

7-10-2001

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## Recommended Citation

Rines, K.; Mahdavi, A.; Geller, M. J.; Diaferio, A.; Mohr, J. J.; and Wegner, G., "X-Ray–emitting Groups in the Infall Region of Abell 2199" (2001). *Open Dartmouth: Faculty Open Access Articles*. 2266.  
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## X-RAY-EMITTING GROUPS IN THE INFALL REGION OF ABELL 2199

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Received 2000 December 12; accepted 2001 March 19

### ABSTRACT

Using a large redshift survey covering 95 deg<sup>2</sup>, we demonstrate that the infall region of Abell 2199 contains Abell 2197, one or two X-ray-emitting groups, and up to five additional groups identified in redshift surveys. Our survey shows that the X-ray-emitting systems, located at projected radii of 1°.4, 1°.9, and 5°.1 (2.2, 3.1, and 8.0  $h^{-1}$  Mpc), are connected kinematically to A2199. A2197 is itself an optically rich cluster; its weak X-ray emission suggests that it is much less massive than A2199. The absence of a sharp peak in the infall pattern at the position of A2197 supports this hypothesis. The outermost group is well outside the virial region of A2199, and it distorts the infall pattern in redshift space. The two X-ray-emitting groups are roughly colinear, suggesting the existence of an extended (8.0  $h^{-1}$  Mpc) filament. The identification of these infalling groups provides direct support of hierarchical structure formation; studies of these systems will provide insights into structure evolution. Groups in the infall regions of nearby clusters may offer a unique probe of the physics of the warm/hot ionized medium (WHIM), which is difficult to observe directly with current instruments.

*Subject headings:* cosmology: observations — galaxies: clusters: general — galaxies: kinematics and dynamics

### 1. INTRODUCTION

Hierarchical structure formation predicts that rich clusters of galaxies should be surrounded by less massive, infalling clusters or groups. Many authors have discussed substructure within the virial region (see, e.g., Geller & Beers 1982; Dressler & Shectman 1988; White, Briel, & Henry 1993; Mohr, Fabricant, & Geller 1993). Reisenegger et al. (2000, hereafter RQCM) recently demonstrated dynamical connections between many clusters in the Shapley supercluster and A3558, the central cluster. We describe similar observations of the infall region around A2199 (Abell 1958) and show that it contains A2197, one or two X-ray-emitting groups, and up to five additional groups identified in redshift surveys, all of which lie outside the virial region of A2199.

In redshift space, the infall regions of clusters form a characteristic trumpet-shaped pattern (bounded by caustics) as galaxies fall into the cluster when the cluster potential overwhelms the Hubble flow (Kaiser 1987; Regös & Geller 1989; Diaferio & Geller 1997). Galaxies outside the caustics are outside the infall region; within the caustics, there are few interlopers. Here, we use the observed caustics around A2199 to identify members of the infall region and to compare the structure of the supercluster as seen in galaxies with the X-ray structure revealed by the *ROSAT* All-Sky Survey (RASS;<sup>6</sup> Voges et al. 1999). We use the X-ray data to show that the infall region contains bound systems destined for accretion. The physical scale at the

redshift of the supercluster ( $cz = 9156 \text{ km s}^{-1}$ ) is  $1^\circ = 1.53 h^{-1} \text{ Mpc}$  ( $H_0 = 100 h \text{ km s}^{-1}$ ,  $\Omega_m = 0.3$ ,  $\Omega_\Lambda = 0.7$ ).

### 2. OBSERVATIONS

We have collected 1216 redshifts (957 new or remeasured) in a large region (95 deg<sup>2</sup>) surrounding A2199/A2197. Two of us (J. J. M. and G. W.) used the Decaspec (Fabricant & Hertz 1990) at the 2.4 m MDM telescope on Kitt Peak to obtain 249 (248 unique) redshifts of galaxies in the central  $1^\circ$  of this region for a separate Jeans analysis of the central region of A2199 (Mohr et al. 2001, in preparation).

