z/X)_{\text{surf}}$, L_{H}/L_{\odot} , L_{He}/L_{\odot} , $M_{\text{He core}}$, and $M_{\text{CO core}}$. The quantity L_{H} is the total H-fusion luminosity combining all branches of the p - p chain and CNO cycle luminosities, while L_{He} combines luminosities from the triple- α and $^{12}\text{C}(\alpha, \gamma)^{16}\text{O}$ reactions.

The naming convention for the stellar evolution tracks is to use the prefix `m` (for mass) followed by three integers representing the stellar mass (1 $M_{\odot} = \text{m}100$) followed by composition information and either the suffix `.trk` for theoretical quantities or the appropriate suffix for a color transformation specified above. The He-burning tracks use the same naming convention, except the `m` has been replaced with `hb`.

The header and first eight data lines of the 1 M_{\odot} , $[\text{Fe}/\text{H}] = 0$, $[\alpha/\text{Fe}] = 0$, $Y = 0.27$ stellar evolution track are given for both the theoretical file `m100fehp00afep0.trk` in Table 6 and the observable file with the PHOENIX color transformation `m100fehp00afep0.phx` in Table 7. In Table 7, the columns labeled F606W and F814W refer to the ACS bands, and the columns labeled F6062 and F8142 refer to the WFPC2 bands.

C. ISOCHRONES AND LUMINOSITY FUNCTIONS

Isochrone files begin with a header that lists composition and color transformation details, as well as the number of ages in

the file. The isochrone files distributed with this paper contain 27 ages from 2 to 15 Gyr in increments of 0.5 Gyr. The individual isochrones have their own headers consisting of two lines. The first line gives the age of the isochrone and the number of data points (EEPs) it contains. The second line lists the quantities provided on each line: EEP number, stellar mass, $\log T_{\text{eff}}$, $\log g$, $\log L/L_{\odot}$, and the absolute magnitudes in all seven filters. The header and first six data lines of the $[\text{Fe}/\text{H}] = 0$, $[\alpha/\text{Fe}] = 0$ isochrone file `fehp00afep0.cmd` are given in Table 8. The columns labeled F606W and F814W refer to the ACS bands, and the columns labeled F6062 and F8142 refer to the WFPC2 bands.

Output from the luminosity function program (`isolf.f`) contains the same header as the isochrone file it was created from. Likewise, each luminosity function has a two-line header that lists the age and number of data points followed by a description of the information printed in each line. The data lines list a bin number, the average absolute magnitude of the bin, the average color of the bin, the cumulative (logged) number of stars, and the (logged) number of stars in that bin. The cumulative star count begins with the brightest bin. The header and first six data lines of a luminosity function generated from the isochrone file (Table 8) are printed in Table 9. The luminosity function is created using the F606W (ACS) filter with color F606W – F814W, bin size of 0.1 mag, and a Salpeter IMF.

D. SYNTHETIC HORIZONTAL-BRANCH MODELS AND LUMINOSITY FUNCTIONS

Output from the SHB code (`shb.f`) consists of a file containing one line for each star plus a short header. The header lists number of stars, $[\text{Fe}/\text{H}]$, $[\alpha/\text{Fe}]$, and Y_{init} . The data lines include stellar mass, age, $\log T_{\text{eff}}$, $\log g$, $\log L/L_{\odot}$, and absolute magnitudes in all seven filters. Table 10 shows the header and first eight

