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THE ASYMMETRIC EXPLOSION OF TYPE Ia SUPERNOVAE AS SEEN FROM NEAR-INFRARED OBSERVATIONS¹

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ABSTRACT

We present near-infrared spectra of late-phase (>200 days) Type Ia supernovae (SNe Ia) taken at the Subaru Telescope. The [Fe II] line of SN 2003hv shows a clear flat-topped feature, while that of SN 2005W shows a less prominent flatness. In addition, a large shift in their line center, varying from -3000 to 1000 km s⁻¹ with respect to the host galaxies, is seen. Such a shift suggests the occurrence of an off-center, nonspherical explosion in the central region and provides important, new constraints on the explosion models of SNe Ia.

Subject headings: infrared: stars — supernovae: general —
supernovae: individual (SN 2003du, SN 2003hv, SN 2005W)

1. INTRODUCTION AND SUMMARY

The brightness and approximate uniformity of Type Ia supernovae (SNe Ia) enable us to use them as reliable and distant standard candles, and they provide evidence of an accelerating universe (Riess et al. 1998; Perlmutter et al. 1999). For precision cosmology, it is critical to understand the origin of diversity as well as the explosion mechanism of SNe Ia, which is still open to debate (see, e.g., Hillebrandt & Niemeyer 2000 for a review). Possible sources of diversity include the asymmetry of the explosion (Wang et al. 2003), due to the rotation of the progenitor white dwarfs (Piersanti et al. 2003; Saio & Nomoto 2004; Uenishi et al. 2003; Yoon & Langer 2003), and/or the turbulent behavior of the deflagration flame (e.g., Gamezo et al. 2003; Röpke et al. 2006). Therefore, it is important to observationally investigate the distribution of the synthesized elements and the kinematical structure of the ejecta in order to constrain the explosion models.

In this respect, late-phase (~1 yr since the explosion) spectroscopy at the near-infrared (NIR) wavelength provides important diagnostics. Because the ejecta become optically thin in late phases, spectroscopy provides an unbiased, direct view of the innermost regions. Well-isolated [Fe II] emission lines at 1.257 and 1.644 μm enable us to trace the distribution of the most important isotopes synthesized in SNe Ia, i.e., ⁵⁶Ni (which decays into ⁵⁶Co and then ⁵⁶Fe) and other iron isotopes such as ⁵⁴Fe. NIR observations of the Type Ia SN 1991T show a strongly peaked [Fe II] 1.644 μm line that appears symmetric

with respect to the rest frame of the host galaxy (Spyromilio et al. 1992; Bowers et al. 1997). The NIR [Fe II] line profile of SN 1998bu observed by Spyromilio et al. (2004) appears to have a somewhat less centrally peaked emission profile.

This approach was highlighted by the discovery of a flat-topped [Fe II] 1.644 μm emission line in SN 2003du (Höflich et al. 2004). Such a line profile requires a central hole in the kinematic distribution of the radioactive ejecta. This sort of hollow radioactivity distribution is predicted in one-dimensional (1D) explosion models (see, e.g., Nomoto et al. 1994 for a review), in which the innermost regions are burned under high densities, $\rho > 10^9$ g cm⁻³, and electron capture produces stable isotopes of ⁵⁸Ni, ⁵⁴Fe, and ⁵⁶Fe rather than radioactive ⁵⁶Ni. However, such a distribution is contrary to the predictions of state-of-the-art 3D deflagration simulations that predict large-scale turbulent mixing in the inner layers; thus, these simulations do not confine ⁵⁶Ni outside the central core. The Höflich et al. (2004) NIR spectrum of SN 2003du also indicates that the [Fe II] 1.257 and 1.644 μm emission lines show a blueshift of 500–1000 km s⁻¹ with respect to the host galaxy.

However, since the observed sample size of late-phase NIR spectra is so small, it is not clear whether or not the flat-topped and blueshifted lines are general properties of SNe Ia. Such observations are extremely difficult because even relatively nearby and bright SNe Ia are faint in NIR at late phases (e.g., $H = 20$ – 21 for SN 2003du ~1 yr after the explosion). We therefore have conducted NIR observations of SNe Ia using the Subaru Telescope with an OH airglow suppressor (Table 1) and obtained two more late-phase 1–2 μm spectra of SN 2003hv and SN 2005W.