We used the FAST spectrograph (Fabricant et al. 1998) on the 1.5 m Tillinghast Telescope of the Fred Lawrence Whipple Observatory (FLWO) to obtain 684 spectra of galaxies within  $5.5^\circ$  ( $\approx 8.4 h^{-1} \text{ Mpc}$ ) of the center of A2199. The observations are similar to those described in Rines et al. (2000) and will appear in Rines et al. (2001, in preparation). We observed infall galaxy candidates in two campaigns. We selected targets from digitized images of the POSS I from the Automated Plate Scanner (APS).<sup>7</sup> We initially selected galaxies using the automatic classification system of APS; we visually inspected these targets to eliminate stars. The first campaign yielded a deep sample in the central  $4^\circ \times 4^\circ$  region around A2199. This sample is complete to 103a-E magnitude  $E < 16.5$  (roughly equivalent to  $R < 16.0$ ) and consists of 305 redshifts. The second campaign (379 redshifts) was a shallower survey ( $E < 16.1$  or  $R \lesssim 15.6$ ) of all galaxies within  $5.5^\circ \approx 8.4 h^{-1} \text{ Mpc}$  of A2199. The completeness limits are not very precise because the magnitudes come from multiple plate scans and because we could not obtain redshifts for some low surface brightness galaxies. We include 84 redshifts, associated with the groups NRGs385 and NRGs388, obtained by FAST for a separate study of the X-ray and optical properties of groups of galaxies (59 published in Mahdavi et al. 1999; see also Mahdavi et al. 2000, hereafter MBGR).

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<sup>6</sup> The RASS is available at <http://www.xray.mpe.mpg.de/cgi-bin/rosat/rosat-survey>.

<sup>7</sup> The APS databases are supported by the National Aeronautics and Space Administration and the University of Minnesota and are available at <http://aps.umn.edu>.

We collected the remaining 200 redshifts from the CfA Redshift Catalog<sup>8</sup> and/or the NASA/IPAC Extragalactic Database.<sup>9</sup> X-ray data for these systems are available from the RASS.

### 3. DEFINING THE INFALL REGION WITH CAUSTICS

Figure 1 displays the projected radii and redshifts of galaxies surrounding A2199. The expected caustic pattern is easily visible; we calculate the shape with the technique described in Diaferio (1999), using a smoothing parameter

<sup>8</sup> Available at <http://cfa-www.harvard.edu/~huchra/zcat>.

<sup>9</sup> The NASA/IPAC Extragalactic Database is available at <http://nedwww.ipac.caltech.edu/index.html>.

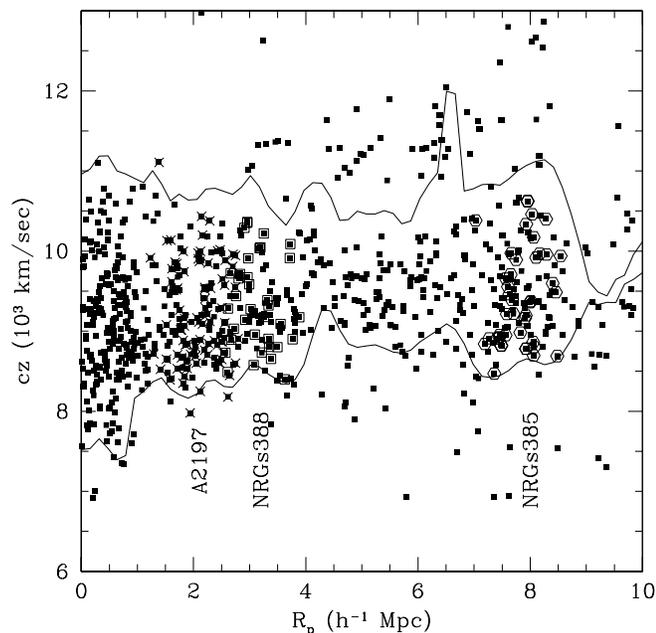


FIG. 1.—Projected radius vs. redshift for galaxies surrounding A2199. Lines indicate our estimate of the caustics. Crosses, hexagons, and open squares indicate galaxies in A2197, NRGs385, and NRGs388, respectively. There are 743 galaxies within the caustics.