TABLE 7
EXAMPLE OF A STELLAR EVOLUTION TRACK WITH OBSERVABLE QUANTITIES

$M/M_{\odot} = 1.00$, $[\text{Fe}/\text{H}] = 0.07$, $[\alpha/\text{Fe}] = 0.00$, $Y = 0.274$										
Age (yr)	$\log T$	$\log g$	$\log L$	B	V	I	F606W	F814W	F6062	F8142
5.27000000E+02.....	3.62873	2.41053	1.49574	3.1495	1.7757	0.3959	1.4607	0.4112	1.3968	0.4286
5.94200000E+02.....	3.62878	2.41191	1.49456	3.1517	1.7781	0.3988	1.4632	0.4140	1.3994	0.4315
6.64760000E+02.....	3.62884	2.41335	1.49335	3.1539	1.7806	0.4017	1.4657	0.4169	1.4019	0.4344
7.38848000E+02.....	3.62890	2.41487	1.49209	3.1562	1.7832	0.4047	1.4684	0.4199	1.4046	0.4374
8.16640400E+02.....	3.62897	2.41645	1.49077	3.1585	1.7859	0.4079	1.4712	0.4231	1.4074	0.4405
8.98322420E+02.....	3.62904	2.41810	1.48940	3.1609	1.7886	0.4112	1.4740	0.4263	1.4102	0.4438
9.84088541E+02.....	3.62911	2.41983	1.48796	3.1635	1.7916	0.4146	1.4771	0.4298	1.4133	0.4472
1.07414297E+03.....	3.62919	2.42163	1.48646	3.1662	1.7946	0.4182	1.4802	0.4334	1.4164	0.4508

TABLE 8
EXAMPLE OF AN ISOCHRONE FILE

Number of Ages = 29											
MIX-LEN = 1.9380			Y = 0.2741	Z = 1.8850E − 02	Z _{eff} = 1.8850E − 02	[Fe/H] = 0.07	[α/Fe] = 0.00				
Photometric System: V&C2003 ^a + Sirianni et al. (2005)											
AGE = 2.000 EEPs = 238											
EEP	M/M _⊙	log T _{eff}	log G	log L/L _⊙	B	V	I	F606W	F814W	F6062	F8142
4.....	0.099162	3.5048	5.3118	−2.9048	16.5676	14.8000	11.7423	14.1827	11.6499	14.0308	11.6953
5.....	0.115037	3.5048	5.2473	−2.7758	16.2406	14.4760	11.4311	13.8600	11.3397	13.7098	11.3849
6.....	0.134056	3.5048	5.1810	−2.6432	15.9058	14.1439	11.1106	13.5292	11.0200	13.3805	11.0652
7.....	0.169755	3.5076	5.1016	−2.4501	15.3348	13.5914	10.6076	12.9822	10.5209	12.8399	10.5654
8.....	0.223896	3.5162	5.0395	−2.2333	14.5387	12.8438	9.9844	12.2493	9.9068	12.1222	9.9497
9.....	0.276573	3.5236	4.9959	−2.0684	13.9167	12.2621	9.5074	11.6808	9.4371	11.5657	9.4787

^a V&C2003 refers to Vandenberg & Clem (2003).

TABLE 9
EXAMPLE OF A LUMINOSITY FUNCTION FILE

Number of Ages = 29					
MIX-LEN = 1.9380	$Y = 0.2741$	$Z = 1.8850E - 02$	$Z_{\text{eff}} = 1.8850E - 02$	$[\text{Fe}/\text{H}] = 0.07$	$[\alpha/\text{Fe}] = 0.00$
Photometric System: V&C (2003) + Sirianni et al. (2005)					
AGE = 2.000 BINS = 158 IMF $x = -2.35$					
N	MF606W	F606W – F814W	$\log N$	$\log dN$	
1.....	–1.517	1.7545	0.1086	0.1086	...
2.....	–1.417	1.7064	0.3138	–0.1105	...
3.....	–1.317	1.4011	0.4090	–0.2966	...
4.....	–1.217	1.3380	0.4921	–0.2671	...
5.....	–1.117	1.2829	0.5655	–0.2426	...
6.....	–1.017	1.2381	0.6276	–0.2478	...