A detailed explanation and discussion of the spectra will be presented elsewhere (K. Motohara et al. 2006, in preparation). We briefly report three interesting results in this Letter. (1) The [Fe II] lines of SN 2003hv also show a clear flat-topped feature. (2) The [Fe II] lines of SN 2005W show less prominent flatness. (3) A large velocity shift of the line center, varying from -3000 to 1000 km s⁻¹ with respect to the host galaxies, exists.

Considering the uniformity in optical bands around the peak luminosity, the existence of such inhomogeneity and asymmetry in SNe Ia is surprising and provides important, new constraints on the explosion models.

¹ Based on data collected at Subaru Telescope, which is operated by the National Astronomical Observatory of Japan.

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TABLE 1
INFRARED SPECTROSCOPY OBSERVING LOG

Name	Host	z_{host}	Obs. Date (UT)	Epoch ^a (days)	$\lambda\lambda$ (Å)	$\lambda/\Delta\lambda$	Exposure (s)	Slit (arcsec)
2003du	UGC 9391	0.006384	2004 Feb 27.5	+297.2	11080–18040	400	4000	0.5
2003hv	NGC 1201	0.005604	2004 Oct 06.5	+394	11080–18040	400	12000	0.5
2005W	NGC 691	0.008889	2005 Sep 12.6	+214	11080–18040	400	4000	0.5

^a Epochs are with respect to the epoch of *B*-band maximum light ($t_{B_{\text{max}}} = 0$).

2. OBSERVATIONS

2.1. SN 2003du

SN 2003du was discovered by LOTOSS (Lick Observatory and the Tenagra Observatory Supernova Searches) on 2003 April 22.4 in UGC 9391 at about 15.9 mag (Schwartz & Holvorcem 2003). It was confirmed to be a SN Ia by Kotak et al. (2003), who reported that the optical spectrum resembles that of SN 2002bo about 2 weeks before maximum, and reached maximum light in the *B* band ($t_{B_{\text{max}}} = 0$ days) on 2003 May 6.3 UT (JD 2,452,766.3 \pm 0.5; Anupama et al. 2005).

As reported in Höflich et al. (2004), *JH*-band spectroscopy of SN 2003du was carried out using OHS/CISCO (OH-airglow Suppressor/Cooled Infrared Spectrograph and Camera for OHS; Iwamuro et al. 2001; Motohara et al. 2002) at the Subaru Observatory. For this work, we re-reduced the data taken on 2004 February 27.5 (+297 days), which were processed in a standard procedure of flat-fielding, sky subtraction, bad-pixel correction, and residual sky subtraction. The flux was scaled using the *H*-band photometry that was taken on 2003 April 2.5, assuming that the change in the brightness is negligible at the late phase. The wavelength was calibrated using the standard pixel-wavelength relation of CISCO, of which the systematic error is estimated to be less than 0.5 pixels (<3 Å).

2.2. SN 2003hv

SN 2003hv was discovered by LOTOSS on 2003 September 9.5 (UT) in NGC 1201 at about 12.5 mag (Beutler & Li 2003) and confirmed to be a SN Ia by the spectrum taken on 2003 September 10.4, which resembles that of SN 1994D 2 days after maximum (Dressler et al. 2003). We therefore assume maximum light in the *B* band to be JD 2,452,891.

Our *JH*-band spectroscopy of SN 2003hv was carried out using OHS/CISCO on 2004 October 6 (epoch +394 days). The *JH*-band spectroscopy consists of six frames of 2000 s exposure with an 0".5 slit, providing a wavelength resolution of ~ 400 . The A2 star SAO 169939 was observed after the target to correct the atmospheric and instrumental absorption pattern. The data were reduced in the same manner as that of SN 2003du.

2.3. SN 2005W

SN 2005W was discovered on 2005 February 1.4 (UT) in NGC 691 at about 15.2 mag (Nakano & Li 2005) and confirmed to be a SN Ia about a week before maximum on 2005 February 2.7 (UT) (Elias-Rosa et al. 2005). The expansion velocity is measured to be $\sim 11,600$ km s⁻¹ from the Si II line. We assume maximum light in the *B* band to be JD 2,453,413 (2005 February 10).