of  $q = 25$ . We use a hierarchical clustering analysis to determine the center of the largest system, A2199; we use this as the center of the supercluster. The caustics shown in Figure 1 do not include the symmetry and first derivative constraints of Diaferio (1999). We will determine the caustics and the supercluster mass profile in more detail in Rines et al. (2001, in preparation). Preliminary results indicate that the total supercluster mass is in the range  $(5-10) \times 10^{14} h^{-1} M_{\odot}$  within a radius of  $\sim 8 h^{-1}$  Mpc. For comparison, Markevitch et al. (1999, hereafter MVFS) use X-ray data to estimate a mass of  $1.7 \times 10^{14} h^{-1} M_{\odot}$  within  $0.6 h^{-1}$  Mpc for A2199. The A2199 supercluster is located within the Great Wall (Geller & Huchra 1989; Fig. 6b of Falco et al. 1999), which may complicate the interpretation of the dynamics of the system.

The sky positions of galaxies in the infall region (crosses in Fig. 2) show several groups in addition to the main cluster. Table 1 lists the coordinates and basic properties of these systems (MBGR). Figure 2 displays X-ray contours from the RASS (*red*), with contours of the local galaxy density overlaid (*blue*). Purple areas in the figure therefore show regions with both significant galaxy overdensities and X-ray emission; X-ray emission confirms that a system is bound. The largest purple area is A2199; at least one background X-ray source and A2197 are to the north of A2199, and NRGs388 and NRGs385 are to the southwest. The X-ray emission from A2199 is quite symmetric relative to other clusters (Mohr et al. 1995; MVFS), which suggests that the inner region of A2199 has not undergone any recent major mergers. The infalling groups are all located at projected radii significantly larger than the virial radius ( $R_{\text{vir}} \approx 1.6 h^{-1}$  Mpc, Girardi et al. 1998;  $R_{\text{vir}} \approx 1.8 h^{-1}$  Mpc, MVFS). We identify galaxies within  $1 h^{-1}$  Mpc of these systems in Figure 1.

NRGs385, NRGs388, and A2199 are roughly colinear (Fig. 2). This alignment may be coincidental, or it may indicate the presence of a filament of galaxies and/or dark matter. A2197 and an optical group with no X-ray counterpart in the RASS lie roughly along the extension of this line to the northeast. The apparent X-ray excess in a northeast-southwest band, coincidentally along the broad apparent

TABLE 1

PROPERTIES OF SYSTEMS ASSOCIATED WITH A2199

SYSTEM	COORDINATES		$cz$ ( $\text{km s}^{-1}$ )	$\sigma_p$ ( $\text{km s}^{-1}$ )	$R_p$ ( $h^{-1}$ Mpc)	$\log L_X$ ( $h^{-2}$ ergs $\text{s}^{-1}$ )
	R.A. (J2000)	Decl. (J2000)				
A2199 .....	16 28 38	39 33 05	8963	$801 \pm 92^a$	...	43.9
A2197W .....	16 27 41	40 55 40	9300	$612 \pm 56^a$	2.2	42.9
A2197E .....	16 29 43	40 49 12	...	...	2.0	42.7
NRGs385 .....	16 17 15	34 55 00	9478	$525 \pm 216$	8.0	42.9
NRGs388 .....	16 23 01	37 55 21	9788	$468 \pm 94$	3.1	42.6
NRGs389 .....	16 21 57	36 02 22	10,096	$132 \pm 38$	5.8	<41.8
NRGs396 <sup>b</sup> .....	16 36 50	44 13 00	9540	$480 \pm 105$	7.6	42.0 <sup>c</sup>
NRGs395 .....	16 37 05	36 09 01	9957	$331 \pm 67$	5.8	<41.9
NRGs399 .....	16 41 42	39 47 07	9677	$407 \pm 94$	3.8	<41.9
NRGs400 .....	16 48 13	35 59 23	10,085	$724 \pm 424$	8.1	<41.9

NOTE.—Units of right ascension are hours, minutes, and seconds, and units of declination are degrees, arcminutes, and arcseconds.