TABLE 10
EXAMPLE OF A SYNTHETIC HORIZONTAL-BRANCH FILE

$N = 1000$, $[\text{Fe}/\text{H}] = 0.00$, $[\alpha/\text{Fe}] = 0.00$, $Y_i = 0.274$											
M/M_{\odot}	Age (yr)	$\log T_{\text{eff}}$	$\log g$	$\log L/L_{\odot}$	B	V	I	F606W	F814W	F6062	F8142
0.7771.....	1.0639E+08	3.6677	2.2732	1.6791	2.1224	0.9687	–0.1202	0.7106	–0.1177	0.7032	–0.1000
0.7920.....	9.7970E+07	3.6727	2.3536	1.6274	2.1900	1.0672	0.0076	0.8166	0.0106	0.8100	0.0279
0.8825.....	6.9989E+07	3.6760	2.4404	1.6004	2.2174	1.1160	0.0741	0.8700	0.0774	0.8639	0.0944
0.7812.....	5.2351E+07	3.6792	2.4334	1.5671	2.2678	1.1813	0.1560	0.9397	0.1596	0.9339	0.1763
0.8525.....	1.0007E+08	3.6694	2.3350	1.6638	2.1377	0.9968	–0.0823	0.7412	–0.0796	0.7341	–0.0620
0.8959.....	1.1958E+08	3.6137	1.5453	2.2529	1.4486	–0.0369	–1.5517	–0.3977	–1.5599	–0.4197	–1.5359
0.7885.....	9.2760E+07	3.6747	2.3783	1.6087	2.2134	1.1021	0.0536	0.8544	0.0568	0.8480	0.0739
0.8229.....	1.0110E+08	3.6695	2.3239	1.6602	2.1459	1.0050	–0.0734	0.7495	–0.0707	0.7424	–0.0532

TABLE 11
EXAMPLE OF A SYNTHETIC HORIZONTAL-BRANCH LUMINOSITY FUNCTION FILE

$N = 1000$, $[\text{Fe}/\text{H}] = 0.00$, $[\alpha/\text{Fe}] = 0.00$, $Y_i = 0.274$, Bin Size = 0.0368, Number of Bins = 50					
No. per Bin	Bin Center	$\langle \text{F606} - \text{F814} \rangle$	σ	$\langle \text{F606W} \rangle$	σ
5.....	–0.82520300	1.62138000	0.11453257	–0.82572000	0.01088511
2.....	–0.78840900	1.65370000	0.11840000	–0.79715000	0.00715000
4.....	–0.75161500	1.69665000	0.24597608	–0.74417500	0.00733941
3.....	–0.71482100	1.41780000	0.02011765	–0.71806667	0.00549201
2.....	–0.67802700	1.36825000	0.03005000	–0.68070000	0.01010000
2.....	–0.64123300	1.66125000	0.34805000	–0.64785000	0.01055000
3.....	–0.60443900	1.55713333	0.35709536	–0.60706667	0.00978445

data lines from an SHB model with 1000 stars at $[\text{Fe}/\text{H}] = 0$, $[\alpha/\text{Fe}] = 0$, and $Y = 0.274$.

Output from `shb.f` may be processed by `shblf.f` to create a fiducial line and luminosity function from the SHB model. The `shblf` program creates luminosity functions and fiducials in any of the seven bands and with either a set number of bins or a set bin size. It scans through an SHB model file and calculates the number of stars, the average and standard deviation of the absolute magnitude, and the average and standard deviation of the color in each bin, and finally, it prints these data along with a short header. Table 11 shows the header and first seven data lines of an SHB luminosity function derived from the SHB model described above (Table 10).

E. COMPUTER PROGRAMS AND SUMMARY

Five computer programs, written in FORTRAN 77, are available along with the data files. These programs allow for the following manipulations of the data files:

`iso_interp_feh.f`: Isochrone interpolation in $[\text{Fe}/\text{H}]$ at fixed $[\alpha/\text{Fe}]$, Y , and age.

`iso_split.f`: Creation of individual age isochrone files from one of the large, many-aged files.

`isolf.f`: Luminosity function construction from an isochrone file.

`shb.f`: SHB model construction with $[\text{Fe}/\text{H}]$ interpolation.

`shblf.f`: Luminosity function/fiducial line construction from an SHB model file.

An accompanying README text file contains the text of this appendix.

F. HOW TO OBTAIN THE DATA FILES AND PROGRAMS

The entire distribution of stellar evolution tracks, isochrones, and accompanying computer programs may be downloaded from the Dartmouth Stellar Evolution Database⁷ and the Multimission Archive at the Space Telescope Science Institute⁸ as part of the ACS Survey of Galactic Globular Clusters (GO-10775).

⁷ See <http://stellar.dartmouth.edu/~models/>.

⁸ See <http://archive.stsci.edu>.

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