Our *JH*-band spectroscopy of SN 2005W was carried out on 2005 September 12 (epoch +214 days) by OHS/CISCO. It consists of four frames of 1000 s exposure with an 0".5 slit. The A2 star HIP 20091 was observed after the target to correct

the atmospheric and instrumental absorption pattern. The data were reduced in the same manner as that of SN 2003du.

3. RESULTS

The NIR spectra are shown in Figure 1. It can be seen that all the observed SNe Ia exhibit strong [Fe II] 1.257 and 1.644 μm lines. We discovered two important features.

First, the line center of [Fe II], corrected for the redshift of the host galaxies, is not identical for the three events. With respect to the rest wavelength of the line (1.644 μm), two SNe Ia (2003du, 2003hv) show a blueshift, while the other (2005W) shows a redshift. The corresponding velocity shifts relative to the hosts' rest frame are -2600 , -2100 , and $+1400$ km s⁻¹ for SNe 2003hv, 2003du, and 2005W, respectively.

Such a velocity shift is confirmed by a mid-infrared (MIR) spectrum of SN 2003hv taken by the Infrared Spectrograph (IRS) on the *Spitzer Space Telescope* $\sim +360$ days after maximum brightness (Fig. 2; L. Gerardy et al. 2006, in preparation). One of the strongest emission features is identified as the ground-state fine-structure line of [Co III] at 11.89 μm ($a^4F_{9/2} - a^4F_{7/2}$). The well-isolated [Co III] line also shows a velocity shift that is consistent with that seen in the NIR [Fe II] 1.644 μm line, to within the noise level of the MIR feature.

Second, we find that the observed [Fe II] lines show a large variety in their shape (Fig. 1). The spectrum of SN 2003hv clearly shows a flat-topped boxy profile like that seen at a much lower signal-to-noise ratio (S/N) in SN 2003du. Indeed, the higher S/N of the SN 2003hv spectrum places a much stronger constraint on the flatness of the core of the [Fe II] line. In contrast, the [Fe II] line profile of SN 2005W shows no evidence of a central flat top.

The boxy profile of [Fe II] lines seen in SN 2003du and SN 2003hv suggests the absence of radioactive ⁵⁶Ni below ~ 3000 km s⁻¹ (measured from the center of the ⁵⁶Ni distribution). We note that the observed epoch of SN 2003hv is more advanced than SN 2003du. At a later epoch, the line profile should follow the ⁵⁶Ni distribution more closely (§ 4). Effects of possible line deformation due to electron and line scatterings will also effectively vanish. The present result of SN 2003hv confirms the existence of the ⁵⁶Ni empty hole more strongly than that of SN 2003du.

4. DISCUSSION

In this section we discuss the impact of our findings for understanding the explosion mechanism. It is widely agreed that the explosion of a SN Ia starts from a deflagration (Nomoto et al. 1976). In the spherical deflagration models, electron capture leads to the synthesis of ⁵⁸Ni, ⁵⁴Fe, and ⁵⁶Fe (not via ⁵⁶Ni decay), thus creating an almost ⁵⁶Ni empty hole (e.g., Nomoto et al. 1984). In multidimensional models, the ignition of the deflagration may be off-center, producing a nonspherical burning region (Wunsch & Woosley 2004; Plewa et al. 2004). In