<sup>a</sup> Girardi et al. 1998.

<sup>b</sup> Identified in Ramella, Pisani, & Geller 1997; too poor to be included in MBGR, but named using their convention.

<sup>c</sup> Detection significance =  $2.7 \sigma$ .

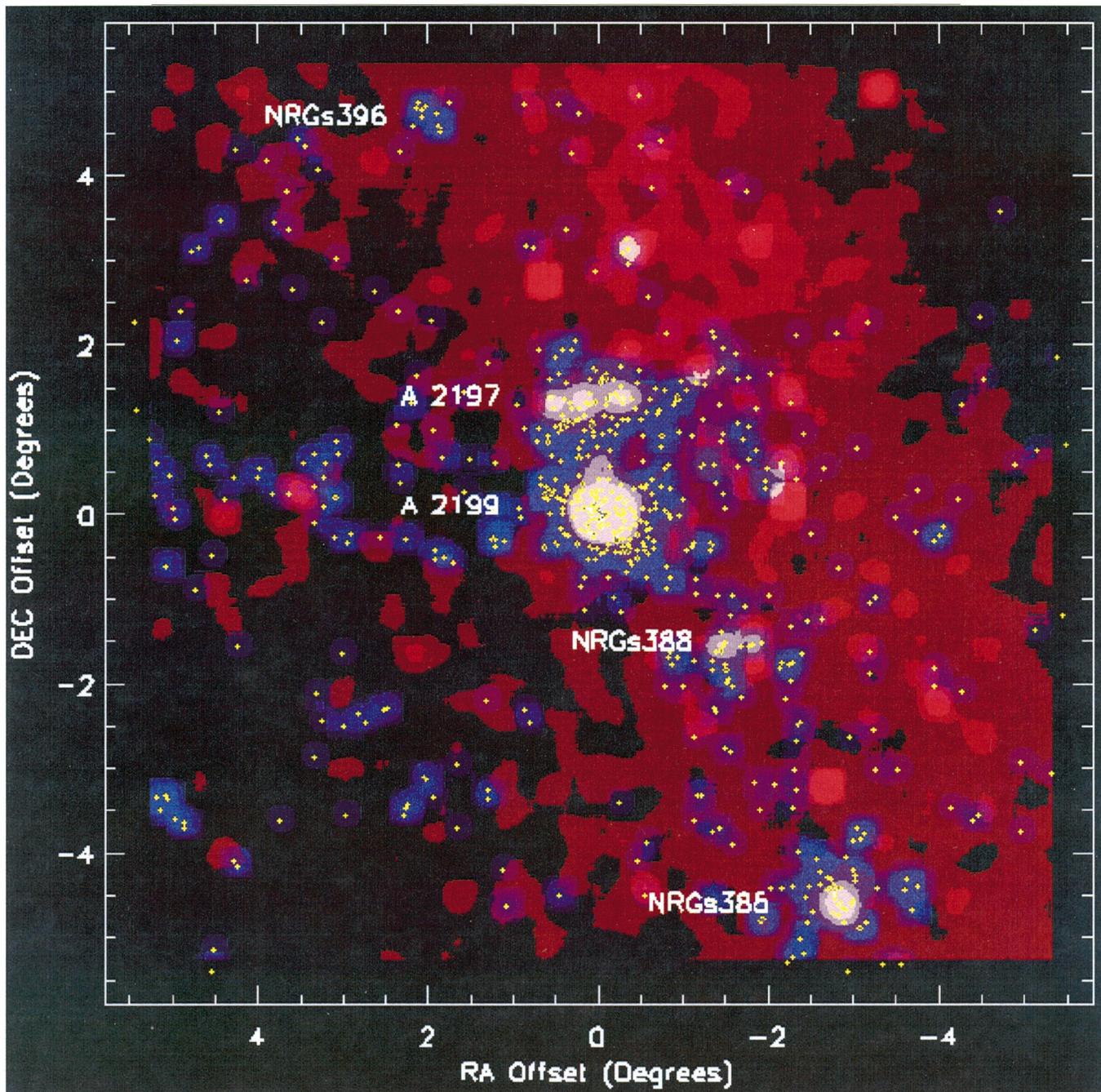


FIG. 2.—Sky positions of 743 galaxies in the infall region. Red shows X-ray intensity from the RASS; blue shows smoothed infall galaxy density. Regions with X-ray emission and large galaxy density thus appear purple. Several (most likely background) X-ray point sources are present and appear as red circles. Some groups of galaxies have no associated X-ray emission detectable in the RASS.