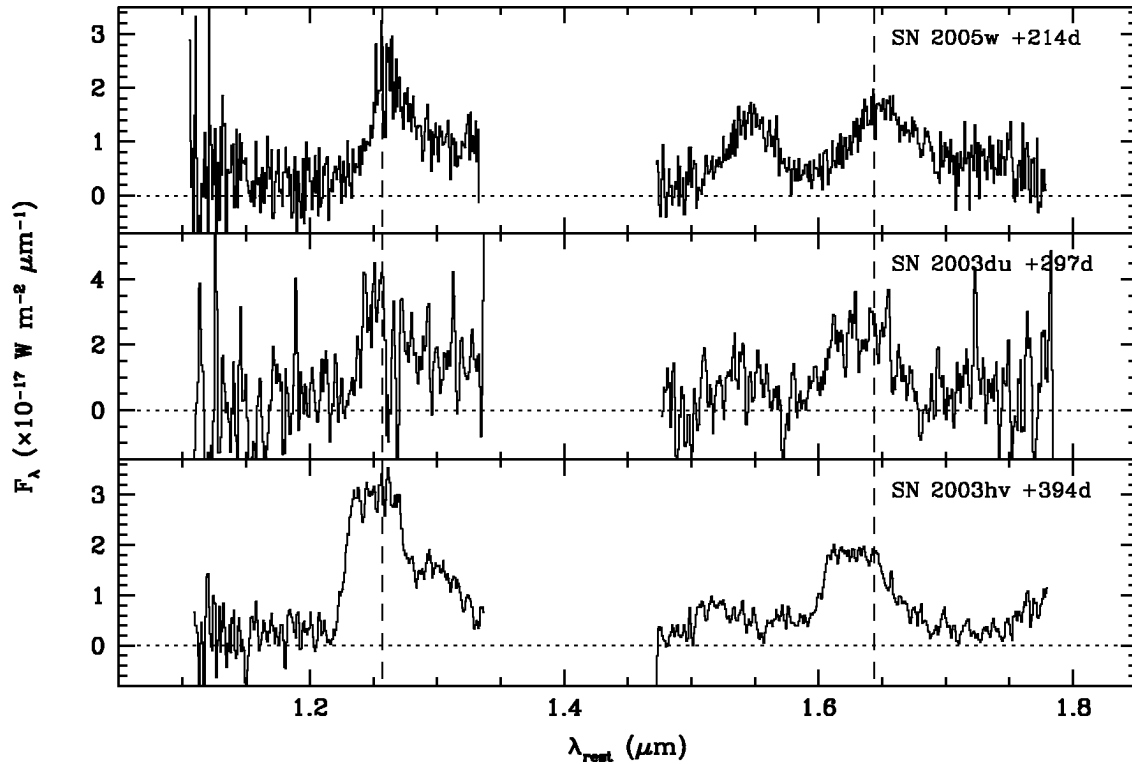


FIG. 1.—NIR 1–2 μm spectra of late-phase SNe Ia, converted to the rest wavelength at the host galaxies. All the spectra are smoothed with a 3 pixel boxcar filter. The vertical dashed lines show the position of [Fe II] 1.257 and 1.644 μm lines.

some cases, the deflagration to detonation transition (DDT) may occur (delayed detonation models; Khokhlov 1991). The DDT may also take place nonspherically (Livne 1999), even if the deflagration does not have a bulk kinematical offset.

The flat-topped [Fe II] NIR lines in SNe 2003hv and 2003du indicate that the highest density region occupied by

neutron-rich Fe-peak isotopes is not mixed with the surrounding region where the dominant isotope is ^{56}Ni (Höflich et al. 2004). Figure 2 shows a spectrum for the spherical deflagration explosion model W7 (Nomoto et al. 1984) 400 days after the explosion, compared with that of SN 2003hv. This is calculated by solving the transport of γ -rays produced in the decay chain $^{56}\text{Ni} \rightarrow ^{56}\text{Co} \rightarrow ^{56}\text{Fe}$ and by iteratively solving non-LTE rate equations (Mazzali et al. 2001; Maeda et al. 2006). Positrons produced by the ^{56}Co -decay are assumed to be trapped on the spot, since the mean free path of positrons is expected to be small at this epoch (Milne et al. 2001).

The model shows that the flat-topped profile is consistent with the [Fe II] 1.644 μm emission for the reason described above. The asymmetry in the profile, mildly peaking in the red, is due to the weak contributions of [Fe II] 1.664 and 1.677 μm .

The shift of [Fe II] lines suggests that the distribution of ^{56}Ni (which decays to ^{56}Fe via ^{56}Co) produced at the explosion is asymmetric, showing a bulk kinematical offset of $\geq 2000 \text{ km s}^{-1}$ with respect to the SN rest frame. This suggests that, unlike the spherical models, the carbon ignition may take place off-center. The flat-topped [Fe II] line profiles suggest that electron capture in the high-density off-center deflagration region creates the neutron-rich hole as in 1D models, at least in SNe 2003du and 2003hv. The neutron-rich hole is offset along with the bulk ^{56}Ni distribution. This indicates that the highest density burning in these SNe took place quite far away from the center of the progenitor star.