filament, is probably an artifact of the RASS scan pattern. Archival pointed PSPC observations indicate that the X-ray surface brightness of the filament is less than  $4 \times 10^{-16} \text{ ergs cm}^{-2} \text{ s}^{-1} \text{ arcmin}^{-2}$ . Briel & Henry (1995) obtain similar limits for other clusters.

#### 4. INFALLING SYSTEMS

##### 4.1. *Abell 2197*

Rood (1976) first suggested that, because of their close proximity on the sky and in redshift space, A2197 and A2199 are part of a supercluster. Gregory & Thompson (1984) analyzed the pair as a binary cluster. A2197 is an

optically rich Abell cluster with an apparent velocity dispersion comparable to that of A2199 (Girardi et al. 1998), but X-ray data suggest that A2197 is significantly less massive than A2199 (see, e.g., White, Jones & Forman 1997; Jones & Forman 1999). The absence of a sharp spike in the caustic pattern at the projected radius of A2197 (Fig. 1) indicates that the infall pattern is dominated by A2199 and that the apparently large velocity dispersion of A2197 is due in part to its proximity to A2199.

A2199 and A2197 provide an interesting contrast. A2199 is a rich, regular cluster with both a centrally concentrated galaxy distribution and X-ray emission. The galaxies around A2197 are significantly elongated to the southeast

and northwest, and they are much less centrally concentrated than in A2199. Archival *ROSAT* data reveal that the X-ray emission has at least three components (all three are evident in Fig. 2). The central component is centered on NGC 6173, the more luminous component to the west is centered on NGC 6160, and the source to the east is not centered on any bright galaxies and is probably a background source (C. Jones 2000, private communication; note also that a point source is evident between A2199 and A2197). The galaxy distribution in A2197 is quite different from the X-ray emission, though not as dramatically as in A754 (Zabludoff & Zaritsky 1995). We suggest that A2197 is either forming from the merger of two groups or being disrupted by A2199.

#### 4.2. NRGs388

The X-ray emission from NRGs388 is centered on a bright elliptical that dominates the group (see Fig. 6 of MBGR). There is no significant substructure in either the X-ray emission or the galaxy distribution within NRGs388. The apparent absence of ram pressure stripping or tidal distortions may constrain the distribution of intercluster gas or the shape of the gravitational potential of the supercluster.

#### 4.3. NRGs385

NRGs385 is located  $5^{\circ}.1 = 8.0 h^{-1}$  Mpc or  $\sim 5$  virial radii from A2199. Although other investigators have found X-ray-emitting groups around rich clusters (see, e.g., Briel, Henry, & Böhringer 1992; Wang, Connolly, & Brunner 1997; Kull & Böhringer 1999; Donnelly et al. 1999; RQCM), to our knowledge no other X-ray group so far outside the virial radius of a cluster has been linked to it kinematically. While the kinematic connection is clear, NRGs385 may not be bound to the supercluster (Rines et al. 2001, in preparation).

NRGs385 is sufficiently massive that it distorts the caustic pattern in its vicinity (Fig. 1). This effect is expected; subclustering increases the amplitude of the caustics at all radii (Diaferio 1999), but spikes in the amplitude of the caustics can reveal the presence of massive subclusters. The existence of this distortion confirms the hypothesis that subclusters can alter the sharp caustics expected from spherical infall. It is striking that the expected distortion is evident from this group but no such distortions are evident in the infall region of A3558 (see Fig. 1 of RQCM), even though it contains infalling clusters that are probably more massive than NRGs385. The reason for this difference is unclear, although it might depend on the galaxy population included by the selection criteria of RQCM.

The peak of the X-ray emission in NRGs385 is offset (at roughly the  $3\sigma$  level) from the center of the galaxy distribution. The X-ray center is located  $\approx 218 h^{-1}$  kpc southwest of the optical center, away from A2199. If this offset is not due to a projection effect, it could indicate either that the galaxies do not trace the gravitational potential of the group or that the group has recently undergone a merger. We rule out the possibility of ram pressure stripping (see, e.g., Gunn & Gott 1972) by the warm/hot ionized medium (WHIM) because the required gas density in the WHIM would produce X-ray emission above our upper limit on the surface brightness of any filament (§ 3).

The shape of the X-ray contours may provide information about the physical processes occurring in the infalling

groups. NRGs385 is visibly elongated in the direction of A2199 and along the possible filament; the width of the  $10\sigma$  contour is approximately  $26' = 720 h^{-1}$  kpc northeast-southwest and approximately  $19' = 525 h^{-1}$  kpc southeast-northwest (see Fig. 6 of Mahdavi et al. 1999). This elongation may be caused by tidal effects and/or the presence of a filament.

#### 4.4. X-Ray-Faint Groups

Several small knots of galaxies are evident in Figure 2. Many of these are probably chance superpositions rather than physical groups. Four of them are contained in the RASSCALs optical catalog (MBGR), though none has significant extended X-ray emission (we list upper limits from the RASS in Table 1). A fifth possible group, NRGs396 (Ramella, Pisani, & Geller 1997; named according to the convention of MBGR), is located  $5^{\circ}.0 = 7.6 h^{-1}$  Mpc northeast of A2199, but is too poor (four group members in the CfA redshift survey) for inclusion in the RASSCALs catalog (MBGR require a minimum of five group members in the CfA redshift survey). The X-ray-faint group NRGs389 lies between NRGs385 and NRGs388 and is aligned with the possible filament.

We find marginal evidence ( $2.7\sigma$ ) of X-ray emission near NRGs396 in the RASS; pointed observations of these groups with *XMM-Newton* or *Chandra* would place tighter constraints on their X-ray properties. Like NRGs385, NRGs396 and NRGs400 are sufficiently distant from the center of the supercluster that they may not be gravitationally bound.

### 5. DISCUSSION

The infall pattern in redshift space (delineated by the caustics) indicates that A2197 and one or two X-ray groups (NRGs388 and NRGs385) are connected kinematically to A2199. The three systems are located at projected radii of  $1^{\circ}.4$ ,  $1^{\circ}.9$ , and  $5^{\circ}.1$  ( $2.2$ ,  $3.1$ , and  $8.0 h^{-1}$  Mpc), respectively; we also identify five X-ray-faint groups in the infall region. All three systems with X-ray emission are significantly less massive than A2199.

To our knowledge, NRGs388 (and possibly NRGs385) is the first X-ray group demonstrably bound to a cluster even though it lies well outside the virial radius (RQCM demonstrate similar connections between X-ray clusters in the Shapley supercluster). The identification of caustics delineating the infall regions of clusters reveals the kinematic connections between clusters and distant, infalling subclusters.

At large radii from rich clusters, it is difficult to observe the WHIM directly with current X-ray instruments (Kull & Böhringer 1999; Pierre, Bryan, & Gastaud 2000; Cen & Ostriker 1999). The presence of hot gas in at least some infalling groups shows that the intracluster medium in clusters of galaxies can be accreted from infalling groups. Like groups at smaller radii (White et al. 1993, Donnelly et al. 1999), these groups may interact with the WHIM and thus yield insight into its properties and processes, such as cooling and feedback (Pierre et al. 2000).