The blueshifted and flat-topped profile of [Fe II] can also be reproduced by a symmetric and opaque dusty core, like [O I] $\lambda\lambda 6300, 6364$ observed in a late-phase Type II SN (Elmhamdi et al. 2003). If this is the case in the present results, the MIR [Co III] line, which will be far less affected by dust extinction,

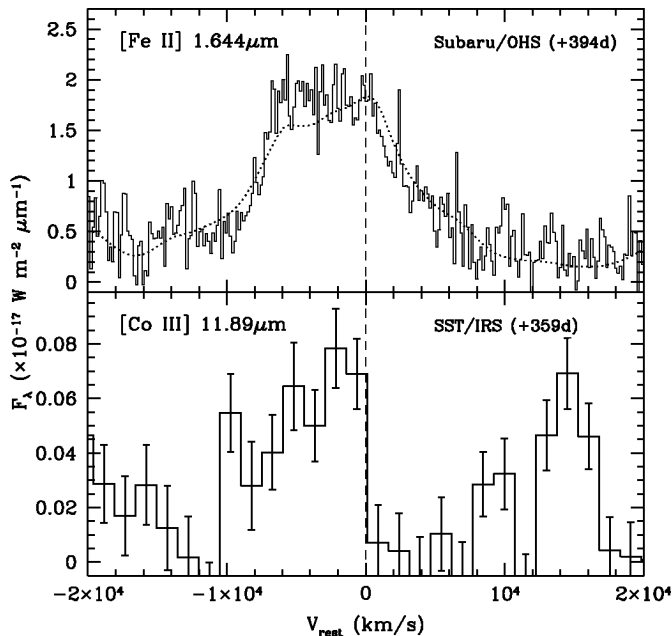


FIG. 2.—[Fe II] 1.644 μm (top) and [Co III] 11.89 μm (bottom) emission lines of SN 2003hv (solid histograms). Overlapping the [Fe II] line is the 1D explosion model W7 400 days after the explosion (dotted line). The model flux is arbitrarily scaled to fit the observed flux as well as shifted to the blue by 2600 km s^{-1} .

may show a much smaller blueshift than NIR [Fe II]. However, the [Co III] line shows almost the same blueshift as [Fe II] (Fig. 2). Therefore, it is unlikely that the flat-topped profile is caused by the opaque dust core.

Adding the result of SN 1991T (Spyromilio et al. 1992; Bowers et al. 1997), two SNe (SN 2003du, SN 2003hv) out of four SNe Ia show the boxy [Fe II] profile (Fig. 1), while SN 1991T and SN 2005W do not. This could indicate that the distribution of ^{56}Ni in the innermost region of the SN Ia explosion might differ from object to object. SNe Ia with a non-boxy line profile might experience more mixing than those with a boxy profile (e.g., SN 2003hv) in the innermost region. Another possibility is that the density of the progenitor, where the ignition takes place, differs from object to object, and a SN Ia with the peaked [Fe II] explodes at a low density so as not to produce electron-capture isotopes.

Alternatively, the different line shape could be due to an age effect, i.e., due to a variation in the extent to which the energy deposition from the radioactivity is kept local. At earlier epochs, γ -rays are the dominant heating source. Since the γ -ray penetration is not a local process, the innermost region, even without ^{56}Ni , can be heated effectively. At later epochs, the contribution from the local positron energy input becomes larger, and the line profile should follow the ^{56}Ni distribution more closely. Therefore, we expect that the line shape could evolve from a peaked to a flat-topped profile. This may partially

explain the fact that the most aged SN Ia (2003hv) shows the clean flat-topped boxy [Fe II] line. Indeed, this provides an observational test of our model.

We suggest a future test to estimate the age effect and to confirm our interpretation of the ^{56}Ni : a temporal series of NIR nebular spectra for an individual SN Ia. The shape of an [Fe II] line is expected to start getting flat at ~ 250 days and may become totally flat at ~ 500 days. Thereafter, it may start showing a peaked profile at ~ 1000 days again, depending on the amount of positrons that can escape the ejecta (Milne et al. 2001). Also, spectra at different wavelengths from optical to the MIR will be useful for investigating the distribution of different ions and heating radioactive isotopes (e.g., Fig. 2). Because the line profiles from different ions and from different energy levels are dependent on the ionization and thermal structure of the ejecta, further detailed theoretical study is necessary to make use of these observations efficiently.

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