This project would not have been possible without the assistance of Perry Berlind and Michael Calkins, the remote observers at FLWO, and Susan Tokarz, who processed the spectroscopic data. We thank Lars Hernquist and Christine Jones for helpful discussions. K. R., A. M., and M. G. are

supported in part by the Smithsonian Institution. J. J. M. is supported by Chandra Fellowship grant PF8-1003, awarded through the Chandra Science Center. The Chandra Science Center is operated by the Smithsonian Astrophysical Observatory for NASA under contract NAS 8-39073. We thank the Max-Planck-Institut für Astro-

physik, where some of the computing for this work was done. This research has made use of the NASA/IPAC Extragalactic Database (NED), which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

## REFERENCES

- Abell, G. O. 1958, *ApJS*, 3, 211  
 Briel, U. G., & Henry, J. P. 1995, *A&A*, 302, L9  
 Briel, U. G., Henry, J. P., & Böhringer, H. 1992, *A&A*, 259, L31  
 Cen, R., & Ostriker, J. P. 1999, *ApJ*, 514, 1  
 Diaferio, A. 1999, *MNRAS*, 309, 610  
 Diaferio, A., & Geller, M. 1997, *ApJ*, 481, 633  
 Donnelly, R. H., Markevitch, M., Forman, W., Jones, C., Churazov, E., & Gilfanov, M. 1999, *ApJ*, 513, 690  
 Dressler, A., & Shectman, S. A. 1988, *AJ*, 95, 985  
 Fabricant, D., Cheimets, P., Caldwell, N., & Geary, J. 1998, *PASP*, 110, 79  
 Fabricant, D., & Hertz, E. 1990, *Proc. SPIE*, 1235, 747  
 Falco, E. E., et al. 1999, *PASP*, 111, 438  
 Geller, M. J., & Beers, T. C. 1982, *PASP*, 94, 421  
 Geller, M. J., & Huchra, J. P. 1989, *Science*, 246, 897  
 Girardi, M., Giuricin, G., Mardirossian, F., Mezzetti, M., & Boschin, W. 1998, *ApJ*, 505, 74  
 Gregory, S. A., & Thompson, L. A. 1984, *ApJ*, 286, 422  
 Gunn, J. E., & Gott, J. R., III 1972, *ApJ*, 176, 1  
 Jones, C. & Forman, W. 1999, *ApJ*, 511, 65  
 Kaiser, N. 1987, *MNRAS*, 227, 1  
 Kull, A., & Böhringer, H. 1999, *A&A*, 341, 23  
 Mahdavi, A., Böhringer, H., Geller, M. J., & Ramella, M. 2000, *ApJ*, 534, 114 (MBGR)  
 Mahdavi, A., Geller, M. J., Böhringer, H., Kurtz, M. J., & Ramella, M. 1999, *ApJ*, 518, 69  
 Markevitch, M., Vikhlinin, A., Forman, W. R., & Sarazin, C. L. 1999, *ApJ*, 527, 545 (MVFS)  
 Mohr, J. J., Evrard, A. E., Fabricant, D. G., & Geller, M. J. 1995, *ApJ*, 447, 8  
 Mohr, J. J., Fabricant, D. G., & Geller, M. J. 1993, *ApJ*, 413, 492  
 Pierre, M., Bryan, G., & Gastaud, R. 2000, *A&A*, 356, 403  
 Ramella, M., Pisani, A., & Geller, M. J. 1997, *AJ*, 113, 483  
 Regös, E., & Geller, M. 1989, *AJ*, 98, 755  
 Reisenegger, A., Quintana, H., Carrasco, E. R., & Maze, J. 2000, *AJ*, 120, 523 (RQCM)  
 Rines, K., Geller, M. J., Diaferio, A., Mohr, J. J., & Wegner, G. 2000, *AJ*, 120, 2338  
 Rood, H. J. 1976, *ApJ*, 207, 16  
 Voges, W., et al. 1999, *A&A*, 349, 389  
 Wang, Q. D., Connolly, A. J., & Brunner, R. J. 1997, *ApJ*, 487, L13  
 White, D. A., Jones, C., & Forman, W. 1997, *MNRAS*, 292, 419  
 White, S. D. M., Briel, U. G., & Henry, J. P. 1993, *MNRAS*, 261, L8  
 Zabludoff, A. I., & Zaritsky, D. 1995, *ApJ*, 447, L21